Neutrino electromagnetic interactions: A window to new physics

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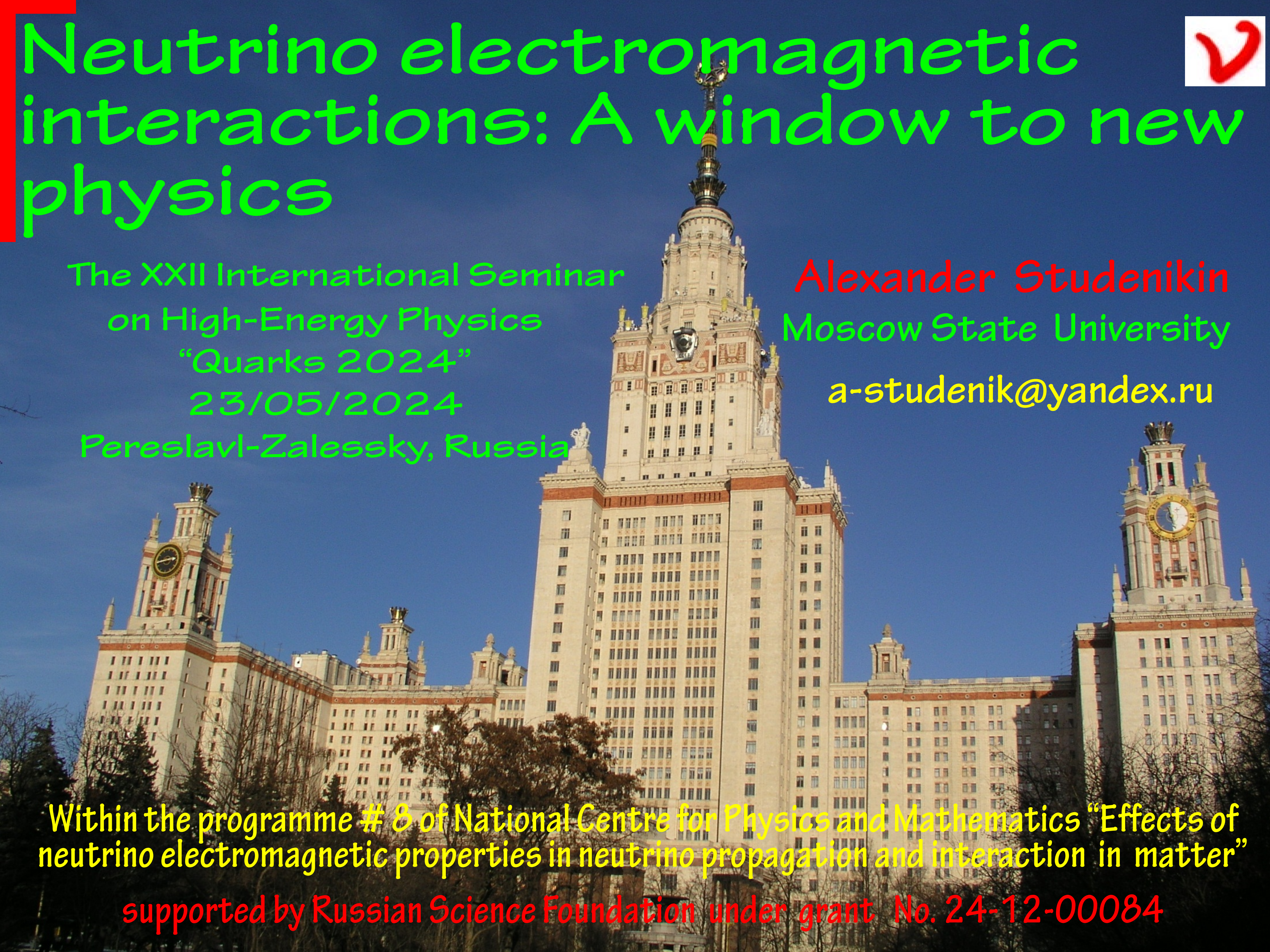
Alexander Studenikin
Moscow State University
a-studenik@yandex.ru



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Outline

1. reminder of ν electromagnetic properties

2. constraints on μ_ν, d_ν, q_ν and $<r^2_\nu>$ from laboratory experiments

3. effects of electromagnetic ν interactions in astrophysics

4. astrophysical probes of electromagnetic ν

5. new effects in ν oscillations related to electromagnetic ν interactions

. . . new phenomena in ν spin (flavor) oscillations in moving and polarized mater and magnetic field of interest for astrophysical applications . . .

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Neutrino electromagnetic interactions: A window to new physics

+ upgrade:

Studenikin, Overview of neutrino electromagnetic properties 2022, arXiv:2301.06071

Studenikin, Overview of neutrino electromagnetic properties (the theory, laboratory experiments and astrophysical probes), PoS (NuFact2021) 402(2022)052

Detailed review on emp of ν

… > 500 citations

Carlo Giunti[*]
INFN, Torino Section, Via P. Giuria 1, I-10125 Torino, Italy

Alexander Studenikin[†]
Department of Theoretical Physics, Faculty of Physics, Moscow State University and Joint Institute for Nuclear Research, Dubna, Russia

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A review is given of the theory and phenomenology of neutrino electromagnetic interactions, which provide powerful tools to probe the physics beyond the standard model. After a derivation of the general structure of the electromagnetic interactions of Dirac and Majorana neutrinos in the one-photon approximation, the effects of neutrino electromagnetic interactions in terrestrial experiments and in astrophysical environments are discussed. The experimental bounds on neutrino electromagnetic properties are presented and the predictions of theories beyond the standard model are confronted.

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CONTENTS

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Neutrino magnetic moment: a window to new physics

A. Studenikin[a] *
Department of Theoretical Physics, Moscow State University, 119991 Moscow, Russia

A short review on a neutrino magnetic moment is presented.

Introduction. Experimental and theoretical studies of flavour conversion in solar, atmospheric, reactor and accelerator neutrino fluxes give strong evidence of non-zero neutrino mass. A massive neutrino can have non-trivial electromagnetic properties [1]. For a recent review on neutrino electromagnetic properties see [2].

The neutrino dipole magnetic moment (along with the electric dipole moment) is the most well studied among neutrino electromagnetic properties. The effective Lagrangian, that is in charge of a neutrino coupling to the electromagnetic field, can be written in the form

$$L_{int} = \frac{1}{2}\bar{\psi}_i\sigma_{\alpha\beta}(\mu_{ij} + \epsilon_{ij}\gamma_5)\psi_j F^{\alpha\beta} + h.c. \quad (1)$$

where the magnetic moments μ_{ij}, in the presence of mixing between different neutrino states, are associated with the neutrino mass eigenstates ν_i. The interplay between magnetic moment and neutrino mixing effects is important. Note that electric (transition) moments ϵ_{ij} do also contribute to the coupling.

A Dirac neutrino may have non-zero diagonal electric moments in models where CP invariance is violated. For a Majorana neutrino the diagonal magnetic and electric moments are zero. Therefore, neutrino magnetic moments can be used to distinguish Dirac and Majorana neutrinos (see [3] and also [2] for a detailed discussion).

Neutrino magnetic moment in a minimal extension of Standard Model. The explicit evaluation of the one-loop contributions to the Dirac neutrino magnetic moment in the leading approximation over small parameters $b_i = \frac{m_i^2}{M_W^2}$ (m_i are the neutrino masses, $i = 1, 2, 3$), that however exactly accounts for $a_l = \frac{m_l^2}{M_W^2}$ ($l = e, \mu, \tau$), leads the following result [4],

$$\mu^D_{ij} = \frac{eG_F m_i}{8\sqrt{2}\pi^2}\left(1+\frac{m_j}{m_i}\right)\sum_{l=e,\mu,\tau} f(a_l)U_{lj}U^*_{li}, \quad (2)$$

$$f(a_l) = \frac{3}{4}\left[1+\frac{1}{1-a_l}-\frac{2a_l}{(1-a_l)^2}-\frac{2a_l^2\ln a_l}{(1-a_l)^3}\right],$$

where U_{li} is the neutrino mixing matrix. The correspondent result in the absence of mixing was confirmed in [5,6]. A Majorana neutrino may also have transition moment of the value $\mu^M_{ij} = 2\mu^D_{ij}$ (see [2] for a detailed discussion and references).

For the diagonal magnetic moment of the Dirac neutrino, from (2) in the limit $a_l \ll 1$ the result [1] can be obtained

$$\mu^D_{ii} = \frac{3eG_F m_i}{8\sqrt{2}\pi^2}\left(1-\frac{1}{2}\sum_{l=e,\mu,\tau} a_l \mid U_{li}\mid^2\right). \quad (3)$$

The magnetic moment for hypothetical heavy neutrino was studied in [6]. In particular, it was obtained

$$\mu_\nu = \frac{eG_F m_\nu}{8\sqrt{2}\pi^2}\begin{cases} 3+\frac{5}{6}b, & m_\ell \ll m_\nu \ll M_W, \\ 1, & m_\ell \ll M_W \ll m_\nu. \end{cases} \quad (4)$$

Note that the LEP data set a limit on number of light neutrinos coupled to Z boson.

The numerical value of the Dirac neutrino magnetic moment within a minimal extension of the Standard Model, as it follows from (3), is

$$\mu^D_{ii} \approx 3.2\times 10^{-19}\left(\frac{m_i}{1\ eV}\right)\mu_B, \quad (5)$$

This is several orders of magnitude smaller than the present experimental limits if to account for the existed constraints on neutrino masses.

*e-mail: studenik@srd.sinp.msu.ru

... Why ν electromagnetic properties are important ?

...Why ν em properties

to new physics ?

... How does it all relate to ν oscillations ?

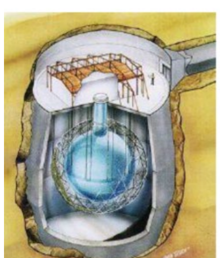
$\Longrightarrow$ $m_\nu \neq 0$ •

magnetic moment $\mu_\nu \neq 0$ •

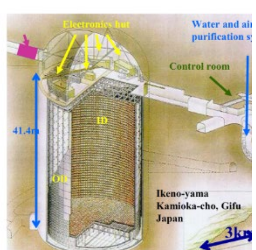
Arthur McDonald

The Nobel Prize in Physics 2015

Takaaki Kajita



«for the discovery of neutrino oscillations, which shows that neutrinos have mass»



in Standard Model

• $m_\nu = 0$!!!

70 years ago ...

C. L. Cowan, F. Reines and F. B. Harrison,
Upper limit on the neutrino magnetic moment,
Phys. Rev. 96 (1954) 1294

ν electromagnetic properties

and possibility of measuring μ_ν

raised before experimental discovery of ν

... problem and puzzle ...

ν electromagnetic properties

up to now nothing has been seen

... in spite of reasonable efforts ...

- results of terrestrial lab experiments on μ_ν (and ν EM properties in general)

- as well as data from astrophysics and cosmology

are in agreement with "ZERO" ν EM properties

... However, in course of recent development of knowledge on ν mixing and oscillations,

In the easiest generalization of SM

$$\mu_{ii}^{D} = \frac{3eG_F m_i}{8\sqrt{2}\pi^2} \approx 3.2 \times 10^{-19}\left(\frac{m_i}{1\ \mathrm{eV}}\right)\mu_B$$

if $m_\nu \leq 0.8\ \ eV$ ⟸ KATRIN limit

then $\mu_{ii}^{D} \sim 3.2 \times 10^{-19}\mu_B$!

K.Fujikawa, R.Shrock, Phys.Rev.Lett. 45 (1980) 963

many orders of magnitude smaller than present experimental limits:

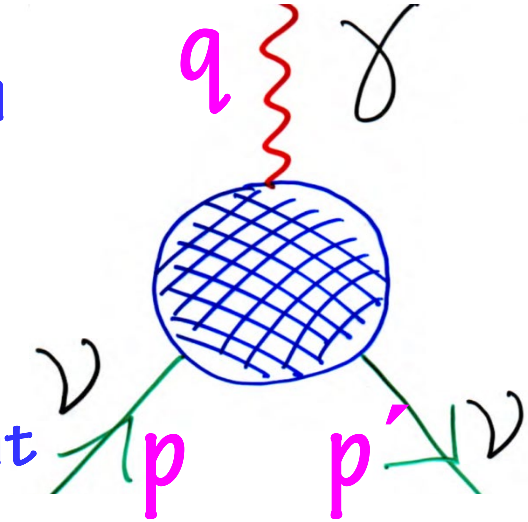
- $\mu_\nu \sim 10^{-11}\mu_B$ reactor ν limits GEMMA 2012
- $\mu_\nu \sim\ 10^{-11} \div\ 10^{-12}\mu_B$ astrophysical (ν_{solar} ν_{SN}, DM) limits
  Borexino 2017 - XENONnT 2023, LUX-ZEPLIN 2023

μ_ν is no less extravagant than possibility of $q_\nu \neq 0$

- limitations imposed by general principles of any theory are very strict
- $q_\nu \leq 3 \times 10^{-21} e$ from neutrality of hydrogen atom
- slightly weaker constraints are imposed by astrophysics $q_\nu \leq 1.3 \times 10^{-19}\ \ e_0$
  Studenikin, Tokarev, NPB, 2014

… a bit of ν electromagnetic properties theory …

✓ electromagnetic vertex function



$$< \psi(p')|J_\mu^{EM}|\psi(p) >= \bar{u}(p')\Lambda_\mu(q,l)u(p)$$

Matrix element of electromagnetic current
is a Lorentz vector

$\Lambda_\mu(q,l)$ should be constructed using

- matrices $\hat{1}, \quad \gamma_5, \quad \gamma_\mu, \quad \gamma_5\gamma_\mu, \quad \sigma_{\mu\nu},$
- tensors $g_{\mu\nu}, \quad \epsilon_{\mu\nu\sigma\gamma}$
- vectors q_μ and l_μ

$$q_\mu = p'_\mu - p_\mu, \; l_\mu = p'_\mu + p_\mu$$

- Lorentz covariance (1) and electromagnetic gauge invariance (2) ⟹

Matrix element of electromagnetic current between neutrino states

$$\langle \nu(p')|J_\mu^{EM}|\nu(p)\rangle = \bar{u}(p')\Lambda_\mu(q)u(p)$$

where vertex function generally contains 4 form factors

$$\Lambda_\mu(q) = f_Q(q^2)\gamma_\mu + f_M(q^2)i\sigma_{\mu\nu}q^\nu - f_E(q^2)\sigma_{\mu\nu}q^\nu\gamma_5 + f_A(q^2)(q^2\gamma_\mu - q_\mu \not{q})\gamma_5$$

1. electric
2. magnetic dipole
3. electric dipole
4. anapole

- Hermiticity and discrete symmetries of EM current J_μ^{EM} put constraints on form factors

Dirac ν

1) CP invariance + Hermiticity ⟹ $f_E = 0$,

2) at zero momentum transfer only electric Charge $f_Q(0)$ and magnetic moment $f_M(0)$ contribute to $H_{int} \sim J_\mu^{EM} A^\mu$!

3) Hermiticity itself ⟹ three form factors are real: $Im f_Q = Im f_M = Im f_A = 0$

Majorana ν

1) from CPT invariance (regardless CP or ~~CP~~).

$$f_Q = f_M = f_E = 0$$

…as early as 1939, W.Pauli…

- EM properties ⟹ a way to distinguish Dirac and Majorana ν

In general case matrix element of J_μ^{EM} can be considered between different initial $\psi_i(p)$ and final $\psi_j(p')$ states of different masses

$$< \psi_j(p')|J_\mu^{EM}|\psi_i(p) >= \bar{u}_j(p')\Lambda_\mu(q)u_i(p)$$

$p^2 = m_i^2, \;\; p'^2 = m_j^2$:

... beyond SM...

and

$$\Lambda_\mu(q) = \Big(f_Q(q^2)_{ij} + f_A(q^2)_{ij}\gamma_5\Big)(q^2\gamma_\mu - q_\mu \not{q}) + f_M(q^2)_{ij} i\sigma_{\mu\nu}q^\nu + f_E(q^2)_{ij}\sigma_{\mu\nu}q^\nu\gamma_5$$

form factors are matrices in ν mass eigenstates space.

Dirac ν (off-diagonal case $i \neq j$) Majorana ν

1) Hermiticity itself does not apply restrictions on form factors,

2) CP invariance + Hermiticity

$f_Q(q^2), \; f_M(q^2), \; f_E(q^2), \; f_A(q^2)$

are relatively real (no relative phases) .

... quite different EM properties ...

1) CP invariance + hermiticity

$\mu_{ij}^M = 2\mu_{ij}^D \;\; and \;\; \epsilon_{ij}^M = 0$ or

$\mu_{ij}^M = 0 \;\; and \;\; \epsilon_{ij}^M = 2\epsilon_{ij}^D$

… importance of μ_ν studies…

If diagonal $\mu_\nu \neq 0$ were confirmed

then ν^{Dirac}

… for ν^{Majorana}
non-diagonal = transitional
$\mu_\nu \neq 0$

… progress in experimental studies of μ_ν $\Rightarrow$

. . . a bit more on ν electromagnetic properties theory

(em properties in gauge models)

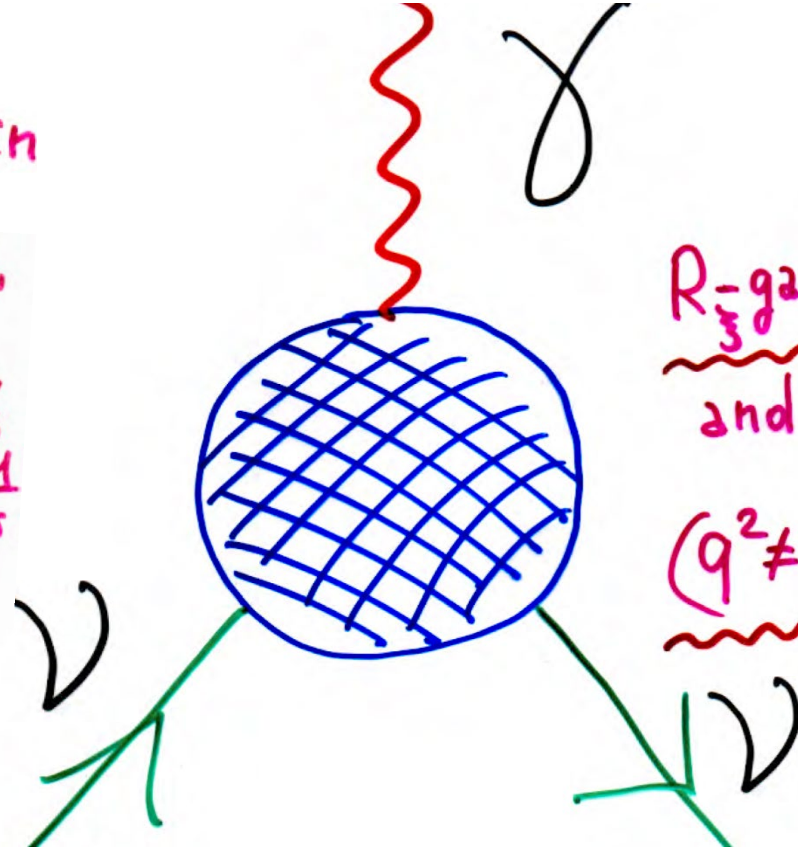
ν em vertex function

The most general study of the **massive neutrino** vertex function (including electric and magnetic form factors) in arbitrary R_ξ gauge in the context of the SM + SU(2)-singlet ν_R accounting for masses of particles in polarization loops

M. Dvornikov, A. Studenikin

⊛ Phys. Rev. D 63, 073001, 2004,
"Electric charge and magnetic moment of massive neutrino";
JETP 126 (2004), N8, 1

⊛ "Electromagnetic form factors of a massiv neutrino."



R_ξ-gauge

and

$(q^2 \neq 0)$

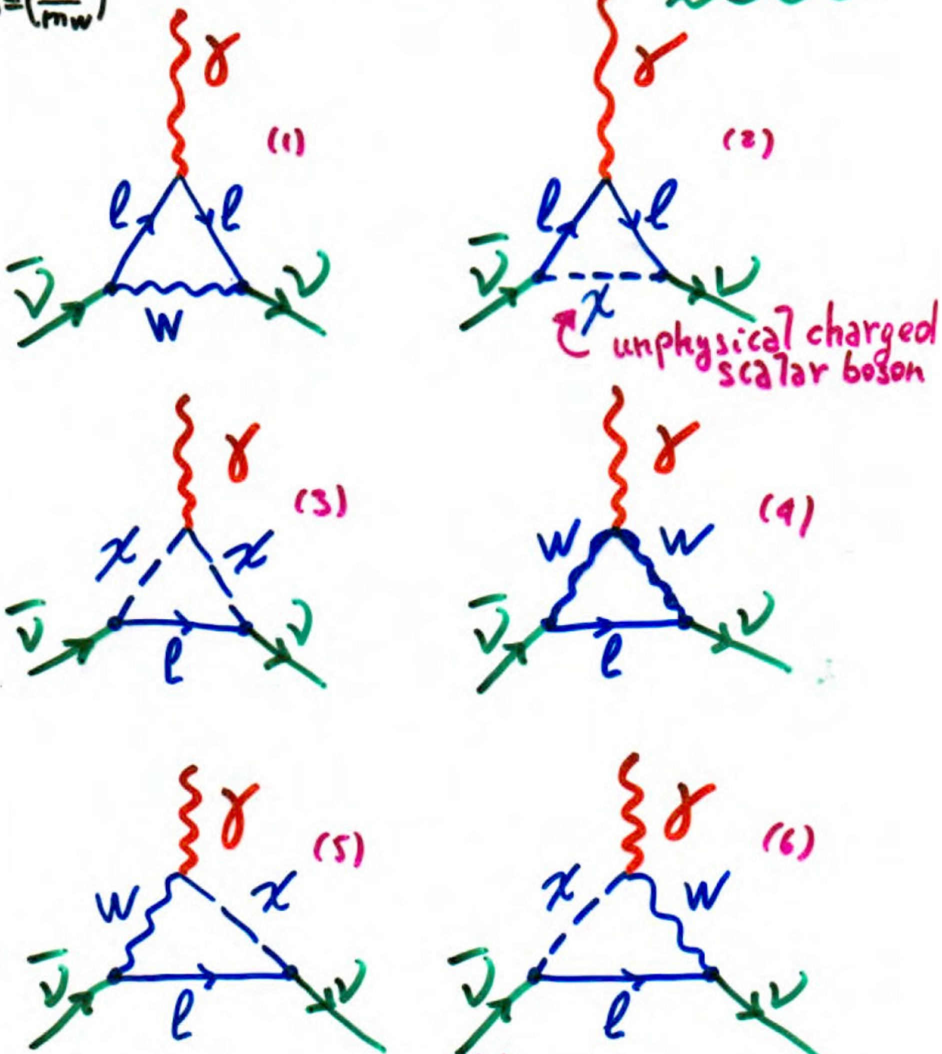
charge   magnetic moment

$$\Lambda_\mu(q) = f_Q(q^2)\gamma_\mu + f_M(q^2)\, i\sigma_{\mu\nu}q^\nu -$$
$$- f_E(q^2)\, i\sigma_{\mu\nu}q^\nu\gamma_5 - f_A(q^2)(q^2\gamma_\mu - q_\mu \not{q})\gamma_5$$

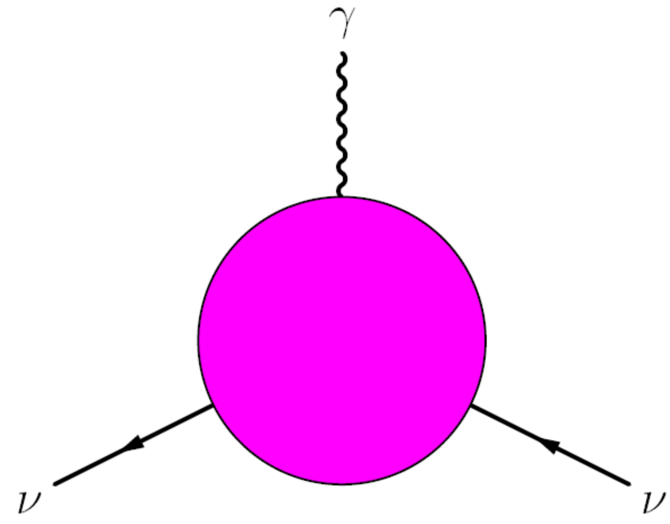
electric moment   anapole moment

$a=\left(\frac{m_\ell}{m_W}\right)^2,\quad b=\left(\frac{m_\nu}{m_W}\right)^2$

Proper vertices R_ξ-gauge



$$\Lambda_\mu(q)=\sum_{i=1}^{14}\Lambda^i_\mu(q)$$

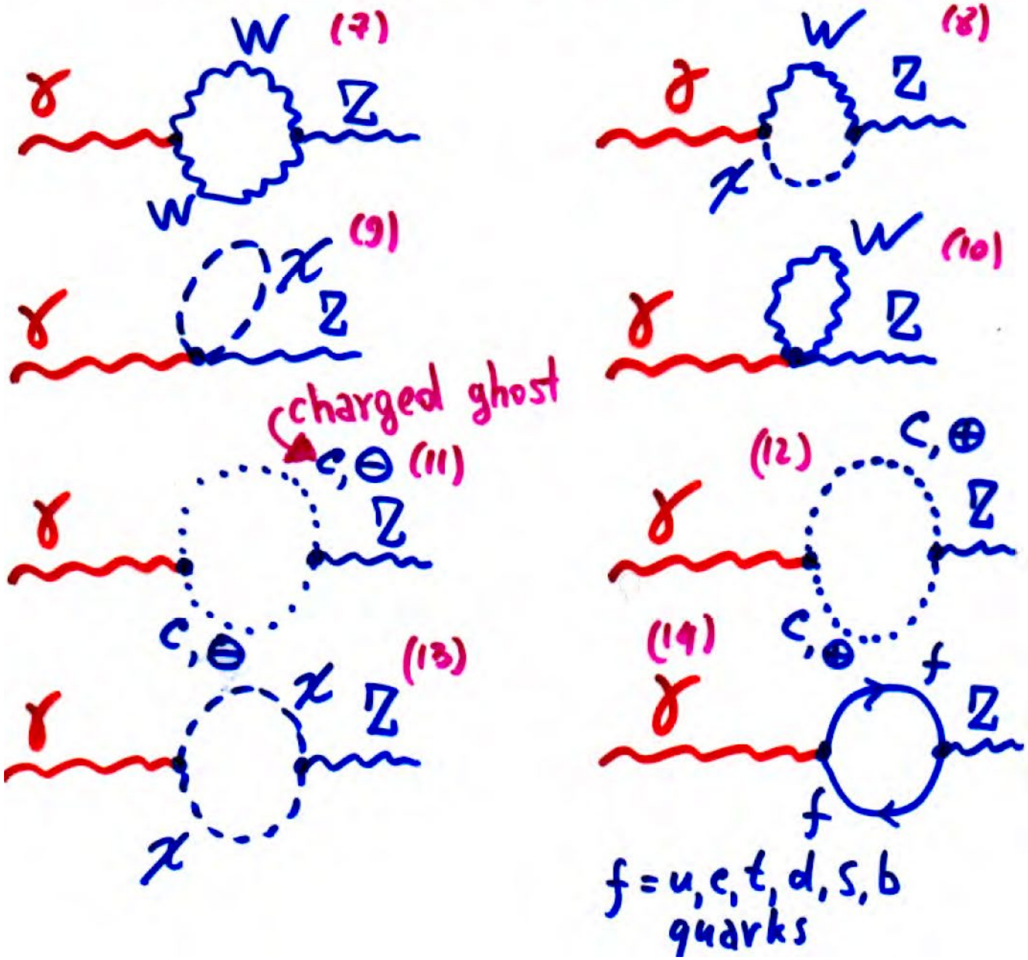


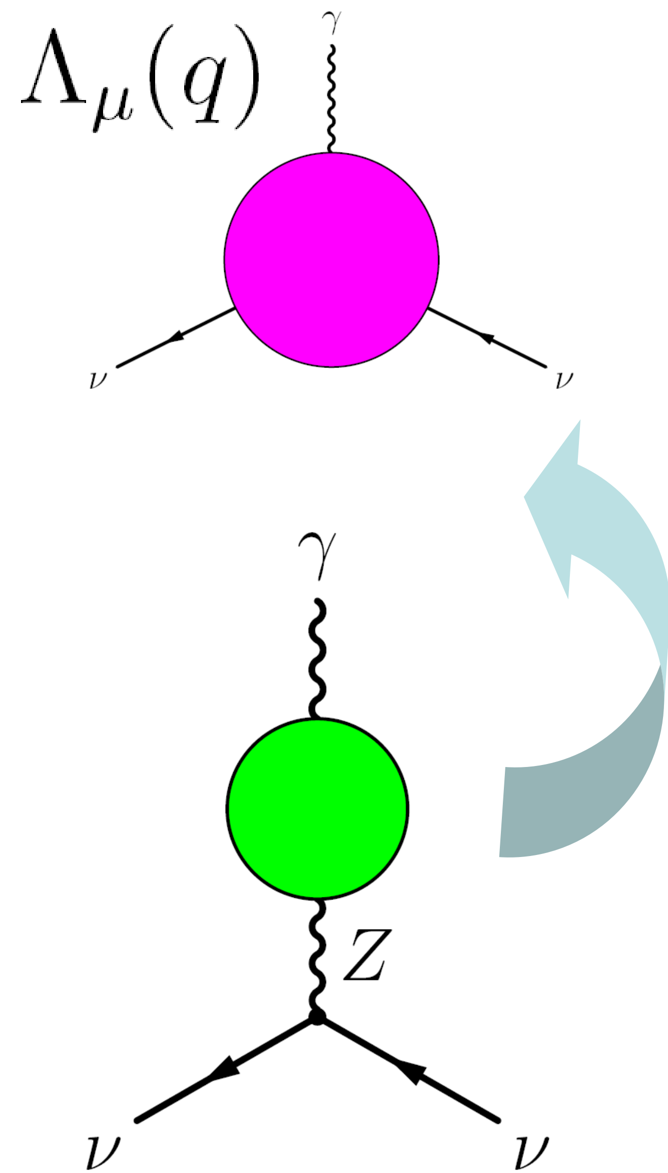
$\Lambda_\mu(q)$

$$\Lambda^{j}_{\mu}(q) = \frac{g}{2\cos\theta_W}\,\Pi^{(j)}_{\mu\nu}(q)\,\frac{1}{q^2 - M_Z^2}$$

$$\times \left\{ g^{\nu\alpha} - (1-\alpha_Z)\frac{q^\nu q^\alpha}{q^2 - \alpha_Z M_Z^2} \right\} \gamma^4_\alpha, \quad j = 7, \ldots, 14$$

γ-Z self-energy diagrams





$\gamma - Z$ self-energy diagrams

Magnetic moment dependence

$$\mu_\nu = \mu_\nu(m_\nu)$$

on neutrino mass

3.2

Calculation of ν magnetic moment (massive ν , arbitrary R_ξ - gauge)

Dvornikov, Studenikin PRD 2004

$$\Lambda_\mu(q) = f_Q(q^2)\gamma_\mu + f_M(q^2) i\sigma_{\mu\nu}q^\nu - f_E(q^2)\sigma_{\mu\nu}q^\nu\gamma_5 + f_A(q^2)(q^2\gamma_\mu - q_\mu \not{q})\gamma_5$$

magnetic moment

$$\mu(a,b,\alpha) = f_M(q^2=0)$$

two mass parameters

$$a = \left(\frac{m_\ell}{M_W}\right)^2 \qquad b = \left(\frac{m_\nu}{M_W}\right)^2$$

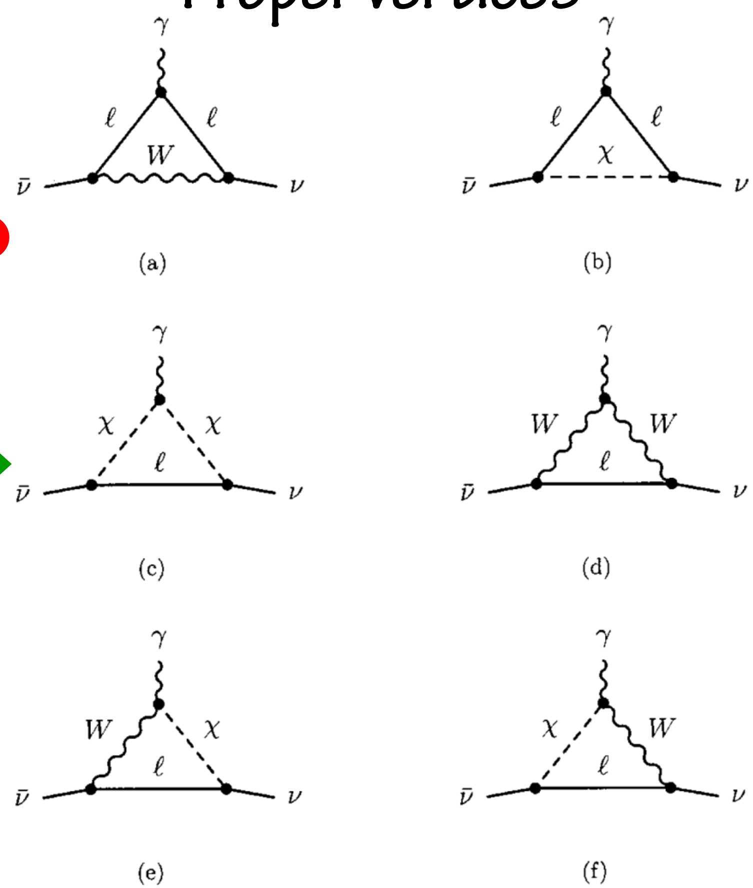
$$\mu(a,b,\alpha) = \sum_{i=1}^{6} \mu^{(i)}(a,b,\alpha)$$

and gauge-fixing parameter

$$\alpha = \frac{1}{\xi}$$

$\xi = 0$ - unitary gauge, $\xi = 1$ - 't Hooft-Feynman gauge

Proper vertices

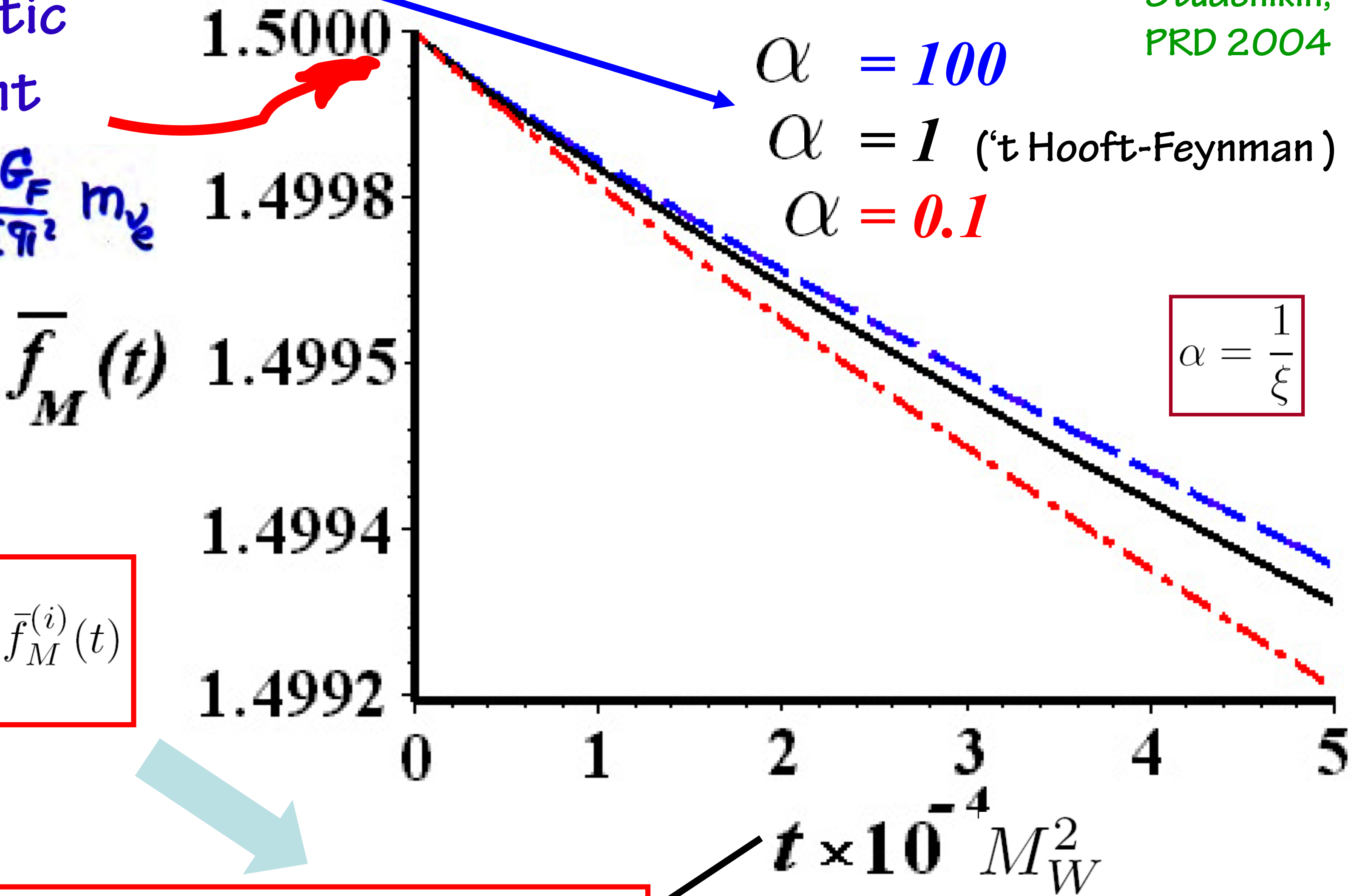


Gauge and q×q dependence ...

Dvornikov, Studenikin, PRD 2004

ν magnetic moment

● $\mu_e = \frac{3eG_F}{8\sqrt{2}\pi^2} m_{\nu_e}$



$$\bar{f}_M(t) = \sum_{i=1}^{6} \bar{f}_M^{(i)}(t)$$

$$f_M(q^2) = \frac{eG_F}{4\pi^2\sqrt{2}} m_\nu \sum_{i=1}^{6} \bar{f}_M^{(i)}(q^2)$$

ν dipole magnetic form factor

$m_\nu \ll m_e \ll M_W$ light ν

$\mu_e = \frac{3eG_F}{8\sqrt{2}\pi^2} m_{\nu_e}$

$$\mu_\nu = \frac{eG_F}{4\pi^2\sqrt{2}} m_\nu \frac{3}{4(1-a)^3}(2 - 7a + 6a^2 - 2a^2 \ln a - a^3) \qquad a = \left(\frac{m_e}{M_W}\right)^2$$

Dvornikov, Studenikin, Phys.Rev.D 69 (2004) 073001; JETP 99 (2004) 254

Gabral-Rosetti, Bernabeu, Vidal, Zepeda, Eur.Phys.J C 12 (2000) 633

$m_e \ll m_\nu \ll M_W$ intermediate ν

$$\mu_\nu = \frac{3eG_F}{8\pi^2\sqrt{2}} m_\nu \left\{1 + \frac{5}{18}b\right\} \qquad b = \left(\frac{m_\nu}{M_W}\right)^2$$

$m_e \ll M_W \ll m_\nu$

$$\mu_\nu = \frac{eG_F}{8\pi^2\sqrt{2}} m_\nu$$

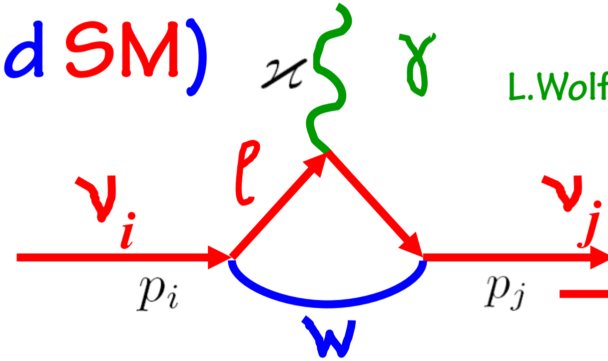
heavy ν

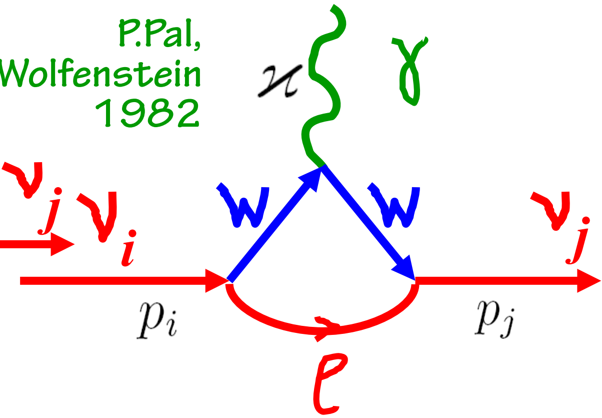
$\sim 10^{-19} \mu_B \left(\frac{m_{\nu_e}}{1\,eV}\right)$

... μ_ν in case of mixing ... $\Longrightarrow$

Neutrino (beyond SM) dipole moments (+ transition moments)

P.Pal, L.Wolfenstein 1982





- Dirac neutrino

$$\left.\begin{array}{c}\mu_{ij}\\ \epsilon_{ij}\end{array}\right\} = \frac{eG_F m_i}{8\sqrt{2}\pi^2}\left(1 \pm \frac{m_j}{m_i}\right) \sum_{l=e,\,\mu,\,\tau} f(r_l) U_{lj} U_{li}^*$$

$$r_l = \left(\frac{m_l}{m_W}\right)^2$$

$m_e = 0.5\ MeV$
$m_\mu = 105.7\ MeV$
$m_\tau = 1.78\ GeV$
$m_W = 80.2\ GeV$

- $m_i, m_j \ll m_l, m_W$

$$f(r_l) \approx \frac{3}{2}\left(1 - \frac{1}{2} r_l\right), \quad r_l \ll 1$$

transition moments vanish because unitarity of U implies that its rows or columns represent orthogonal vectors

- transition moments are suppressed, Glashow – Iliopoulos - Maiani cancellation,
- for diagonal moments there is no GIM cancellation

- Majorana neutrino only for $i \neq j$

$$\mu_{ij}^M = 2\mu_{ij}^D \;\; and \;\; \epsilon_{ij}^M = 0$$

or

$$\mu_{ij}^M = 0 \;\; and \;\; \epsilon_{ij}^M = 2\epsilon_{ij}^D$$

… depending on relative CP phase of ν_i and ν_j

The first nonzero contribution from neutrino transition moments

$$f_{r_l} \to -\frac{3}{2} + \frac{3}{4}\left(\frac{m_l}{m_W}\right)^2 \ll 1$$

GIM cancellation

$$\left.\begin{matrix}\mu_{ij}\\ \epsilon_{ij}\end{matrix}\right\} = \frac{3eG_F m_i}{32\sqrt{2}\pi^2}\left(1 \pm \frac{m_j}{m_i}\right)\left(\frac{m_\tau}{m_W}\right)^2 \sum_{l=e,\,\mu,\,\tau}\left(\frac{m_l}{m_\tau}\right)^2 U_{lj}U_{li}^*$$

$$\mu_B = \frac{e}{2m_e}$$

- $$\left.\begin{matrix}\mu_{ij}\\ \epsilon_{ij}\end{matrix}\right\} = 4\times 10^{-23}\mu_B\left(\frac{m_i \pm m_j}{1\ eV}\right)\sum_{l=e,\,\mu,\,\tau}\left(\frac{m_l}{m_\tau}\right)^2 U_{lj}U_{li}^*$$

→ … neutrino radiative decay is very slow

- Dirac ν diagonal (i=j) magnetic moment

$\epsilon_{ii}^D = 0$ for CP-invariant interactions

$\mu_{ii}^M = \epsilon_{ii}^M = 0$

$$\mu_{ii} = \frac{3eG_F m_i}{8\sqrt{2}\pi^2}\left(1 - \frac{1}{2}\sum_{l=e,\mu,\tau} r_l \mid U_{li}\mid^2\right) \approx 3.2\times 10^{-19}\left(\frac{m_i}{1\ eV}\right)\mu_B$$

$$r_l = \left(\frac{m_l}{m_W}\right)^2$$

Lee, Shrock, Fujikawa, 1977

- no GIM cancellation
- μ_{ii}^D - to leading order - independent on U_{li} and $m_{l=e,\,\mu,\,\tau}$
- $$\mu_e^2 = \sum_{i=1,2,3} \mid U_{ie}\mid^2 \mu_{ii}^2$$ …possibility to measure fundamental μ_{ii}^D

$\mu_{ii}^D = 0$ for massless ν (in the absence of right-handed charged currents) →

...the present status...

to have visible $m_\nu \neq 0$

is not an easy task for

theoreticians

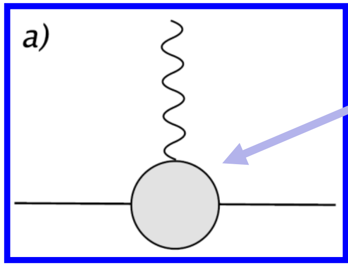
and experimentalists

3.3 Naïve relationship between m_ν and μ_ν

… problem to get large μ_ν and still acceptable m_ν

If μ_ν is generated by physics beyond the SM at energy scale Λ,

P. Vogel e.a., 2006

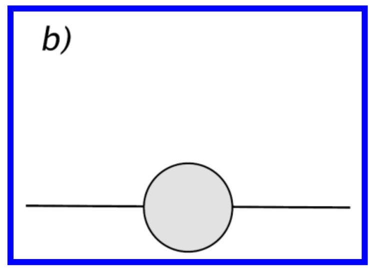


then

$$\mu_\nu \sim \frac{eG}{\Lambda},$$

…combination of constants and loop factors…

contribution to m_ν given by



, then

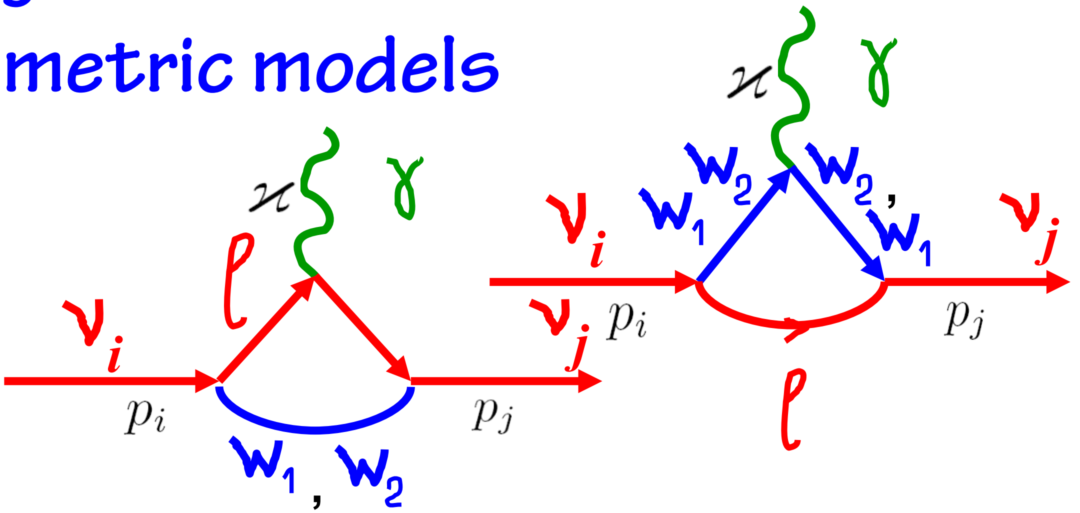
$$m_\nu \sim G\Lambda$$

$$m_\nu \sim \frac{\Lambda^2}{2m_e}\frac{\mu_\nu}{\mu_B} \sim \frac{\mu_\nu}{10^{-18}\mu_B}[\Lambda(\mathrm{TeV})]^2 \ \mathrm{eV}$$

Voloshin, 1988; Barr, Freire, Zee, 1990

3.6 Neutrino magnetic moment in left-right symmetric models

$SU_L(2) \times SU_R(2) \times U(1)$



Gauge bosons mass states $W_1 = W_L \cos\xi - W_R \sin\xi$

$W_2 = W_L \sin\xi + W_R \cos\xi$

with mixing angle ξ of gauge bosons $W_{L,R}$ with pure $(V \pm A)$ couplings

Kim, 1976; Marciano, Sanda, 1977
Beg, Marciano, Ruderman, 1978

$$\mu_{\nu_l} = \frac{eG_F}{2\sqrt{2}\pi^2}\left[m_l\left(1 - \frac{m_{W_1}^2}{m_{W_2}^2}\right)\sin 2\xi + \frac{3}{4}m_{\nu_l}\left(1 + \frac{m_{W_1}^2}{m_{W_2}^2}\right)\right]$$

... charged lepton mass ... ... neutrino mass ...

Large magnetic moment $\mu_\nu = \mu_\nu(m_\nu, m_{B^+}, m_{e^-})$

- In the L-R symmetric models $(SU(2)_L \times SU(2)_R \times U(1))$

Kim, 1976
Beg, Marciano, Ruderman, 1978

- Voloshin, 1988
  "On compatibility of small m_ν with large μ_ν of neutrino", Sov.J.Nucl.Phys. 48 (1988) 512
  … there may be $SU(2)_\nu$ symmetry that forbids m_ν but not μ_ν

Z.Z.Xing, Y.L.Zhou, "Enhanced electromagnetic transition dipole moments and radiative decays of massive neutrinos due to the seesaw-induced non-unitary effects" Phys.Lett.B 715 (2012) 178

- Bar, Freire, Zee, 1990
- supersymmetry
- extra dimensions

considerable enhancement of μ_ν to experimentally relevant range

- model-independent constraint μ_ν for BSM $(\Lambda \sim 1 \text{ TeV})$ without fine tuning and under assumption $\delta m_\nu \leq 1 \text{ eV}$

Bell, Cirigliano, Ramsey-Musolf, Vogel, Wise, 2005

! $\mu_\nu^D \leq 10^{-15}\mu_B$ Dirac versus Majorana $\mu_\nu^M \leq 10^{-14}\mu_B$!

REVIEWS OF MODERN PHYSICS, VOLUME 87, APRIL–JUNE 2015

REVIEWS OF MODERN PHYSICS, VOLUME 87, APRIL–JUNE 2015

Neutrino electromagnetic interactions: A window to new physics

+ upgrade:

Studenikin, Overview of neutrino electromagnetic properties 2022, arXiv:2301.06071

Studenikin, Overview of neutrino electromagnetic properties (the theory, laboratory experiments and astrophysical probes), PoS (NuFact2021) 402(2022)052

Detailed review on emp of ν

... > 500 citations

Carlo Giunti*
INFN, Torino Section, Via P. Giuria 1, I-10125 Torino, Italy

Alexander Studenikin†
Department of Theoretical Physics, Faculty of Physics, Moscow State University and Joint Institute for Nuclear Research, Dubna, Russia

(published 16 June 2015)

A review is given of the theory and phenomenology of neutrino electromagnetic interactions, which provide powerful tools to probe the physics beyond the standard model. After a derivation of the general structure of the electromagnetic interactions of Dirac and Majorana neutrinos in the one-photon approximation, the effects of neutrino electromagnetic interactions in terrestrial experiments and in astrophysical environments are discussed. The experimental bounds on neutrino electromagnetic properties are presented and the predictions of theories beyond the standard model are confronted.

DOI: 10.1103/RevModPhys.87.531 PACS numbers: 14.60.St, 13.15.+g, 13.35.Hb, 14.60.Lm

CONTENTS

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Neutrino magnetic moment: a window to new physics

A. Studenikin[a] *
[a]Department of Theoretical Physics, Moscow State University, 119991 Moscow, Russia

A short review on a neutrino magnetic moment is presented.

Introduction. Experimental and theoretical studies of flavour conversion in solar, atmospheric, reactor and accelerator neutrino fluxes give strong evidence of non-zero neutrino mass. A massive neutrino can have non-trivial electromagnetic properties [1]. For a recent review on neutrino electromagnetic properties see [2].

The neutrino dipole magnetic moment (along with the electric dipole moment) is the most well studied among neutrino electromagnetic properties. The effective Lagrangian, that is in charge of a neutrino coupling to the electromagnetic field, can be written in the form

$$L_{int} = \frac{1}{2}\bar{\psi}_i\sigma_{\alpha\beta}(\mu_{ij} + \epsilon_{ij}\gamma_5)\psi_j F^{\alpha\beta} + h.c. \quad (1)$$

where the magnetic moments μ_{ij}, in the presence of mixing between different neutrino states, are associated with the neutrino mass eigenstates ν_i. The interplay between magnetic moment and neutrino mixing effects is important. Note that electric (transition) moments ϵ_{ij} do also contribute to the coupling.

A Dirac neutrino may have non-zero diagonal electric moments in models where CP invariance is violated. For a Majorana neutrino the diagonal magnetic and electric moments are zero. Therefore, neutrino magnetic moments can be used to distinguish Dirac and Majorana neutrinos (see [3] and also [2] for a detailed discussion).

Neutrino magnetic moment in a minimal extension of Standard Model. The explicit evaluation of the one-loop contributions to the Dirac neutrino magnetic moment in the leading approximation over small parameters $b_i = \frac{m_i^2}{M_W^2}$ (m_i are the neutrino masses, $i = 1, 2, 3$), that however exactly accounts for $a_l = \frac{m_l^2}{M_W^2}$ ($l = e, \mu, \tau$), leads the following result [4],

$$\mu_{ij}^D = \frac{eG_F m_i}{8\sqrt{2}\pi^2}\Big(1+\frac{m_j}{m_i}\Big)\sum_{l=e,\mu,\tau} f(a_l)U_{lj}U_{li}^*, \quad (2)$$

$$f(a_l) = \frac{3}{4}\Big[1+\frac{1}{1-a_l} - \frac{2a_l}{(1-a_l)^2} - \frac{2a_l^2\ln a_l}{(1-a_l)^3}\Big],$$

where U_{li} is the neutrino mixing matrix. The correspondent result in the absence of mixing was confirmed in [5,6]. A Majorana neutrino may also have transition moment of the value $\mu_{ij}^M = 2\mu_{ij}^D$ (see [2] for a detailed discussion and references).

For the diagonal magnetic moment of the Dirac neutrino, from (2) in the limit $a_l \ll 1$ the result [1] can be obtained

$$\mu_{ii}^D = \frac{3eG_F m_i}{8\sqrt{2}\pi^2}\Big(1 - \frac{1}{2}\sum_{l=e,\mu,\tau} a_l \mid U_{li}\mid^2\Big). \quad (3)$$

The magnetic moment for hypothetical heavy neutrino was studied in [6]. In particular, it was obtained

$$\mu_\nu = \frac{eG_F m_\nu}{8\sqrt{2}\pi^2}\begin{cases} 3+\frac{5}{6}b, & m_\ell \ll m_\nu \ll M_W, \\ 1, & m_\ell \ll M_W \ll m_\nu. \end{cases} \quad (4)$$

Note that the LEP data set a limit on number of light neutrinos coupled to Z boson.

The numerical value of the Dirac neutrino magnetic moment within a minimal extension of the Standard Model, as it follows from (3), is

$$\mu_{ii}^D \approx 3.2\times 10^{-19}\Big(\frac{m_i}{1\ eV}\Big)\mu_B, \quad (5)$$

This is several orders of magnitude smaller than the present experimental limits if to account for the existed constraints on neutrino masses.

*e-mail: studenik@srd.sinp.msu.ru

… A remark on electric charge of ν …

Beyond Standard Model

ν neutrality $Q=0$ is attributed to

gauge invariance + anomaly cancellation constraints

imposed in SM of electroweak interactions

- …General proof: $SU(2)_L \times U(1)_Y$

In SM: $Q = I_3 + \frac{Y}{2}$ …from Gell-Mann – Nishijima …

Foot, Joshi, Lew, Volkas, 1990;
Foot, Lew, Volkas, 1993;
Babu, Mohapatra, 1989, 1990
Foot, He (1991)

In SM (without ν_R) triangle anomalies cancellation constraints → certain relations among particle hypercharges that is enough to fix all Y so that they, and consequently Q, are quantized

- $Q=0$ is proven also by direct calculation in SM within different gauges and methods

$Q=0$

Bardeen, Gastmans, Lautrup, 1972;
Cabral-Rosetti, Bernabeu, Vidal, Zepeda, 2000;
Beg, Marciano, Ruderman, 1978;
Marciano, Sirlin, 1980; Sakakibara, 1981;

- Dvornikov, Studenikin, 2004 (for SM in one-loop calculations)

- … Strict requirements for Q quantization may disappear in extensions of standard $SU(2)_L \times U(1)_Y$ EW model if ν_R with $Y \neq 0$ are included: in the absence of Y quantization electric charges Q gets dequantized → millicharged ν

ν charge radius and anapole moment

$$\Lambda_\mu(q) = f_Q(q^2)\gamma_\mu + f_M(q^2)i\sigma_{\mu\nu}q^\nu - f_E(q^2)\sigma_{\mu\nu}q^\nu\gamma_5 + f_A(q^2)(q^2\gamma_\mu - q_\mu \not{q})\gamma_5$$

electric — magnetic — dipole electric — anapole

Although it is usually assumed that ν are electrically neutral (charge quant. Implies $Q \sim \frac{1}{3}e$), ν can be characterized by two ± charge distributions

$$f_Q(q^2) = f_Q(0) + q^2 \frac{df_Q}{dq^2}(0) + \cdots,$$ and $f_Q(q^2) \neq 0 \; for \; q^2 \neq 0$ even for electric charge $f_Q(0) = 0$

ν charge radius is introduced as $\langle r_\nu^2 \rangle = +6 \frac{df_Q}{dq^2}(0)$

- … it is often claimed for SM massless ν anapole moment $a_\nu = f_A(q^2) = \frac{1}{6}\langle r_\nu^2 \rangle$? ? ?

- for two-component massless left-handed Weyl spinors of SM

to be correct ⟹

$$\Lambda^{Q,A}_{\mathrm{SM}\mu}(q) = (\gamma_\mu q^2 - q_\mu \not{q})\mathfrak{f}^{\mathrm{SM}}(q^2)$$

Giunti, Studenikin Rev.Mod.Phys.2015

$$\mathfrak{f}^{\mathrm{SM}}(q^2) = \tilde{\mathfrak{f}}_Q(q^2) - \mathfrak{f}_A(q^2) \xrightarrow[q^2\to 0]{} \frac{\langle r^2 \rangle}{6} - \mathfrak{a}$$

… in SM charge radius and anapole moment are not defined separately …

Interpretation of charge radius as an observable is rather delicate issue: $\langle r_\nu^2 \rangle$ represents a correction to tree-level electroweak scattering amplitude between ν and charged particles, which receives radiative corrections from several diagrams (including γ exchange) to be considered simultaneously ⟹ calculated CR is infinite and gauge dependent quantity. For ν with m=0, $\langle r_\nu^2 \rangle$ and a_ν can be defined (finite and gauge independent) from scattering cross section.

Bernabeu, Papavassiliou, Vidal, Nucl.Phys. B 680 (2004) 450

? ? ? For massive ν ? ? ?

The definition of the neutrino charge radius follows an analogy with the elastic electron scattering off a static spherically symmetric charged distribution of density $\rho(r)$ ($r = |\mathbf{x}|$), for which the differential cross section is determined [79–81] by the point particle cross section $\frac{d\sigma}{d\Omega}|_{point}$,

$$\frac{d\sigma}{d\Omega} = \frac{d\sigma}{d\Omega}_{|point} |f(q^2)|^2, \tag{90}$$

where the correspondent form factor $f(q^2)$ in the so-called *Breit frame*, in which $q_0 = 0$, can be expressed as

$$f(q^2) = \int \rho(r) e^{i\mathbf{qx}} d^3x = 4\pi \int dr r^2 \rho(r) \frac{\sin(qr)}{qr}, \tag{91}$$

here $q = |\mathbf{q}|$. Thus, one has

$$\frac{df_Q}{dq^2} = \int \rho(r) \frac{qr\cos(qr) - \sin(qr)}{2q^{3/2} r} d^3x. \tag{92}$$

In the case of small q, we have $\lim_{q^2 \to 0} \frac{qr\cos(qr) - \sin(qr)}{2q^{3/2} r} = -\frac{r^2}{6}$ and

$$f(q^2) = 1 - |\mathbf{q}|^2 \frac{\langle r^2 \rangle}{6} + \ \dots \ . \tag{93}$$

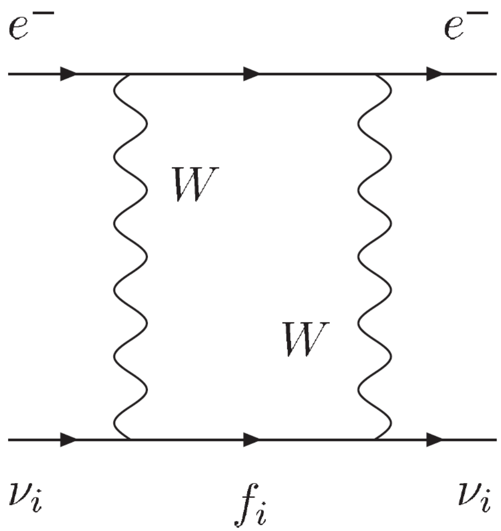
Therefore, the neutrino charge radius (in fact, it is the charge radius squared) is usually defined by

$$\langle r_\nu^2 \rangle = -6 \frac{df_Q(q^2)}{dq^2}|_{q^2=0}. \tag{94}$$

Since the neutrino charge density is not a positively defined quantity, $\langle r_\nu^2 \rangle$ can be negative.

To obtain ν electroweak radius as physical (finite, not divergent) quantity

Bernabeu, Papavassiliou, Vidal, 2004



$$\langle r^2_{\nu_i} \rangle = \frac{G_F}{4\sqrt{2}\pi^2}\left[3 - 2\log\left(\frac{m_i^2}{m_W^2}\right)\right] \qquad i = e, \mu, \tau$$

$$\langle r^2_{\nu_e} \rangle = 4 \times 10^{-33}\,\mathrm{cm}^2$$

…contribution to ν - e scattering experiments through

Contribution of box diagram to

$$\nu_l + l' \to \nu_l + l'.$$

$$g_V \to \frac{1}{2} + 2\sin^2\theta_W + \frac{2}{3} m_W^2 \langle r^2_{\nu_e} \rangle \sin^2\theta_W$$

… theoretical predictions and present experimental limits are in agreement within one order of magnitude…

Experimental constraints

(from studies of terrestrial and astrophysical ν fluxes)

on μ_ν , q_ν and $<r_\nu^2>$

magnetic moment millicharge charge radius

Particle Data Group

Review of Particle Physics 2022 and update 2023

R.L.Workman et al.,
Progress of Theoretical and Experimental Physics,
vol. 2022, no. 8, 083C01

ν magnetic moment

... most easily accepted are
dipole magnetic and electric moments

however most accessible for experimental studies are charge radii $<r_\nu^2>$

Studies of ν-e scattering - most sensitive method for experimental investigation of μ_ν

Cross-section:

- $$\frac{d\sigma}{dT}(\nu + e \to \nu + e) = \left(\frac{d\sigma}{dT}\right)_{SM} + \left(\frac{d\sigma}{dT}\right)_{\mu_\nu}$$

where the Standard Model contribution

- $$\left(\frac{d\sigma}{dT}\right)_{SM} = \frac{G_F^2 m_e}{2\pi}\left[(g_V + g_A)^2 + (g_V - g_A)^2\left(1 - \frac{T}{E_\nu}\right)^2 + (g_A^2 - g_V^2)\frac{m_e T}{E_\nu^2}\right],$$

T is the electron recoil energy and

- $$\left(\frac{d\sigma}{dT}\right)_{\mu_\nu} = \frac{\pi\alpha_{em}^2}{m_e^2}\left[\frac{1 - T/E_\nu}{T}\right]\mu_\nu^2$$

$$\mu_\nu^2(\nu_l, L, E_\nu) = \sum_j \left|\sum_i U_{li} e^{-iE_i L}\mu_{ji}\right|^2$$

$$\mu_{ij} \to |\mu_{ij} - \epsilon_{ij}|$$

$$g_V = \begin{cases} 2\sin^2\theta_W + \frac{1}{2} & \text{for } \nu_e, \\ 2\sin^2\theta_W - \frac{1}{2} & \text{for } \nu_\mu, \nu_\tau, \end{cases} \qquad g_A = \begin{cases} \frac{1}{2} & \text{for } \nu_e, \\ -\frac{1}{2} & \text{for } \nu_\mu, \nu_\tau \end{cases}$$

for anti-neutrinos $g_A \to -g_A$

- to incorporate charge radius: $g_V \to g_V + \frac{2}{3}M_W^2\langle r^2\rangle\sin^2\theta_W$???

- $$\frac{d\sigma}{dT}(\nu + e \to \nu + e) = \left(\frac{d\sigma}{dT}\right)_{\rm SM} + \left(\frac{d\sigma}{dT}\right)_{\mu_\nu}$$

ν-γ coupling ... valid for ν scattering on free e

- $$\left(\frac{d\sigma}{dT}\right)_{\mu_\nu} = \frac{\pi\alpha_{em}^2}{m_e^2}\left[\frac{1 - T/E_\nu}{T}\right]\mu_\nu^2$$

with change of helicity, contrary to SM

T is the electron recoil energy: $0 \le T \le \frac{2E_\nu^2}{2E_\nu + m_e}$

If neutrino has electric dipole moment,
or electric or magnetic transition moments,
these quantities also contribute to scattering cross section

$$\mu_{\nu_\ell}^2 \simeq \mu_{\bar{\nu}_\ell}^2 \simeq \sum_j \left| \sum_k U_{\ell k}^* (\mu_{jk} - i\epsilon_{jk}) \right|^2$$

effective flavour magnetic moment for short-baseline experiments

Kouzakov Studenikin PRD 2017

Possibility of *distractive interference* between magnetic and electric transition moments of Dirac neutrino (Majorana neutrino has only magnetic or electric transition moment, but not both if CP is conserved)

Effective ν_e magnetic moment measured in ν-e scattering experiments ?

$$\mu_e^2$$

Two steps:

1) consider ν_e as superposition of mass eigenstates ($i=1,2,3$) at some distance L from the source, and then sum up magnetic moment contributions to ν-e scattering amplitude (of each of mass components) induced by their magnetic moments

$$A_j \sim \sum_i U_{ei} e^{-iE_i L} \mu_{ji}$$

J.Beacom, P.Vogel, 1999

2) amplitudes combine incoherently in total cross section

$$\sigma \sim \mu_e^2 = \sum_j \left| \sum_i U_{ei} e^{-iE_i L} \mu_{ji} \right|^2$$

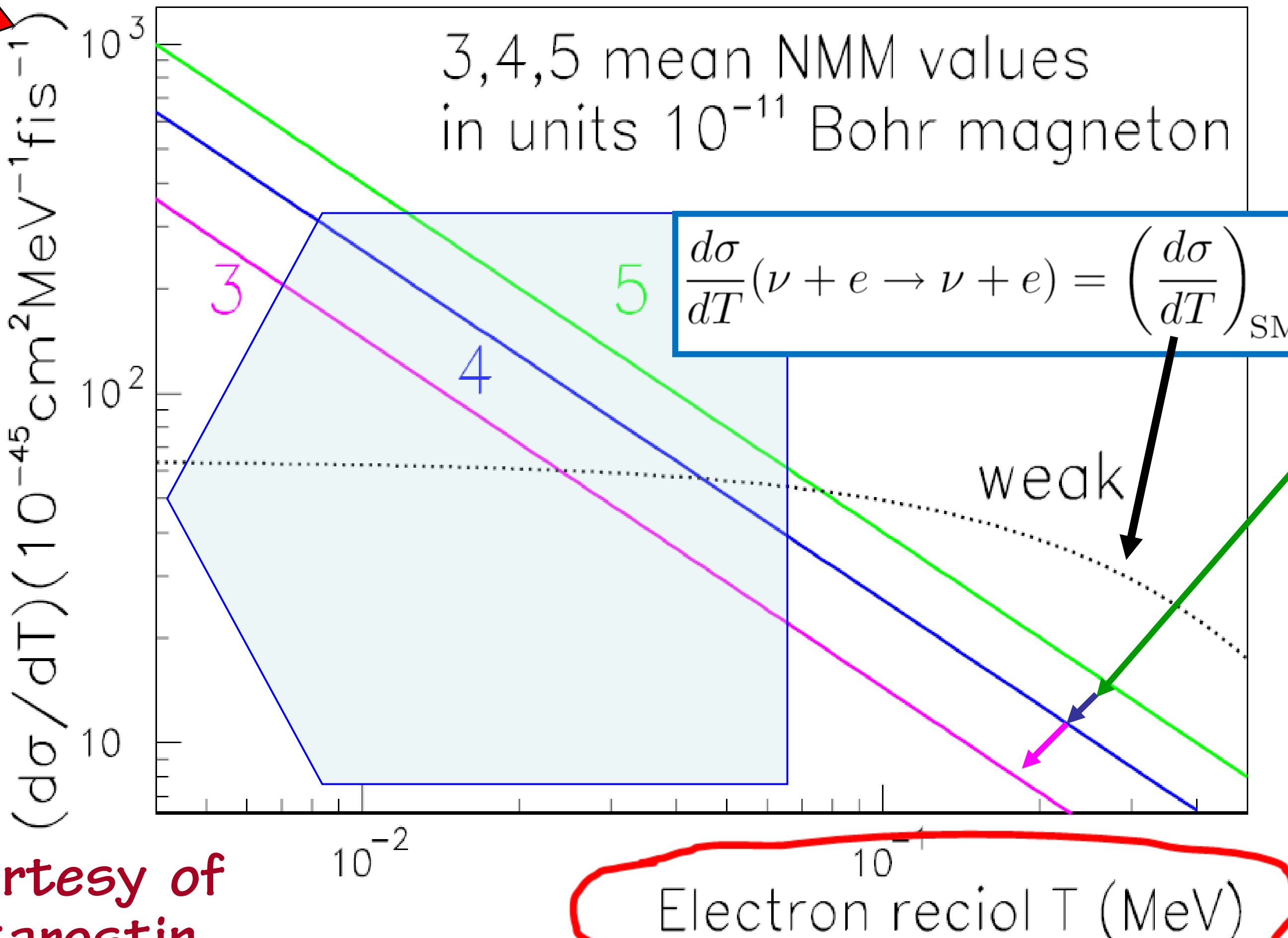
C.Giunti, A.Studenikin, 2009

K.Kouzakov, A.Studenikin, 20018

NB! Summation over $j=1,2,3$ is outside the square because of incoherence of different final mass states contributions to cross section

Magnetic moment contribution dominates at low electron recoil energies when $\left(\frac{d\sigma}{dT}\right)_{\mu_\nu} > \left(\frac{d\sigma}{dT}\right)_{SM}$ and $\frac{T}{m_e} < \frac{\pi^2 \alpha_{em}}{G_F^2 m_e^4} \mu_\nu^2$

… the lower the smallest measurable electron recoil energy is, smaller values of μ_ν^2 can be probed in scattering experiments …



$$\frac{d\sigma}{dT}(\nu + e \to \nu + e) = \left(\frac{d\sigma}{dT}\right)_{SM} + \left(\frac{d\sigma}{dT}\right)_{\mu_\nu}$$

… courtesy of A.Starostin…

Калининской атомной станции

(Удомля, Тверская область)

•

GEMMA (2005 – 2012 - running)
Germanium Experiment for Measurement of Magnetic Moment of Antineutrino

JINR (Dubna) + ITEP (Moscow) at Kalinin Nuclear Power Plant

! World best experimental (reactor) limit

$$\mu_\nu < 2.9 \times 10^{-11} \mu_B$$

June 2012 ●

A.Beda et al, in:
Special Issue on "Neutrino Physics",
Advances in High Energy Physics (2012) 2012,
editors: J. Bernabeu, G. Fogli, A. McDonald, K. Nishikawa

… quite realistic prospects for future …

● GEMMA-2 / νGeN experiment

… searching for μ_ν and CEνNS unprecedentedly low threshold $T \sim 200 \quad eV$

$$\mu_\nu \sim (5-9) \times 10^{-12} \mu_B$$

2023 + to appear soon ?

● first results on CEνNS : I.Alekseev etal., Phys.Rev.D 106 (2022) 5, L051101

coherent elastic neutrino–nucleus scattering

… claim that

ν-e cross section

should be increased by

Atomic Ionization effect:

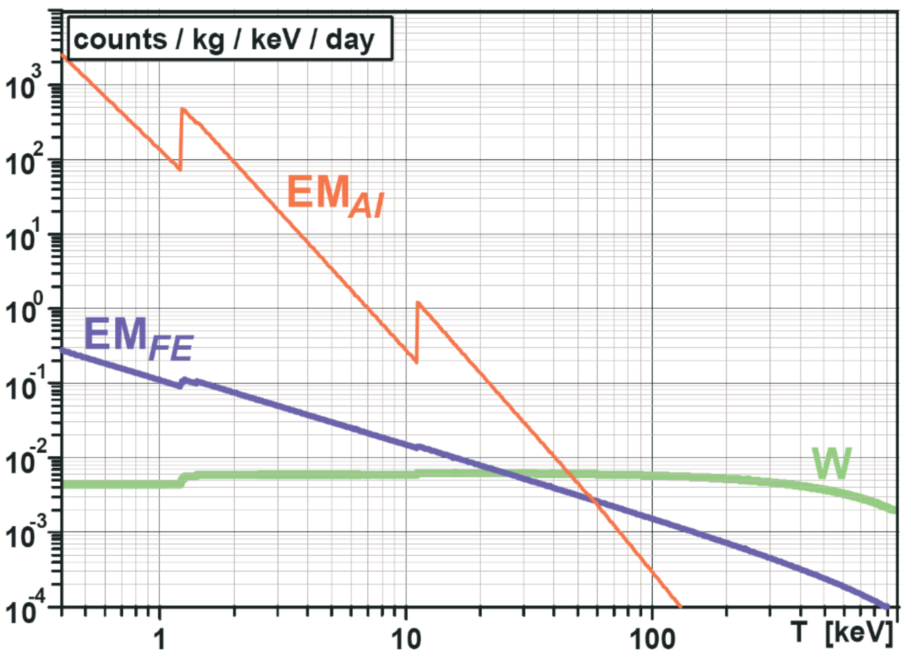
$$\nu + (A,Z) \longrightarrow \nu' + (A,Z)^{+} + e^{-}$$
$$\downarrow \text{recombination}$$
$$(A,Z) \quad + \gamma$$

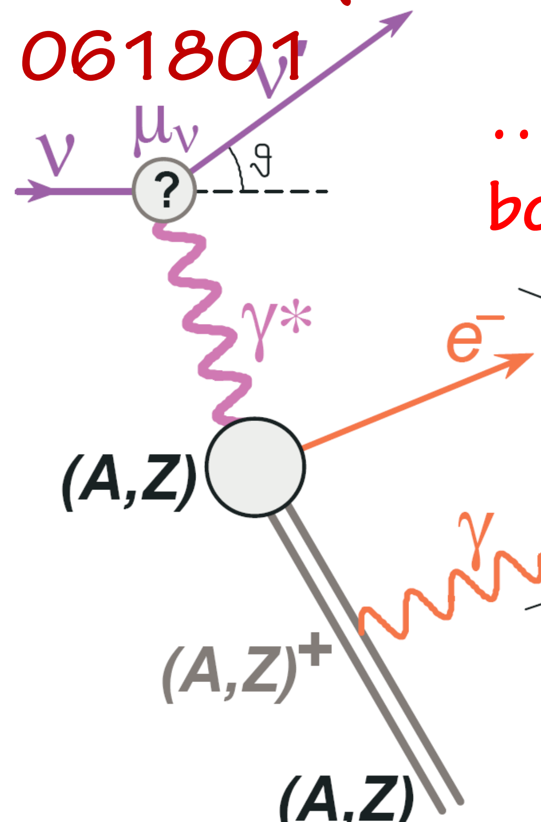
?

H.Wong et al. (TEXONO Coll.), arXiv: 1001.2074, 13 Jan 2010, reported at

Neutrino 2010 Conference (Athens, June 2010),

PRL 105 (2010) 061801





…new ⟹ bounds …

...much better limits on ν effective magnetic moment ...

H.Wong et al.,
(TEXONO Coll.),
arXiv: 1001.2074,
13 Jan 2010,

PRL 105 (2010) 061801

$$\mu_\nu < 1.3 \times 10^{-11} \mu_B$$?

... atomic ionization effect accounted for ...

Neutrino 2010 Conference, Athens

... however ... ⟹

A.Beda et al.
(GEMMA Coll.),
arXiv: 1005.2736,
16 May 2010

$$\mu_\nu < 5.0 \times 10^{-12} \mu_B$$?

... atomic ionization effect accounted for ...

$$\mu_\nu < 3.2 \times 10^{-11} \mu_B$$

... **ν**-*e* scattering on free electrons ...

(without atomic ionization)

K.Kouzakov, A.Studenikin

- Magnetic neutrino scattering on atomic electrons revisited, **Phys.Lett. B** 105 (2011) 061801,
- Electromagnetic neutrino-atom collisions: The role of electron binding, **Nucl.Phys.** (Proc.Suppl.) 217 (2011) 353

K.Kouzakov, A.Studenikin, M.Voloshin

- Neutrino electromagnetic properties and new bounds on neutrino magnetic moments, **J.Phys.: Conf.Ser.** 375 (2012) 042045
- Neutrino-impact ionization of atoms in search for neutrino magnetic moment, **Phys.Rev.D** 83 (2011) 113001
- On neutrino-atom scattering in searches for neutrino magnetic Moments, **Nucl.Phys.B** (Proc.Supp.) 2011 (Proc. of Neutrino 2010 Conf.)
- Testing neutrino magnetic moment in ionization of atoms by neutrino impact, **JETP Lett.** 93 (2011) 699

M.Voloshin

- Neutrino scattering on atomic electrons in search for neutrino magnetic moment, **Phys.Rev.Lett.** 105 (2010) 201801

No important effect of Atomic Ionization on cross section in μ_ν experiments once all possible final electronic states accounted for

…free electron approximation …

M.Voloshin, 23 Aug 2010;
K.Kouzakov, A.Studenikin, 26 Nov 2010;
H.Wong et al, arXiv: 1001.2074 V3, 28 Nov 2010

K. Kouzakov, A. Studenikin,
"Theory of neutrino-atom collisions:
the history, present status, and BSM physics",

in: Special issue
"Through Neutrino Eyes: The Search for New Physics",
Adv. in High Energy Phys. 2014 (2014) 569409 (37pp)

editors: J. Bernabeu, G. Fogli, A. McDonald, K. Nishikawa

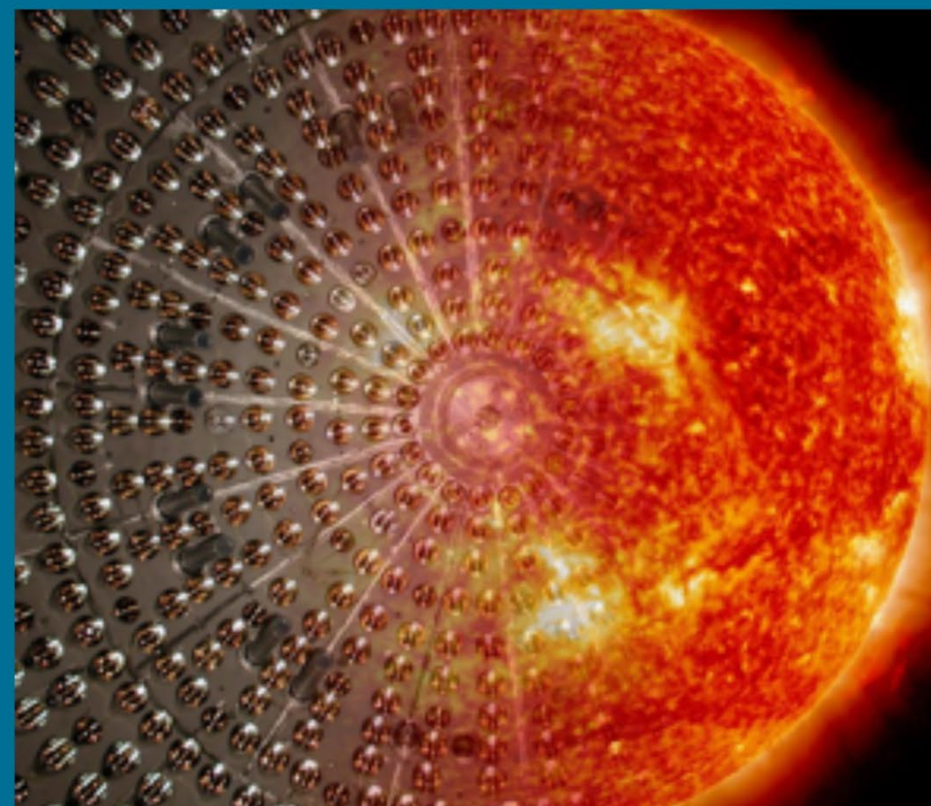
TAUP 2017 Topics in Astroparticle and Underground Physics

JÜLICH FORSCHUNGSZENTRUM

Limiting the effective magnetic moment of solar neutrinos with the Borexino detector

Livia Ludhova
on behalf of
the Borexino collaboration

IKP-2 FZ Jülich,
RWTH Aachen,
and JARA Institute, Germany



Phys. Rev. D 96 (2017) 091103
Limiting μ_ν with Borexino Phase-II solar neutrino data

NEW

BOREXINO Collaboration (2017)

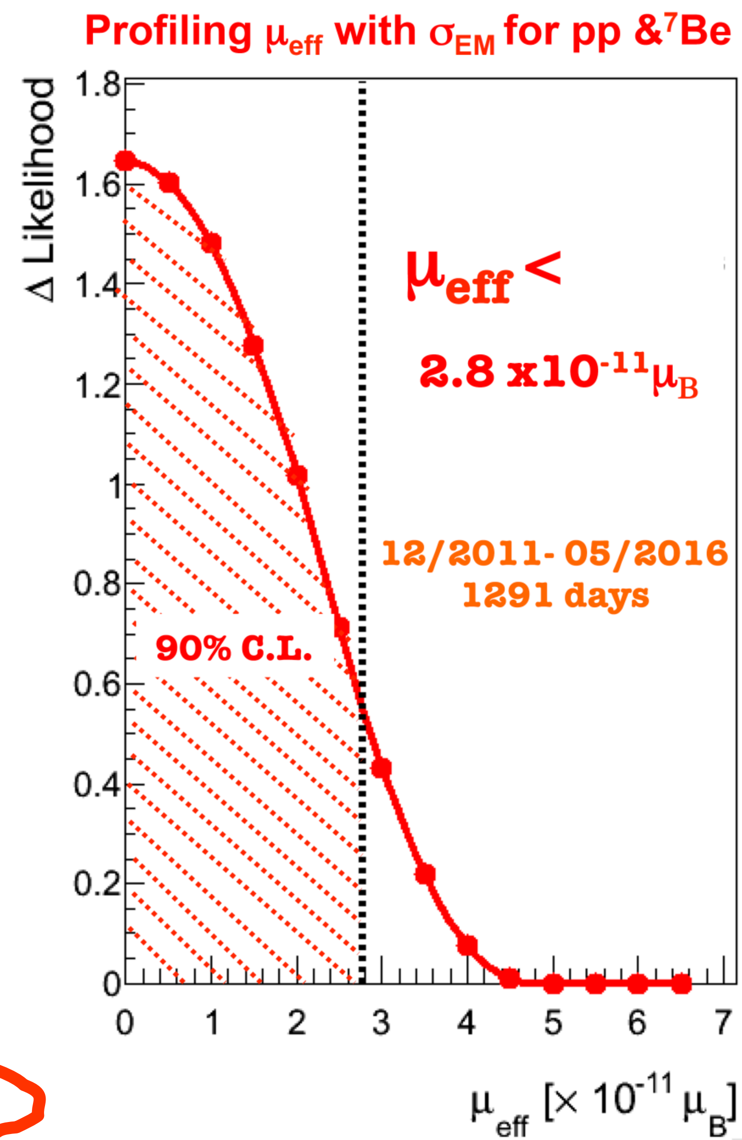
NMM results from Phase 2

JÜLICH FORSCHUNGSZENTRUM

Data selection:
Fiducial volume: R < 3.021 m, |z| < 1.67 m
Muon, ^{214}Bi-^{214}Po, and noise suppression
Free fit parameters: solar-ν (pp, ^{7}Be) and **backgrounds** (^{85}Kr, ^{210}Po, ^{210}Bi, ^{11}C, external bgr.), **response parameters** (light yield, ^{210}Po position and width, ^{11}C edge (2 x 511 keV), 2 energy resolution parameters)
Constrained parameters: ^{14}C, pile up
Fixed parameters: pep-, CNO-, ^{8}B-ν rates
Systematics: treatment of pile-up, energy estimators, pep and CNO constraints with LZ and HZ SSM

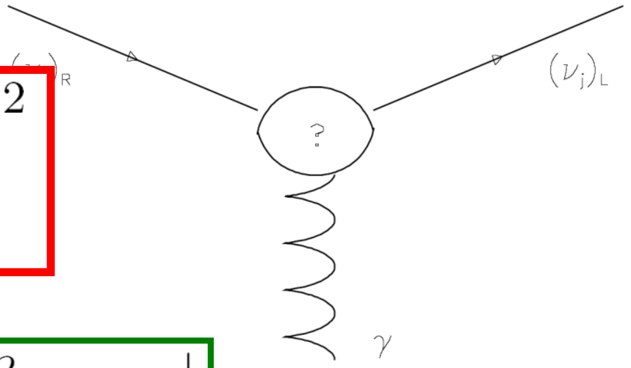
Without radiochemical constraint
$\mu_{eff} < 4.0 \times 10^{-11} \mu_B$ (90% C.L.)

With radiochemical constraint
$\mu_{eff} < 2.6 \times 10^{-11} \mu_B$ (90% C.L.)
adding systematics
$\mu_{eff} < 2.8 \times 10^{-11} \mu_B$ (90% C.L.)



Effective ν magnetic moment in experiments

(for neutrino produced as ν_l with energy E_ν
and after traveling a distance L)



$$\mu_\nu^2(\nu_l, L, E_\nu) = \sum_j \left| \sum_i U_{li} e^{-iE_iL} \mu_{ji} \right|^2$$

where **neutrino mixing matrix** $\quad \mu_{ij} \equiv |\beta_{ij} - \varepsilon_{ij}|$

magnetic and electric moments

- Observable μ_ν is an effective parameter that depends on neutrino flavour composition at the detector.

- **Implications of μ_ν limits from different experiments (reactor, solar $^8\mathrm{B}$ and $^7\mathrm{Be}$) are different.**

... comprehensive analysis of ν-e scattering with account for ν mixing and oscillations ...

PHYSICAL REVIEW D **95,** 055013 (2017)

Electromagnetic properties of massive neutrinos in low-energy elastic neutrino-electron scattering

Konstantin A. Kouzakov*

Department of Nuclear Physics and Quantum Theory of Collisions, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia

Alexander I. Studenikin†

Department of Theoretical Physics, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia and Joint Institute for Nuclear Research, Dubna 141980, Moscow Region, Russia

(Received 11 February 2017; published 14 March 2017)

A thorough account of electromagnetic interactions of massive neutrinos in the theoretical formulation of low-energy elastic neutrino-electron scattering is given. The formalism of neutrino charge, magnetic, electric, and anapole form factors defined as matrices in the mass basis is employed under the assumption of three-neutrino mixing. The flavor change of neutrinos traveling from the source to the detector is taken into account and the role of the source-detector distance is inspected. The effects of neutrino flavor-transition millicharges and charge radii in the scattering experiments are pointed out.

DOI: 10.1103/PhysRevD.95.055013

Short-baselin case $L \ll L_{kk'} = 2E_\nu/|\delta m^2_{kk'}|$ → $e^{-i(\delta m^2_{kk'}/2E_\nu)L} = 1$

- $P_{\nu_\ell \to \nu_e}(L, E_\nu) = \delta_{\ell e}$ $\qquad \mathcal{A}_{\nu_\ell \to \nu_{\ell'}}(L, E_\nu)\mathcal{A}^*_{\nu_\ell \to \nu_{\ell''}}(L, E_\nu) = \delta_{\ell\ell'}\delta_{\ell\ell''}$

effect of ν flavor change is insignificant
($\nu_\ell(L)$ is as in the source)

$$C_1 = (g_V + \delta_{\ell e} + \tilde{Q}_{\ell\ell})^2 + \sum_{\ell'=e,\mu,\tau} (1 - \delta_{\ell'\ell}) \left|\tilde{Q}_{\ell'\ell}\right|^2 \qquad C_2 = (g_A + \delta_{\ell e})^2$$

$$C_3 = (g_V + \delta_{\ell e})(g_A + \delta_{\ell e}) + (g_A + \delta_{\ell e})\tilde{Q}_{\ell\ell}$$

weak-electromagnetic interference term contains only flavour-diagonal millicharges and charge radii

- Effective magnetic moment

$$|\mu_\nu(L, E_\nu)|^2 = \sum_{j=1}^{3} \sum_{k,k'=1}^{3} U^*_{\ell k} U_{\ell k'} (\mu_\nu)_{jk} (\mu_\nu)^*_{jk'} = \sum_{\ell'=e,\mu,\tau} |(\mu_\nu)_{\ell'\ell}|^2$$ where

$$(\mu_\nu)_{\ell'\ell} = \sum_{j,k=1}^{3} U^*_{\ell k} U_{\ell' j} (\mu_\nu)_{jk}$$ is the effective magnetic moment in flavor basis

for GEMMA experiment ...

- Long-baselin case $L \gg L_{kj} = 2E_\nu / |\delta m^2_{kk'}|$ $\exp(-i\delta m^2_{kk'}/2E_\nu) = \delta_{kk'}$

effect of decoherence

$$C_1 = g_V^2 + 2g_V P_{\nu_\ell \to \nu_e} + P_{\nu_\ell \to \nu_e} + \sum_{j,k=1}^{3} |U_{\ell k}|^2 \left|\tilde{Q}_{jk}\right|^2 + 2g_V \sum_{j=1}^{3} |U_{\ell j}|^2 \tilde{Q}_{jj} + 2 \sum_{j,k=1}^{3} |U_{\ell k}|^2 \mathrm{Re}\left\{U_{ej} U_{ek}^* \tilde{Q}_{jk}\right\}$$

$$C_2 = g_A^2 + 2g_A P_{\nu_\ell \to \nu_e} + P_{\nu_\ell \to \nu_e}$$

$$C_3 = g_V g_A + (g_V + g_A + 1) P_{\nu_\ell \to \nu_e} + g_A \sum_{j=1}^{3} |U_{\ell j}|^2 \tilde{Q}_{jj} + 2 \sum_{j,k=1}^{3} |U_{\ell k}|^2 U_{ej} U_{ek}^* \tilde{Q}_{jk}$$

where the flavour transition probability $P_{\nu_\ell \to \nu_e} = \sum_{k=1}^{3} |U_{\ell k}|^2 |U_{ek}|^2$
does not depend on source-detector distance and ν energy

- Effective magnetic moment $|\mu_\nu(L, E_\nu)|^2 = \sum_{j,k=1}^{3} |U_{\ell k}|^2 \left|(\mu_\nu)_{jk}\right|^2$
is independent of *L* and *E*

- for Borexino experiment . . .

Bounds on millicharge q_ν from μ_ν (GEMMA Coll. data)

2

two not seen contributions:

$$\left(\frac{d\sigma}{dT}\right)_{\mu_\nu^a} \approx \pi\alpha^2 \frac{1}{m_e^2 T}\left(\frac{\mu_\nu^a}{\mu_B}\right)^2$$

$$\left(\frac{d\sigma}{dT}\right)_{q_\nu} \approx 2\pi\alpha \frac{1}{m_e T^2} q_\nu^2$$

ν-e cross-section

$$\left(\frac{d\sigma}{dT}\right)_{\nu-e} = \left(\frac{d\sigma}{dT}\right)_{SM} + \left(\frac{d\sigma}{dT}\right)_{\mu_\nu} + \left(\frac{d\sigma}{dT}\right)_{q_\nu}$$

Bounds on q_ν from ... unobserved effects of New Physics

$$R = \frac{\left(\frac{d\sigma}{dT}\right)_{q_\nu}}{\left(\frac{d\sigma}{dT}\right)_{\mu_\nu^a}} = \frac{2m_e}{T}\frac{\left(\frac{q_\nu}{e_0}\right)^2}{\left(\frac{\mu_\nu^a}{\mu_B}\right)^2} \lesssim 1$$

!

Studenikin, Europhys. Lett. 107 (2014) 210011

Particle Data Group, 2016-2022 and update of 2023

Expected new constraints from GEMMA:

now $\mu_\nu < 2.9 \times 10^{-11} \mu_B$ ($T \sim 2.8\ keV$)

2023+ few years data taking

νGeN experiment

$\mu_\nu \sim (5-9) \times 10^{-12} \mu_B$

... low threshold ...

$T \sim 200\ eV$

Constraints on q_ν !

- $|q_\nu| < 1.5 \times 10^{-12} e_0$

in ν Table of Particle Data Group since 2016

- $|q_\nu| < 1.1 \times 10^{-13} e_0$

Particle Data Group collaboration 2016 – 2022 and 2023 update

PDG particle data group **Particle Listings**

Live | Summary Tables | Reviews, Tables, Plots | Particle Listings

2017 Review of Particle Physics

Please use this **CITATION**:
C. Patrignani *et al.* (Particle Data Group), Chin. Phys. C, **40**, 100001 (2016) and 2017 update.

Cut-off date for this update was January 15, 2017.

Particle Listings

Search Listings

Gauge & Higgs Bosons (gamma, g, W, Z, ...)

Leptons (e, mu, tau, neutrinos, heavy leptons ...)

Quarks (u, d, s, c, b, t, ...)

Mesons (pi, K, D, B, psi, Upsilon, ...)

Baryons (p, n, Lambda_b, Xi, ...)

Other Searches (SUSY, Compositeness, ...)

ν CHARGE

VALUE (units: electron charge)	CL%	DOCUMENT ID		TECN	COMMENT
• • • We do not use the following data for averages, fits, limits, etc. • • •					
$<3 \times 10^{-8}$	95	[1] DELLA-VALLE	16	PVLA	Magnetic dichroism
$<2.1 \times 10^{-12}$	90	[2] CHEN	14A	TEXO	Nuclear reactor
$<1.5 \times 10^{-12}$	90	[3] STUDENIKIN	14		Nuclear reactor
$<3.7 \times 10^{-12}$	90	[4] GNINENKO	07	RVUE	Nuclear reactor
$<2 \times 10^{-14}$		[5] RAFFELT	99	ASTR	Red giant luminosity
$<6 \times 10^{-14}$		[6] RAFFELT	99	ASTR	Solar cooling
$<4 \times 10^{-4}$		[7] BABU	94	RVUE	BEBC beam dump
$<3 \times 10^{-4}$		[8] DAVIDSON	91	RVUE	SLAC e^- beam dump
$<2 \times 10^{-15}$		[9] BARBIELLINI	87	ASTR	SN 1987A
$<1 \times 10^{-13}$		[10] BERNSTEIN	63	ASTR	Solar energy losses

[1] DELLA-VALLE 16 obtain a limit on the charge of neutrinos valid for masses of less than 10 meV. For heavier neutrinos the limit increases as a power of mass, reaching 10^{-6} e for $m = 100$ meV.

[2] CHEN 14A use the Multi-Configuration RRPA method to analyze reactor $\overline{\nu}_e$ scattering on Ge atoms with 300 eV recoil energy threshold to obtain this limit.

[3] STUDENIKIN 14 uses the limit on μ_ν from BEDA 13 and the 2.8 keV threshold of the electron recoil energy to obtain this limit.

- Studenikin, New bounds on neutrino electric millicharge from limits on neutrino magnetic moment, Europhysics Letters 107 (2014) 21001

②

Experimental limits for different effective q_ν

C. Giunti, A. Studenikin, Electromagnetic interactions of neutrinos: a window to new physics, Rev. Mod. Phys. 87 (2015) 531

Limit	Method	Reference
$\lvert q_{\nu_\tau} \rvert \lesssim 3 \times 10^{-4}\, e$	SLAC e^- beam dump	Davidson *et al.* (1991)
$\lvert q_{\nu_\tau} \rvert \lesssim 4 \times 10^{-4}\, e$	BEBC beam dump	Babu *et al.* (1994)
$\lvert q_{\nu} \rvert \lesssim 6 \times 10^{-14}\, e$	Solar cooling (plasmon decay)	Raffelt (1999a)
$\lvert q_{\nu} \rvert \lesssim 2 \times 10^{-14}\, e$	Red giant cooling (plasmon decay)	Raffelt (1999a)
$\lvert q_{\nu_e} \rvert \lesssim 3 \times 10^{-21}\, e$	● Neutrality of matter ●	Raffelt (1999a)
$\lvert q_{\nu_e} \rvert \lesssim 3.7 \times 10^{-12}\, e$	Nuclear reactor	Gninenko *et al.* (2007)
$\lvert q_{\nu_e} \rvert \lesssim 1.5 \times 10^{-12}\, e$	Nuclear reactor	Studenikin (2013)

A. Studenikin: New bounds on neutrino electric millicharge from limits on neutrino magnetic moment, Eur.Phys.Lett. 107 (2014) 2100

… since that C.Patrignani et al (Particle Data Group), The Review of Particle Physics 2016 Chinese Physics C 40 (2016) 100001

ν charge radii

. . . most accessible for experimental studies are charge radii $<r_{\nu}^2>$

Bernabeu, Papavassiliou, Vidal, 2004

. . . astrophysical bounds ???

... comprehensive analysis of ν-e scattering ...

PHYSICAL REVIEW D **95**, 055013 (2017)

Electromagnetic properties of massive neutrinos in low-energy elastic neutrino-electron scattering

Konstantin A. Kouzakov*

Department of Nuclear Physics and Quantum Theory of Collisions, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia

Alexander I. Studenikin†

Department of Theoretical Physics, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia and Joint Institute for Nuclear Research, Dubna 141980, Moscow Region, Russia

(Received 11 February 2017; published 14 March 2017)

A thorough account of electromagnetic interactions of massive neutrinos in the theoretical formulation of low-energy elastic neutrino-electron scattering is given. The formalism of neutrino charge, magnetic, electric, and anapole form factors defined as matrices in the mass basis is employed under the assumption of three-neutrino mixing. The flavor change of neutrinos traveling from the source to the detector is taken into account and the role of the source-detector distance is inspected. The effects of neutrino flavor-transition millicharges and charge radii in the scattering experiments are pointed out.

DOI: 10.1103/PhysRevD.95.055013

... all experimental constraints on charge radius should be redone

Concluding remarks

Kouzakov, Studenikin
Phys. Rev. D 95 (2017) 055013

- cross section of ν-e is determined in terms of 3x3 matrices of ν electromagnetic form factors
- in short-baseline experiments one studies form factors in flavour basis
- long-baseline experiments more convenient to interpret in terms of fundamental form factors in mass basis
- ν millicharge when it is constrained in reactor short-baseline experiments (GEMMA, for instance) should be interpreted as

$$|e_{\nu_e}| = \sqrt{|(e_\nu)_{ee}|^2 + |(e_\nu)_{\mu e}|^2 + |(e_\nu)_{\tau e}|^2}$$

- ν charge radius in ν-e elastic scattering can't be considered as a shift $g_V \to g_V + \frac{2}{3} M_W^2 \langle r^2 \rangle \sin^2 \theta_W$, there are also contributions from flavor-transition charge radii

!

Generalized ν charge

Up to now we have used $\tilde{Q}_{jk} = \frac{2\sqrt{2}\pi\alpha}{G_F}\left[\frac{(e_\nu)_{jk}}{q^2} + \frac{1}{6}\langle r_\nu^2\rangle_{jk}\right]$ in mass basis

Finally we have in flavour basis

$$\tilde{Q}_{\ell'\ell} = \sum_{j,k=1}^{3} U_{\ell' j} U_{\ell k}^{*} \tilde{Q}_{jk} = \frac{2\sqrt{2}\pi\alpha}{G_F}\left[\frac{(e_\nu)_{\ell'\ell}}{q^2} + \frac{1}{6}\langle r_\nu^2\rangle_{\ell'\ell}\right]$$

where

$$(e_\nu)_{\ell'\ell} = \sum_{j,k=1}^{3} U_{\ell' j} U_{\ell k}^{*} (e_\nu)_{jk}$$

$$\langle r_\nu^2\rangle_{\ell'\ell} = \sum_{j,k=1}^{3} U_{\ell' j} U_{\ell k}^{*} \langle r_\nu^2\rangle_{jk}$$

millicharge — in ν flavour basis — charge radius

Kouzakov, Studenikin Phys. Rev. D 95 (2017) 055013

PHYSICAL REVIEW D **98**, 113010 (2018)

Editors' Suggestion

Ch – It – Ru
collaboration

Neutrino charge radii from COHERENT elastic neutrino-nucleus scattering

M. Cadeddu*
Dipartimento di Fisica, Università degli Studi di Cagliari, and INFN, Sezione di Cagliari, Complesso Universitario di Monserrato—S.P. per Sestu Km 0.700, 09042 Monserrato (Cagliari), Italy

C. Giunti†
Istituto Nazionale di Fisica Nucleare (INFN), Sezione di Torino, Via P. Giuria 1, I–10125 Torino, Italy

K. A. Kouzakov‡
Department of Nuclear Physics and Quantum Theory of Collisions, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia

Y. F. Li§
Institute of High Energy Physics, Chinese Academy of Sciences, Beijing 100049, China and School of Physical Sciences, University of Chinese Academy of Sciences, Beijing 100049, China

A. I. Studenikin‖
Department of Theoretical Physics, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia and Joint Institute for Nuclear Research, Dubna 141980, Moscow Region, Russia

Y. Y. Zhang¶
Institute of High Energy Physics, Chinese Academy of Sciences, Beijing 100049, China and School of Physical Sciences, University of Chinese Academy of Sciences, Beijing 100049, China

(Received 15 October 2018; published 26 December 2018)

Coherent elastic neutrino-nucleus scattering is a powerful probe of neutrino properties, in particular of the neutrino charge radii. We present the bounds on the neutrino charge radii obtained from the analysis of the data of the COHERENT experiment. We show that the time information of the COHERENT data allows us to restrict the allowed ranges of the neutrino charge radii, especially that of ν_μ. We also obtained for the first time bounds on the neutrino transition charge radii, which are quantities beyond the standard model.

DOI: 10.1103/PhysRevD.98.1130

Physical Review D
– Highlights 2018 –
Editors' Suggestion

"Using data from the COHERENT experiment, the authors put bounds on electromagnetic ν charge radii, including the first bounds on transition charge radii. These results show promising prospects for current and upcoming ν-nucleus experiments"

$$(|\langle r^2_{\nu_{e\mu}}\rangle|, |\langle r^2_{\nu_{e\tau}}\rangle|, |\langle r^2_{\nu_{\mu\tau}}\rangle|) < (22, 38, 27) \times 10^{-32}\ \mathrm{cm}^2$$

K. Kouzakov, A. Studenikin, "Electromagnetic properties of massive neutrinos in low-energy elastic neutrino-electron scattering"
Phys. Rev. D 95 (2017) 055013

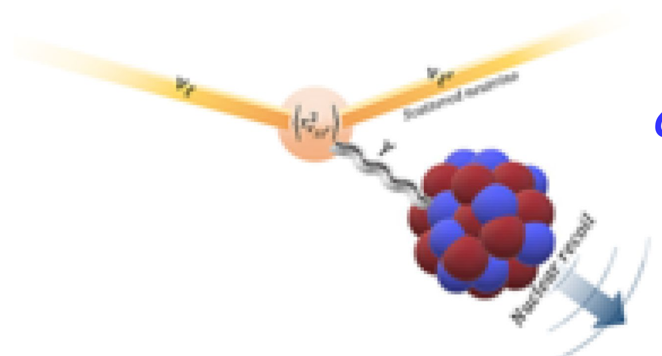
Physical Review D – Highlights 2018 – Editors' Suggestion

Editors' Suggestion

Neutrino charge radii from COHERENT elastic neutrino-nucleus scattering (/prd/abstract/10.1103/PhysRevD.98.113010)

M. Cadeddu, C. Giunti, K. A. Kouzakov, Y. F. Li, A. I. Studenikin, and Y. Y. Zhang

Phys. Rev. D **98**, 113010 (2018) – Published 26 December 2018



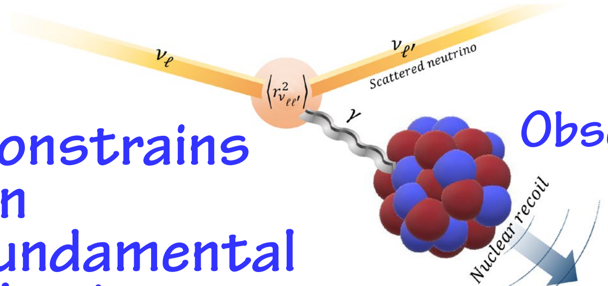
coherent ν scattering
due to charge radius

Using data from the COHERENT experiment, the authors put bounds on neutrino electromagnetic charge radii, including the first bounds on the transition charge radii. These results show promising prospects for current and upcoming neutrino-nucleus scattering experiments.

Show Abstract + ()

Particle Data Group,
Review of Particle Physics (2018-2022),
update of 2023

Coherent elastic ν - nucleous scattering (CEνNS)



constrains on fundamental physics

Predicted in 1974 (Freedman)

Observations: COHERNT (2017 – CsI detector,
2020 – Ar detector)
Dresden-II reactor
(2022 – Ge detector)

An updated review: Carlo Giunti (Neutrino 2022)

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Probing neutrino transition magnetic moments with coherent elastic neutrino-nucleus scattering

O.G. Miranda,[a] D.K. Papoulias,[b] M. Tórtola[b] and J.W.F. Valle[b]

[a] Departamento de Física, Centro de Investigación y de Estudios Avanzados del IPN, Apartado Postal 14-740 07000 Mexico, Distrito Federal, Mexico
[b] AHEP Group, Institut de Física Corpuscular — CSIC/Universitat de València, Parc Científic de Paterna, C/Catedrático José Beltrán 2, E-46980 Paterna, Valencia, Spain
E-mail: omr@fis.cinvestav.mx, dipapou@ific.uv.es, mariam@ific.uv.es, valle@ific.uv.es

ABSTRACT: We explore the potential of current and next generation of coherent elastic neutrino-nucleus scattering (CEνNS) experiments in probing neutrino electromagnetic interactions. On the basis of a thorough statistical analysis, we determine the sensitivities on each component of the Majorana neutrino transition magnetic moment (TMM), $|\Lambda_i|$, that follow from low-energy neutrino-nucleus experiments. We derive the sensitivity to neutrino TMM from the first CEνNS measurement by the COHERENT experiment, at the Spallation Neutron Source. We also present results for the next phases of COHERENT using HPGe, LAr and NaI[Tl] detectors and for reactor neutrino experiments such as CONUS, CONNIE, MINER, TEXONO and RED100. The role of the CP violating phases in each case is also briefly discussed. We conclude that future CEνNS experiments with low-threshold capabilities can improve current TMM limits obtained from Borexino data.

JHEP07(2019)103

- Neutrino, electroweak, and nuclear physics from COHERENT … with refined quenching factor, Cadeddu, Dordei, Giunti, Li, Zhang, PRD 2020

COHERENT data have been used for different purposes:

- nuclear neutron distributions
  Cadeddu, Giunti, Li, Zhang PRL 2018
- weak mixing angle
  Cadeddu & Dordei, PRD 2019
  Huang & Chen 2019
- ν electromagnetic properties
  Papoulias & Kosmas PRD 2018
- ν non-standard interactions
  Coloma, Gonzalez-Garcia, Maltoni, Schwetz PRD 2017
  Liao & Marfatia PLB 2017

Experimental limits on ν charge radius $\langle r_\nu^2 \rangle$

C. Giunti, A. Studenikin, "Electromagnetic interactions of neutrinos: a window to new physics", Rev. Mod. Phys. 87 (2015) 531

Method	Experiment	Limit (cm^2)	C.L.	Reference
Reactor $\bar{\nu}_e$-e^-	Krasnoyarsk	$\lvert\langle r^2_{\nu_e}\rangle\rvert < 7.3 \times 10^{-32}$	90%	Vidyakin *et al.* (1992)
	TEXONO	$-4.2 \times 10^{-32} < \langle r^2_{\nu_e}\rangle < 6.6 \times 10^{-32}$	90%	Deniz *et al.* (2010)[a]
Accelerator ν_e-e^-	LAMPF	$-7.12 \times 10^{-32} < \langle r^2_{\nu_e}\rangle < 10.88 \times 10^{-32}$	90%	Allen *et al.* (1993)[a]
	LSND	$-5.94 \times 10^{-32} < \langle r^2_{\nu_e}\rangle < 8.28 \times 10^{-32}$	90%	Auerbach *et al.* (2001)[a]
Accelerator ν_μ-e^-	BNL-E734	$-4.22 \times 10^{-32} < \langle r^2_{\nu_\mu}\rangle < 0.48 \times 10^{-32}$	90%	Ahrens *et al.* (1990)[a]
	CHARM-II	$\lvert\langle r^2_{\nu_\mu}\rangle\rvert < 1.2 \times 10^{-32}$	90%	Vilain *et al.* (1995)[a]

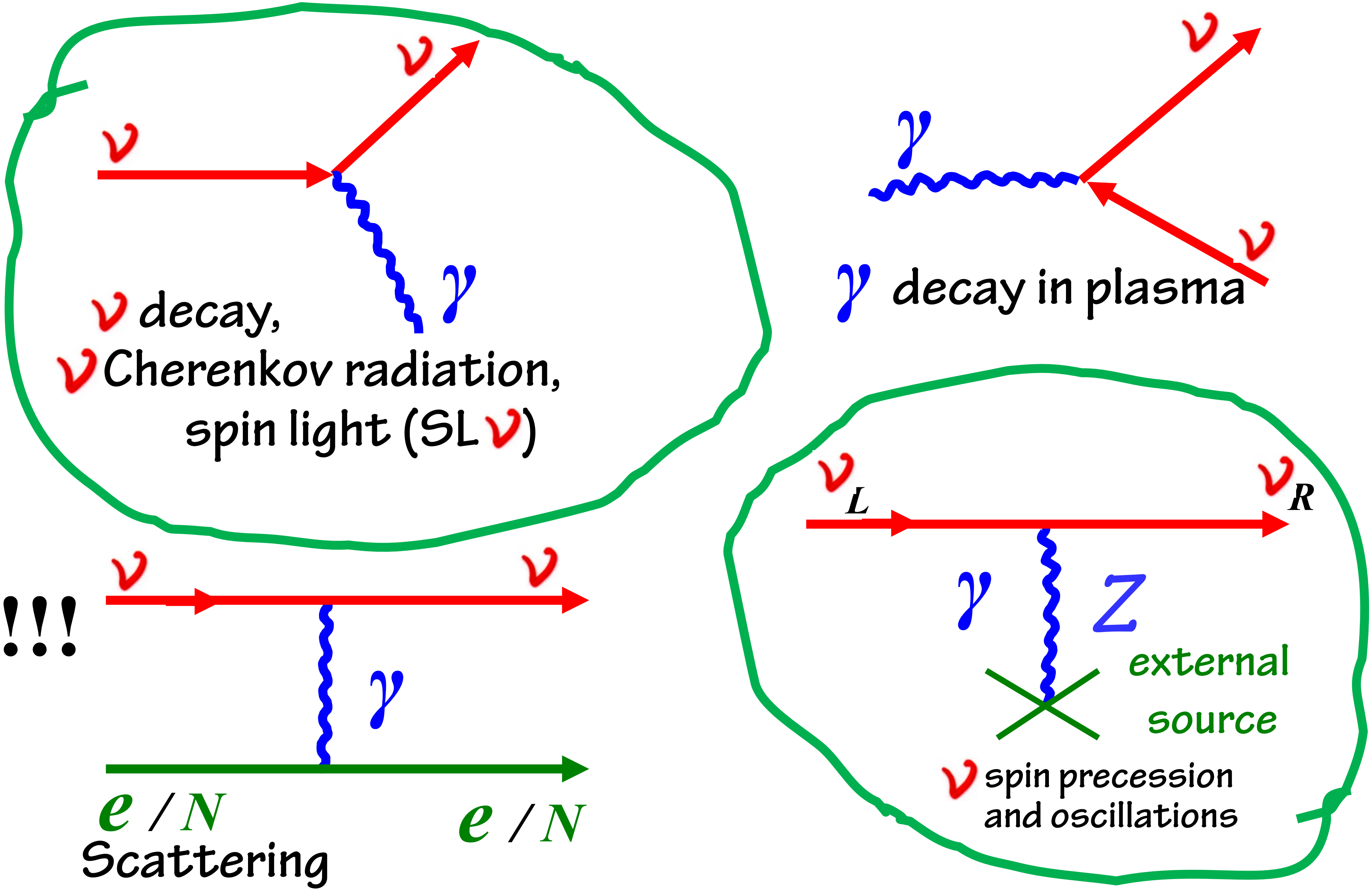
… updated by the recent constraints (effects of physics Beyond Standard Model)

$$(|\langle r^2_{\nu_{e\mu}}\rangle|, |\langle r^2_{\nu_{e\tau}}\rangle|, |\langle r^2_{\nu_{\mu\tau}}\rangle| < (28, 30, 35) \times 10^{-32} cm^2$$

M.Cadeddu, C. Giunti, K.Kouzakov, Yu-Feng Li, A. Studenikin, Y.Y.Zhang, Neutrino charge radii from COHERENT elastic neutrino-nucleus scattering, Phys.Rev.D 98 (2018) 113010

Electromagnetic ν in astrophycis and bounds on μ_ν and q_ν

ν electromagnetic interactions



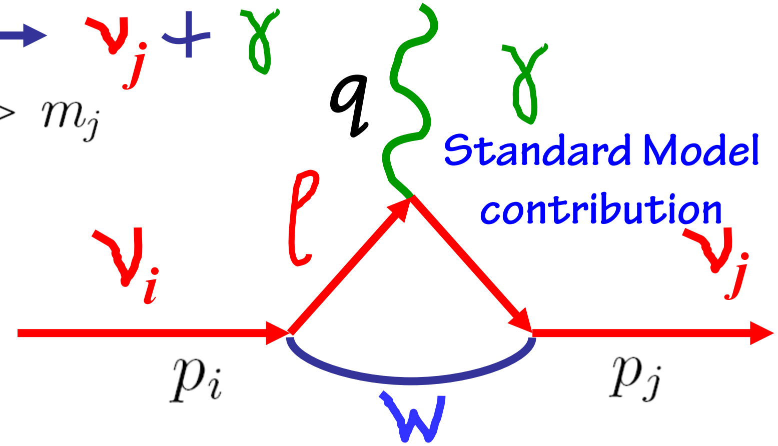
Neutrino radiative decay $\nu_i \longrightarrow \nu_j + \gamma$

$m_i > m_j$

$$\mathcal{L}_{int} = \frac{1}{2}\bar{\psi}_i \sigma_{\mu\nu}(\mu_{if} + i\gamma_5 d_{if})\psi_j F^{\mu\nu} + h.c.$$

$\Downarrow$

$$\Lambda_\mu^{if}(q) = -i\sigma_{\mu\nu}q^\nu(\mu_{if} + i\gamma_5 d_{if})$$



Standard Model contribution

Petkov 1977; Zatsepin, Smirnov 1978;
Bilenky, Petkov 1987; Pal, Wolfenstein 1982

Radiative decay rate

$$\Gamma_{\nu_i \to \nu_j + \gamma} = \frac{\mu_{eff}^2}{8\pi}\left(\frac{m_i^2 - m_j^2}{m_i^2}\right)^3 \approx 5\left(\frac{\mu_{eff}}{\mu_B}\right)^2\left(\frac{m_i^2 - m_j^2}{m_i^2}\right)^3\left(\frac{m_i}{1\ eV}\right)^3 s^{-1}$$

$$\mu_{eff}^2 = |\mu_{ij}|^2 + |\epsilon_{ij}|^2$$

transition magnetic and electric moments
(for Dirac and Majorana ν)

$$\tau^{\rm rf}_{\nu_i \to \nu_f + \gamma} \simeq 0.19\left(\frac{m_i^2}{m_i^2 - m_f^2}\right)^3\left(\frac{\rm eV}{m_i}\right)^3\left(\frac{\mu_{\rm B}}{\mu_{fi}^{\rm eff}}\right)^2 \ {\rm s}$$

$\Downarrow$

- ν life time is indeed huge ...

Radiative decay has been constrained from absence of decay photons:

1) reactor $\bar{\nu}_e$ and solar ν_e fluxes,
2) SN 1987A ν burst (all flavours)
3) spectral distortion of CMBR

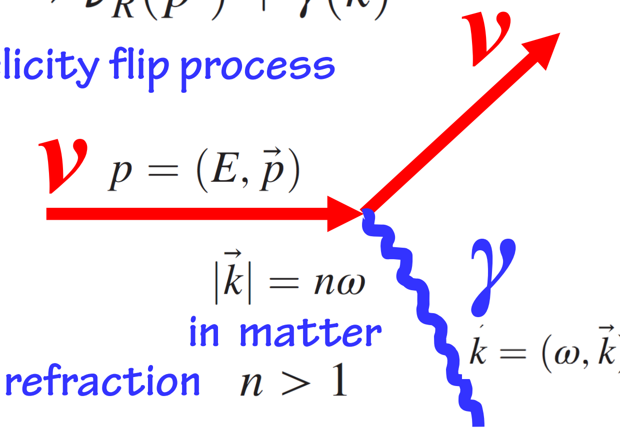
Raffelt 1999
Kolb, Turner 1990
Ressell, Turner 1990

Neutrino Cherenkov radiation $\nu_L(p) \to \nu_R(p') + \gamma(k)$

helicity flip process

ν transition amplitude due to μ_ν

$$M = \frac{\mu}{n}\overline{u^{(+)}}(p')\sigma_{\mu\nu}k^\mu u^{(-)}(p)\varepsilon^\nu(k,\lambda)$$



index of refraction $n > 1$

Cherenkov radiation

Cherenkov process rate

$$\Gamma = \frac{1}{2(2\pi)^2 E}\int \frac{d^3p'}{2E'}\frac{d^3k}{2\omega}|M|^2\delta^4(p - p' - k)$$

after integration

$$\Gamma = \frac{1}{16\pi E^2 v}\int n^2 d\omega d(\cos\theta)|M|^2 \delta\left(\cos\theta - \frac{2\omega E + (n^2-1)\omega^2}{2n\omega E v}\right)$$

$$v = |\vec{p}|/E$$

photon emission angle

$$\cos\theta = \frac{1}{nv}\left(1 + (n^2-1)\frac{\omega}{2E}\right)$$

$$|\cos\theta| \le 1$$

$$\Gamma = \frac{\mu^2}{4\pi E^2 v}\int_{\omega_{\min}}^{\omega_{\max}}\left\{\left[\frac{(n^2-1)^2}{n^2}E^2 + (n^2-1)m_\nu^2\right]\omega^2 - \frac{(n^2-1)^2}{n^2}E\omega^3 - \frac{(n^2-1)^3}{4n^2}\omega^4\right\}d\omega$$

Solar νs with $\mu_\nu \sim 3\times 10^{-11}\mu_B$ emit 5 γ per day in 1 Km^3 water detector

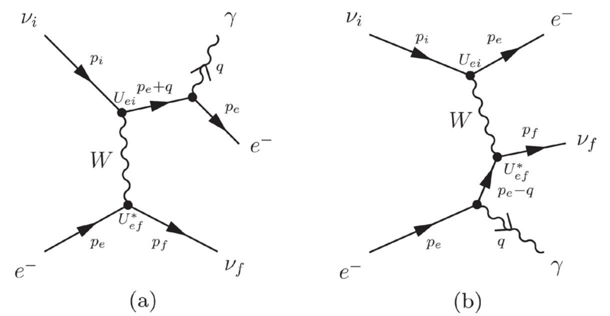
Grimus & Neufeld, 1993

ν radiative decay and Cherenkov radiation in external environments

coherent forward elastic scattering on (electron) background also generates $\nu_i \longrightarrow \nu_j + \gamma$

not suppressed by GIM

D'Olive, Nieves, Pal (1990)



Giunti et al (1992)

Cherenkov radiation by ν in magnetic field

Galtsov, Nikitina (1972)
Ioannisian & Raffelt (1997)

B induces effective ν-γ vertex and modifies γ dispersion relation

(no need for BSM)

ν in medium acquire induce q as a consequence of weak interactions

Oraevsky, Semikoz, Smorodinsky (1986)

another mechanism of Cherenkov radiation in medium

Sawyer (1992)
D'Olive, Nieves, Pal (1996)

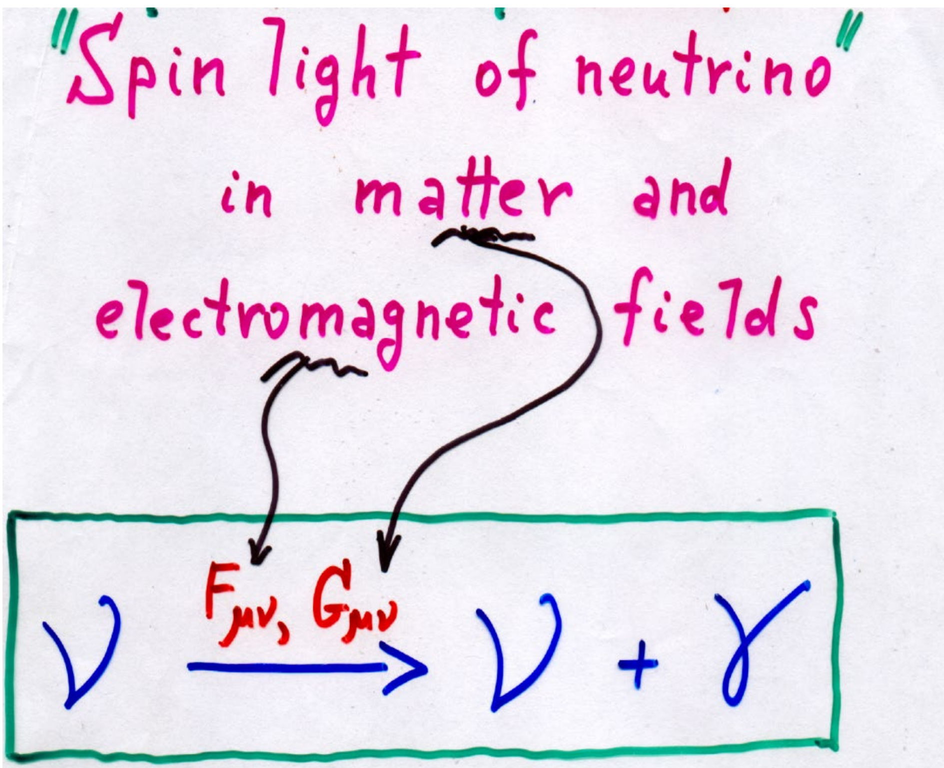
effect for $m_\nu = 0$ in SM (without physics BSM)

- other particular cases for $\nu_i \longrightarrow \nu_j + \gamma$ in em fields and matter

Skobelev (1976)
Borisov, Zhukosky, Ternov (1988)
Ternov (2016)

• New mechanism of electromagnetic radiation

$SL\nu$



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- Phys.Lett. B 515 (2001) 94
Phys.Lett. B 564 (2003) 27
Phys.Lett. B 601 (2004) 171

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J.Phys.A: Math.Gen. 39 (2006) 6769
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Phys.Lett.B 718 (2012) 512

- … quasi-classical approach to ν spin evolution in an external electromagnetic fields

Journal of Cosmology and Astroparticle Physics
An IOP and SISSA journal

Spin light of neutrino in astrophysical environments

Alexander Grigoriev,[b,c] **Alexey Lokhov,**[d] **Alexander Studenikin**[a,e,1] **and Alexei Ternov**[c]

[a]Department of Theoretical Physics, Moscow State University,
119992 Moscow, Russia
[b]Skobeltsyn Institute of Nuclear Physics, Moscow State University,
119992 Moscow, Russia
[c]Department of Theoretical Physics, Moscow Institute of Physics and Technology,
141701 Dolgoprudny, Russia
[d]Institute for Nuclear Research, Russian Academy of Sciences,
117312 Moscow, Russia
[e]Dzhelepov Laboratory of Nuclear Problems, Joint Institute for Nuclear Research,
141980 Dubna, Russia

E-mail: ax.grigoriev@mail.ru, lokhov.alex@gmail.com, studenik@srd.sinp.msu.ru,
ternov.ai@mipt.ru

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#4

New mechanism of

e.m. radiation by ν in matter and e.m. fields, and gravitational fields

$$\nu \xrightarrow[\text{environment}]{\text{background}} \nu + \gamma,$$

|| "Spin Light of Neutrino": "$SL\nu$"

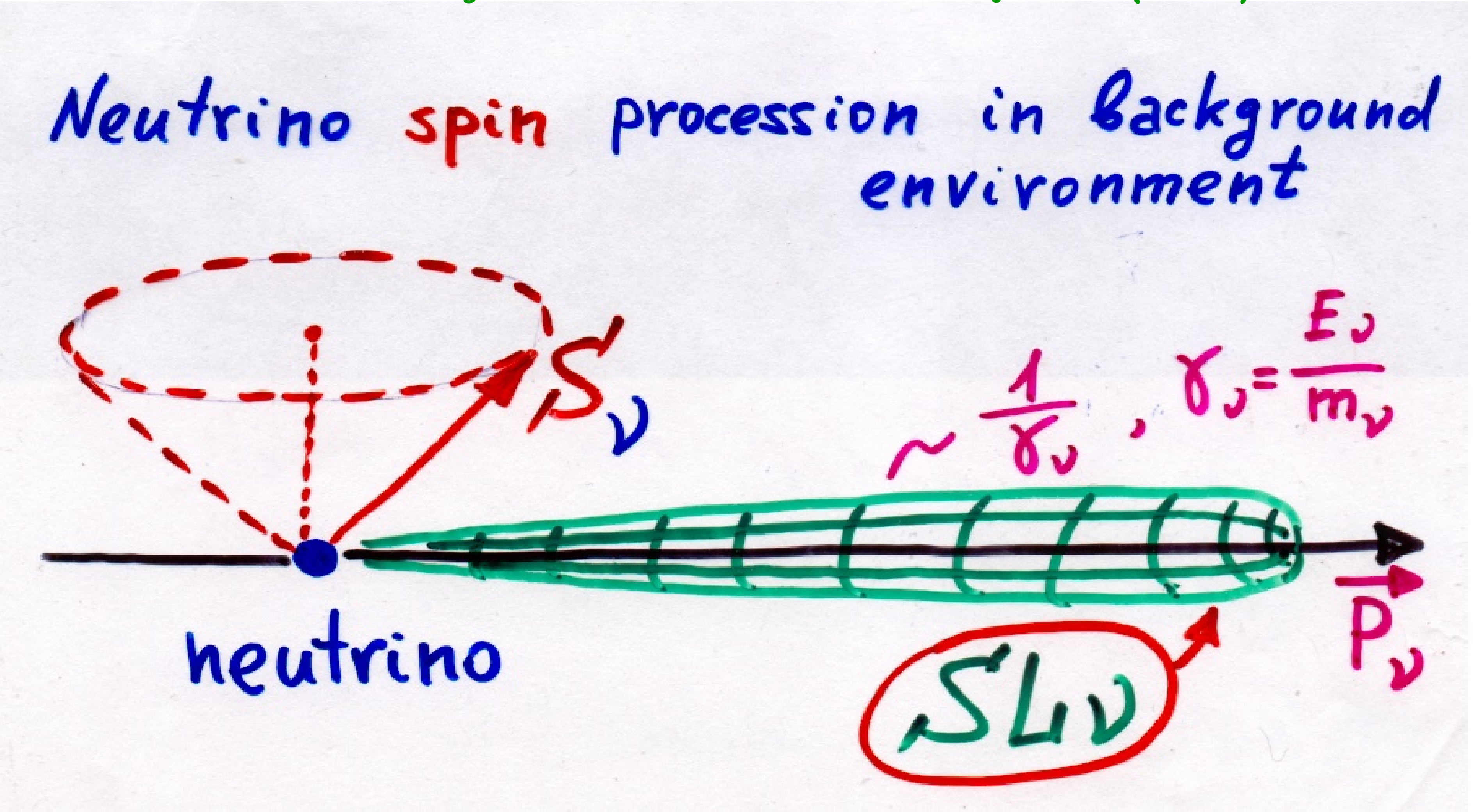
A.Lobanov, A.Studenikin, Phys.Lett.B 564 (2003) 27

… quasi-classical approach to ν spin evolution in an external electromagnetic fields

Quasi-classical theory of spin light of neutrino in matter and gravitational field

$SL\nu$

A.Lobanov, A.Studenikin, Phys.Lett. B 564 (2003) 27,
Phys.Lett. B 601 (2004) 171;
M.Dvornikov, A.Grigoriev, A.Studenikin, Int.J.Mod.Phys. D 14 (2005) 309



New mechanism of electromagnetic radiation

? Why Spin Light of neutrino $SL\nu$ / of electron SLe in matter

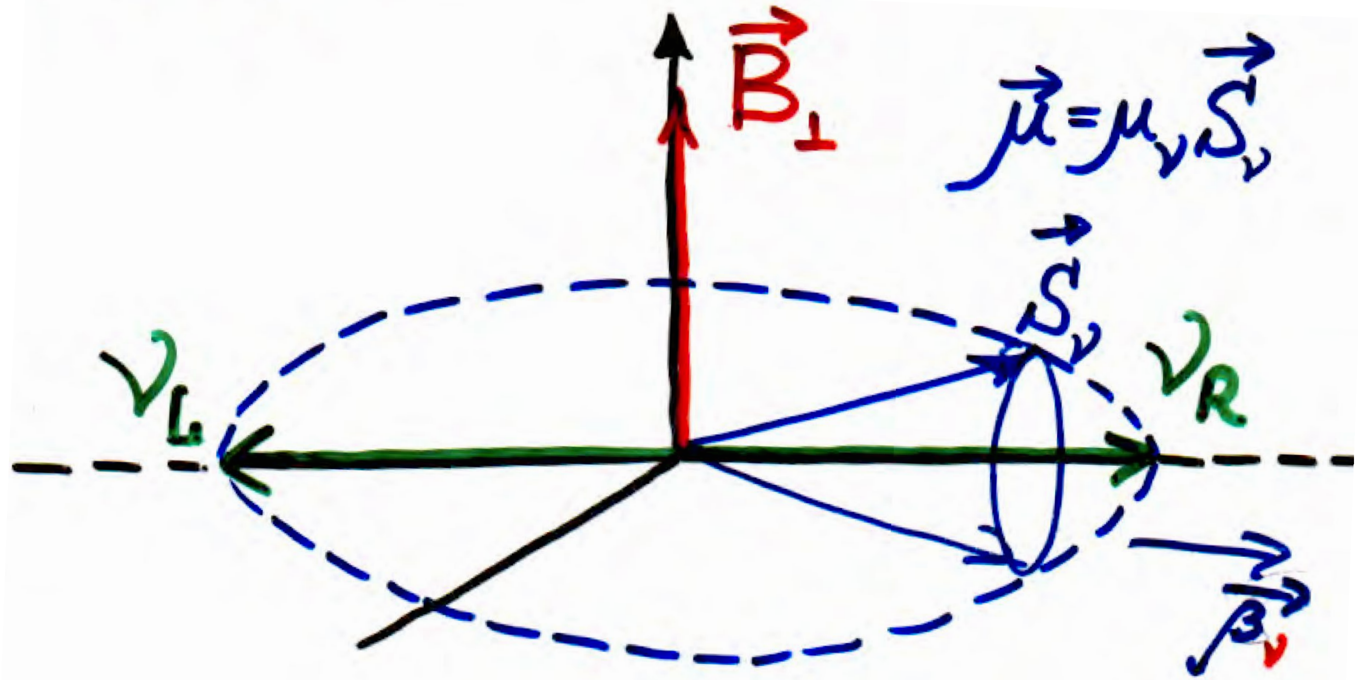
Analogies with:

* classical electrodynamics

an object with charge $Q = 0$ and magnetic moment $\vec{m} = \frac{1}{2}\sum_i e_i [\vec{r}_i \times \vec{v}_i] \neq 0$

$$I^{cl.el.} = \frac{2}{3}\ddot{\vec{m}}^2$$

← magnetic dipole radiation power



$$\frac{d\vec{S}_\nu}{dt} = 2\mu_\nu [\vec{S}_\nu \times \vec{B}] + 2\mu_\nu [\vec{S}_\nu \times \vec{G}]$$

electromagnetic interacton with e.m. field

weak interaction with matter

ν spin evolution in presence of general external fields

M.Dvornikov, A.Studenikin,
JHEP 09 (2002) 016

General types non-derivative interaction with external fields

$$-\mathcal{L} = g_s s(x)\bar{\nu}\nu + g_p \pi(x)\bar{\nu}\gamma^5\nu + g_v V^\mu(x)\bar{\nu}\gamma_\mu\nu + g_a A^\mu(x)\bar{\nu}\gamma_\mu\gamma^5\nu + \\ + \frac{g_t}{2} T^{\mu\nu}\bar{\nu}\sigma_{\mu\nu}\nu + \frac{g_t'}{2}\Pi^{\mu\nu}\bar{\nu}\sigma_{\mu\nu}\gamma_5\nu,$$

scalar, pseudoscalar, vector, axial-vector, tensor and pseudotensor fields: $s, \pi, V^\mu = (V^0, \vec{V}), A^\mu = (A^0, \vec{A}),$ $T_{\mu\nu} = (\vec{a}, \vec{b}), \Pi_{\mu\nu} = (\vec{c}, \vec{d})$

Relativistic equation (quasiclassical) for ν spin vector:

$$\dot{\vec{\zeta}}_\nu = 2g_a \left\{ A^0[\vec{\zeta}_\nu \times \vec{\beta}] - \frac{m_\nu}{E_\nu}[\vec{\zeta}_\nu \times \vec{A}] - \frac{E_\nu}{E_\nu + m_\nu}(\vec{A}\vec{\beta})[\vec{\zeta}_\nu \times \vec{\beta}] \right\} \\ +2g_t \left\{ [\vec{\zeta}_\nu \times \vec{b}] - \frac{E_\nu}{E_\nu + m_\nu}(\vec{\beta}\vec{b})[\vec{\zeta}_\nu \times \vec{\beta}] + [\vec{\zeta}_\nu \times [\vec{a} \times \vec{\beta}]] \right\} + \\ +2ig_t' \left\{ [\vec{\zeta}_\nu \times \vec{c}] - \frac{E_\nu}{E_\nu + m_\nu}(\vec{\beta}\vec{c})[\vec{\zeta}_\nu \times \vec{\beta}] - [\vec{\zeta}_\nu \times [\vec{d} \times \vec{\beta}]] \right\}.$$

- Neither S nor π nor V contributes to spin evolution

- Electromagnetic interaction

$$T_{\mu\nu} = F_{\mu\nu} = (\vec{E}, \vec{B})$$

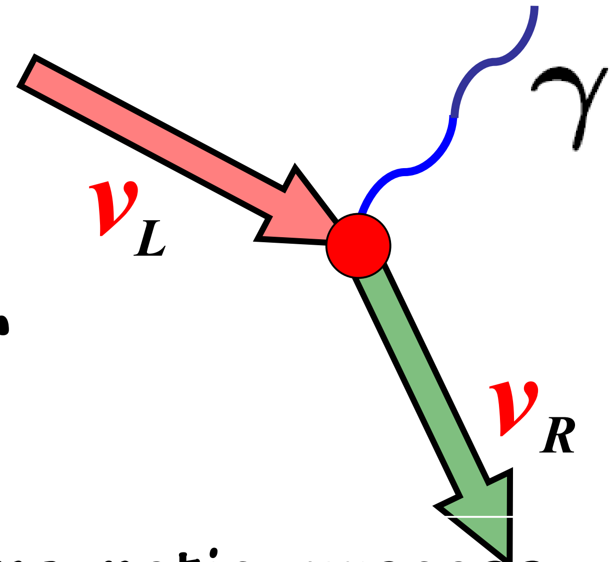
- SM weak interaction

$$G_{\mu\nu} = (-\vec{P}, \vec{M})$$

$$\vec{M} = \gamma(A^0\vec{\beta} - \vec{A})$$
$$\vec{P} = -\gamma[\vec{\beta} \times \vec{A}],$$

… quantum theory of $SL\nu$

Spin light of neutrino in matter



new mechanism of the electromagnetic process stimulated by the presence of matter, in which neutrino with nonzero magnetic moment emits light

A.Lobanov, A.Studenikin, Phys.Lett. B 564 (2003) 27,
Phys.Lett. B 601 (2004) 171

A.S., A.Ternov, Phys.Lett. B 608 (2005) 107

A.Grigoriev, A.S., A.Ternov, Phys.Lett. B 622 (2005) 199

A.S., J.Phys.A: Math.Theor. 41 (2008) 16402

A.S., J.Phys.A: Math.Gen. 39 (2006) 6769

A.Grigoriev, A.Lokhov, A.Ternov, A.Studenikin,
Phys. Lett. B 718 (2012) 512
JCAP 11 (2017) 024

«method of exact solutions»

Interaction of particles in external electromagnetic fields

(Furry representation in quantum electrodynamics)

Potential of electromagnetic field

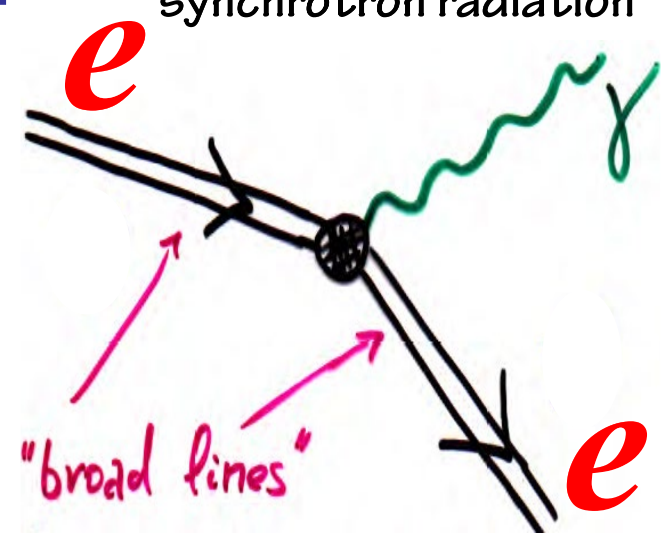
$$A_\mu(x) = A^q_\mu(x) + A^{ext}_\mu(x)$$

quantized part of potential

$$\mathbf{B}_\perp$$

$$e \longrightarrow e + \gamma$$

synchrotron radiation



evolution operator

$$U_F(t_1,t_2) = Texp\Big[-i\int_{t_1}^{t_2} j^\mu(x)A^q_\mu(x)dx\Big]$$

charged particles **current**

$$j_\mu(x) = \frac{e}{2}\left[\overline{\Psi}_F\gamma_\mu, \Psi_F\right]$$

Dirac equation in **external classical (nonquantized) field** $A^{ext}_\mu(x)$

$$\Big\{\gamma^\mu\Big(i\partial_\mu - eA^{ext}_\mu(x)\Big) - m_e\Big\}\Psi_F(x) = 0$$

- ...beyond perturbation series expansion, strong fields and non linear effects...

ν in external electromagnetic fields → ν in dense matter

…«method of exact solutions »…

Studenikin

- Quantum treatment of neutrino in background matter
  J. Phys. A: Math. Gen. 39 (2006) 6769-6776
- Method of wave equations exact solutions in studies of neutrinos and electron interactions in dense matter
  J.Phys.A: Math.Theor. 41 (2008) 164047 (20 p)
- Neutrinos and electrons in background matter: A new approach
  Ann.Fond. de Broglie 31 (2006) 289-316

- ν quantum states in dense magnetized matter

… new effect of …

Spin Light of ν in matter

$SL\nu$

$\Rightarrow$

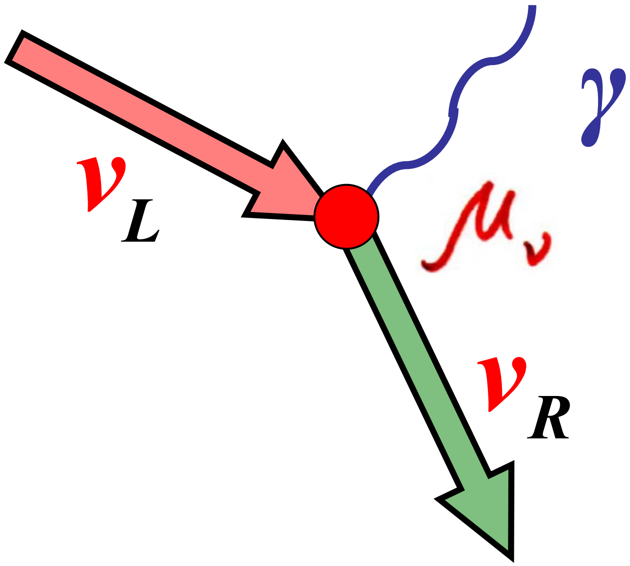
ν energy quantization in rotating matter

… pheno-meno-logical consequences in astrophysics (pulsars)

ν in matter treated within «method of exact solutions»

(Dirac equation with matter potential for ν)

Neutrino – photon coupling



broad neutrino lines
account for interaction
with environment

"Spin light of neutrino in matter"

$SL\nu$

- … within the quantum treatment based on method of exact solutions …

Modified Dirac equation for neutrino in matter

Addition to the vacuum neutrino Lagrangian

$$\Delta L_{eff} = \Delta L_{eff}^{CC} + \Delta L_{eff}^{NC} = -f^{\mu}\left(\bar{\nu}\gamma_{\mu}\frac{1+\gamma^5}{2}\nu\right)$$

matter current

where $$f^{\mu} = \frac{G_F}{\sqrt{2}}\Big((1+4\sin^2\theta_W)j^{\mu} - \lambda^{\mu}\Big)$$

matter polarization

$$\left\{i\gamma_{\mu}\partial^{\mu} - \frac{1}{2}\gamma_{\mu}(1+\gamma_5)f^{\mu} - m\right\}\Psi(x) = 0$$

A.Studenikin, A.Ternov, hep-ph/0410297; Phys.Lett.B 608 (2005) 107

It is supposed that there is a macroscopic amount of electrons in the scale of a neutrino de Broglie wave length. Therefore, **the interaction of a neutrino with the matter (electrons) is coherent.**

L.Chang, R.Zia,'88; J.Panteleone,'91; K.Kiers, N.Weiss, M.Tytgat,'97-'98; P.Manheim,'88; D.Nötzold, G.Raffelt,'88; J.Nieves,'89; V.Oraevsky, V.Semikoz, Ya.Smorodinsky,89; W.Naxton, W-M.Zhang'91; M.Kachelriess,'98; A.Kusenko, M.Postma,'02.

This is the most general equation of motion of a neutrino in which the effective potential accounts for both the **charged** and **neutral-current** interactions with the background matter and also for the possible effects of the matter **motion** and **polarization**.

Quantum theory of spin light of neutrino

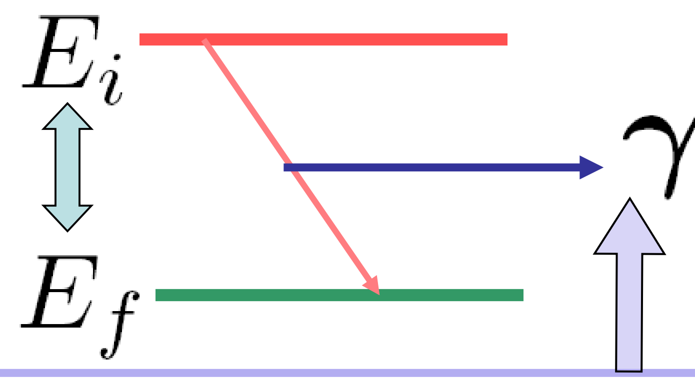
$SL\nu$

Quantum treatment of *spin light of neutrino* in matter
showns that this process originates from the two subdivided phenomena:

the shift of the neutrino energy levels in the presence of the background matter, which is different for the two opposite **neutrino helicity states**,

$$E = \sqrt{\mathbf{p}^2\left(1 - s\alpha\frac{m}{p}\right)^2 + m^2} + \alpha m$$

$$s = \pm 1$$



the radiation of the photon in the process of the neutrino transition from the **"excited" helicity state** to the **low-lying helicity** state in matter

A.Studenikin, A.Ternov, Phys.Lett.B 608 (2005) 107;

A.Grigoriev, A.Studenikin, A.Ternov, Phys.Lett.B 622 (2005) 199;
Grav. & Cosm. 14 (2005) 132;

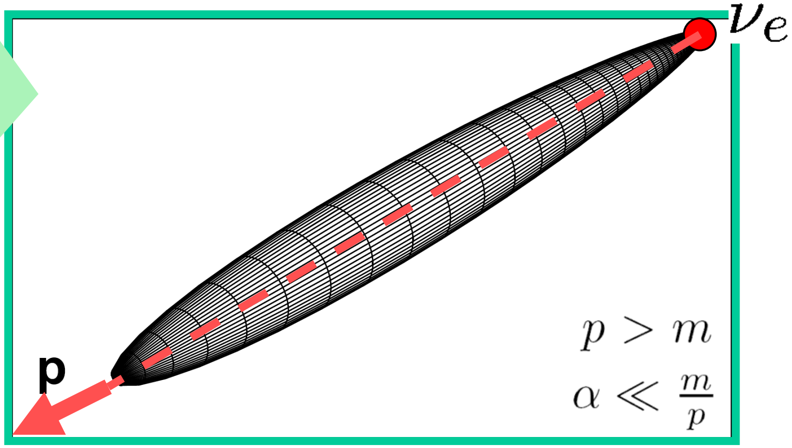
neutrino-spin self-polarization effect in the matter

A.Lobanov, A.Studenikin, Phys.Lett.B 564 (2003) 27;
Phys.Lett.B 601 (2004) 171

Spatial distribution of radiation power

From the angular distribution of $SL\nu$

$$I = \mu^2 \int_0^{\pi} \omega^4 \left[(\tilde\beta\tilde\beta' + 1)(1 - y\cos\theta) - (\tilde\beta + \tilde\beta')(\cos\theta - y)\right] \frac{\sin\theta}{1 + \tilde\beta' y} d\theta$$



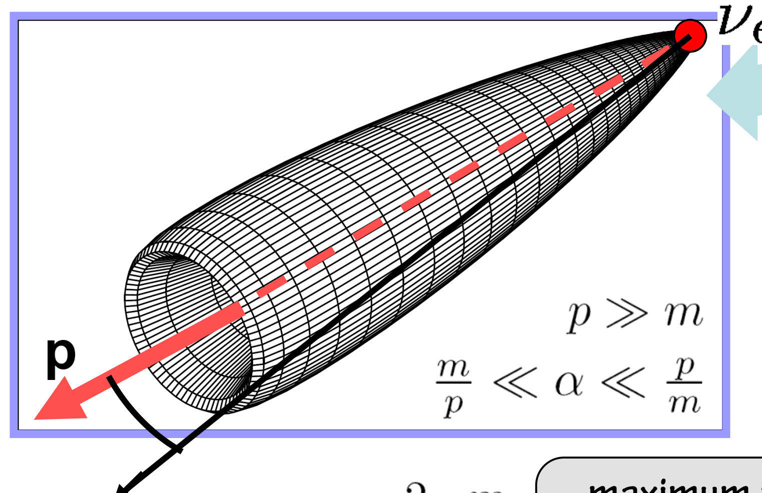
for $p/m = 5$ and $\alpha = 0.01$

$n \approx 10^{35} cm^{-3}$

neutrino momentum

mass

matter density



$$\cos\theta_{max} \simeq 1 - \frac{2}{3}\alpha\frac{m}{p}$$

maximum in radiation power distribution

for $p/m = 10^3$ and $\alpha = 100$

$n \approx 10^{39} cm^{-3}$

increase of matter density

projector-like distribution → cap-like distribution

It is possible to have $\tau = \frac{1}{\Gamma_{SL\nu}} << $ age of the Universe ?

For ultra-relativistic ν

with momentum $p \sim 10^{20} eV$

and magnetic moment $\mu \sim 10^{-10} \mu_B$

in very dense matter $n \sim 10^{40} cm^{-3}$

from

$$\Gamma_{SL\nu} = 4\mu^2 \alpha^2 m_\nu^2 p$$

$$p \gg m_{plasmon}$$

also discussed by A.Kuznetsov, N.Mikheev, IJMP A 2007

$$\alpha m_\nu = \frac{1}{2\sqrt{2}} G_F n (1 + sin^2\theta_W)$$

A.Lobanov, A.S., PLB 2003; PLB 2004
A.Grigoriev, A.S., PLB 2005
A.Grigoriev, A.S., A.Ternov, PLB 2005
A.Grigories, A.Lokhov, A.S., A.Ternov, PLB 2012

it follows that

$$\tau_{SL\nu} = \frac{1}{\Gamma} = 1.5 \times 10^{-8} s$$

A.Grigoriev, A.Lokhov, A.Ternov, A.Studenikin

The effect of plasmon mass on Spin Light of Neutrino in dense matter

Phys.Lett. B 718 (2012) 512-515

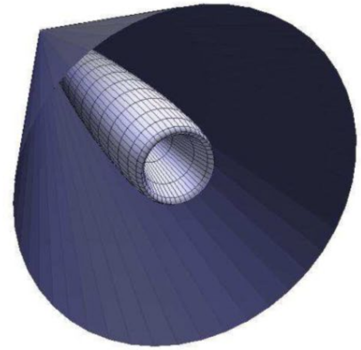


Figure 1: 3D representation of the radiation power distribution.

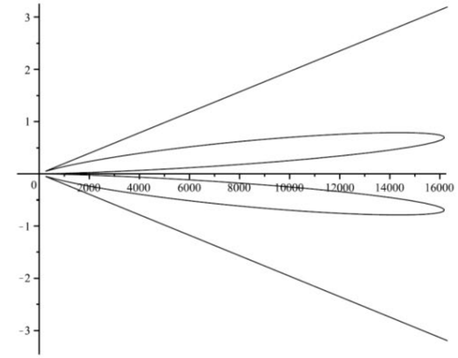


Figure 2: The two-dimensional cut along the symmetry axis. Relative units are used.

4. Conclusions

We developed a detailed evaluation of the spin light of neutrino in matter accounting for effects of the emitted plasmon mass. On the base of the exact solution of the modified Dirac equation for the neutrino wave function in the presence of the background matter the appearance of the threshold for the considered process is confirmed. The obtained exact and explicit threshold condition relation exhibit a rather complicated dependance on the matter density and neutrino mass. The dependance of the rate and power on the neutrino energy, matter density and the angular distribution of the $SL\nu$ is investigated in details. It is shown how the rate and power wash out when the threshold parameter $a = m_\gamma^2/4\tilde{n}p$ approaching unity. From the performed detailed analysis it is shown that the $SL\nu$ mechanism is practically insensitive to the emitted plasmon mass for very high densities of matter (even up to $n = 10^{41} cm^{-3}$) for ultra-high energy neutrinos for a wide range of energies starting from $E = 1$ TeV. This conclusion is of interest for astrophysical applications of $SL\nu$ radiation mechanism in light of the recently reported hints of $1 \div 10$ PeV neutrinos observed by IceCube [17].

Journal of Cosmology and Astroparticle Physics
An IOP and SISSA journal

Spin light of neutrino in astrophysical environments

Alexander Grigoriev,[b,c] Alexey Lokhov,[d] Alexander Studenikin[a,e,1] and Alexei Ternov[c]

[a]Department of Theoretical Physics, Moscow State University, 119992 Moscow, Russia
[b]Skobeltsyn Institute of Nuclear Physics, Moscow State University, 119992 Moscow, Russia
[c]Department of Theoretical Physics, Moscow Institute of Physics and Technology, 141701 Dolgoprudny, Russia
[d]Institute for Nuclear Research, Russian Academy of Sciences, 117312 Moscow, Russia
[e]Dzhelepov Laboratory of Nuclear Problems, Joint Institute for Nuclear Research, 141980 Dubna, Russia

E-mail: ax.grigoriev@mail.ru, lokhov.alex@gmail.com, studenik@srd.sinp.msu.ru, ternov.ai@mipt.ru

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JCAP11(2017)024

A.Grigoriev, A.Lokhov, A.Studenikin, A.Ternov, Spin light of neutrino in astrophysical environments, J. Cosm. Astropart. Phys. 11 (2017) 024

$SL\nu$ in neutron matter of real astrophysical objects [4]

❑ **Plasma effects [5]**

• Photon dispersion with plasmon mass in the degenerate electron gas:

$$\omega = \sqrt{\mathbf{k}^2 + m_\gamma^2}$$

$$m_\gamma = \left(\frac{2\alpha}{\pi}\right)^{1/2} \mu_e \simeq 8.87 \times \left(\frac{n_e}{10^{37}\,\mathrm{cm}^{-3}}\right)^{1/3} \mathrm{MeV}$$

• Threshold condition for the $SL\nu$ [10]: ($Y_e = n_e/n_n$)

$$\frac{m_\gamma^2 + 2\, m_\gamma m_\nu}{4\tilde{n}p} < 1$$

• Neutron matter: (antineutrinos act)

$$\tilde{n} = \frac{1}{2\sqrt{2}} G_F n_n \simeq 3.2 \times \left(\frac{n_n}{10^{38}\,\mathrm{cm}^{-3}}\right) \mathrm{eV},$$

$$E > p_{\mathrm{th}} \simeq 28.5 \times \frac{Y_e^{2/3}}{1 - Y_e} \left(\frac{10^{38}\,\mathrm{cm}^{-3}}{n_n}\right)^{1/3} \mathrm{TeV} \quad \Rightarrow \quad E_{\mathrm{th}} \simeq 6.82\ \mathrm{TeV}.$$

$n_n = 10^{38}\,\mathrm{cm}^{-3}$, $Y_e = 0.1$

• Mean photon energy near the threshold: $\langle\omega\rangle = I/\Gamma \simeq p \simeq E_\nu$.

For most favorable conditions as low density of the charged matter component is needed as possible

❑ **W boson production** $\bar{\nu}_e + e^- \to W^-$ [4]

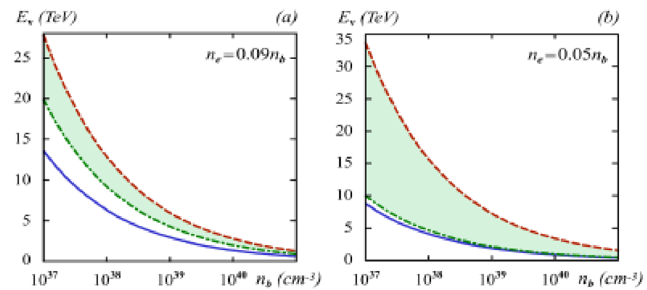


Figure 2. The allowed range of electron antineutrino energies for the $SL\nu$ in the matter of a neutron star depending on the neutron density. Solid line: the $SL\nu$ process threshold without account for the $\bar{\nu}_e e$-scattering; dash-dotted line: the $SL\nu$ process threshold with account for the $\bar{\nu}_e e$-scattering; dashed line: the threshold for the W boson production. (a) $Y_e = 0.09$; (b) $Y_e = 0.05$. The allowed regions are marked in green.

W-boson threshold energy $\varepsilon_W = \frac{m_W^2}{4\mu_e} \simeq 5.77 \times \left(\frac{10^{38}\,\mathrm{cm}^{-3}}{Y_e n_n}\right)^{1/3} \mathrm{TeV}$

• Electron antineutrinos: s-channel interaction with matter through W-boson, importance of the propagator effects $\Rightarrow$ correction to the effective potential of neutrino motion → antineutrino energy shift up → $SL\nu$ is suppressed at Y_e=0.1, but allowed already for Y_e=0.09

• μ and τ antineutrinos: only t-channel interaction with matter through Z-boson, no propagator effects $\Rightarrow$ the $SL\nu$ is allowed if neutrino energy is greater than the W-boson threshold ε_W

Neutrino lifetime with respect to the $SL\nu$ for most optimistic set of parameters:

$$\tau_{SL\nu} = 10^{-4} - 10^{3}\, s, \quad \text{for} \quad n_b = 10^{41} - 10^{38}\,\mathrm{cm}^{-3}$$

The $SL\nu$ in short Gamma-Ray Bursts (SGRBs)

Factors for best $SL\nu$ generation efficiency

• High neutrino energy and density

• High background neutral matter density

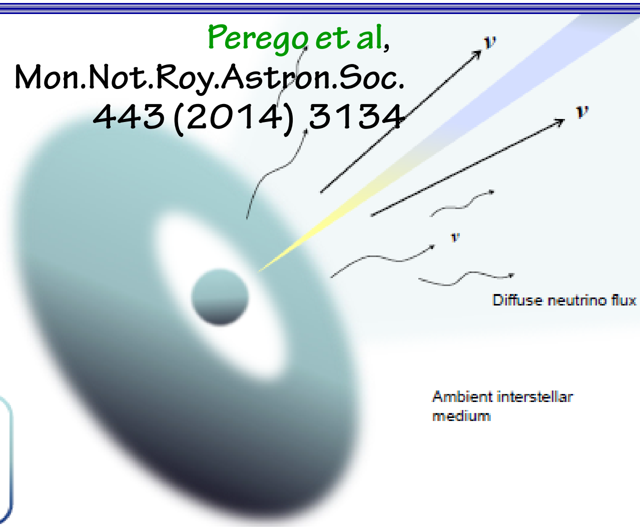
• Low density of the matter charged component

• Low temperature of the charged component

• Considerable extension of the medium

$SL\nu$ radiation by ultra high-energy neutrino in the diffuse neutrino wind blown during neutron stars merger

Perego et al, Mon.Not.Roy.Astron.Soc. 443 (2014) 3134



Matter characteristics[6]:

• neutrinos

$n_\nu \sim 10^{32}\ \mathrm{cm}^{-3}$

• electrons

$Y_e = 0.01$

$T = 0.1\ \mathrm{MeV}$

$\rho = 5 \times 10^3\ \mathrm{g/cm}^3$

$\Rightarrow$ $n_e \simeq 3 \times 10^{25}\ \mathrm{cm}^{-3}$, $m_\gamma \simeq 10^{-3}\ \mathrm{MeV}$, $E_{\mathrm{th}} \simeq 1\ \mathrm{GeV}$

Radiation time

$$\tau_{SL\nu} \simeq 5.4 \times 10^{15} \left(\frac{10^{-11}\mu_B}{\mu}\right)^2 \left(\frac{10^{32}\,\mathrm{cm}^{-3}}{n_{\nu_e}}\right)^2 \left(\frac{1\,\mathrm{PeV}}{E_\nu}\right) \mathrm{s}$$

Neutrino parameters: $\mu \simeq 2.9 \times 10^{-11}\mu_B$, $E_\nu \sim 10^{12} - 10^{18}\ \mathrm{eV}$

$$\tau_{SL\nu} \simeq 6.4 \times (10^{11} - 10^{17})\ \mathrm{s} = 2 \times (10^4 - 10^{10})\ \mathrm{years}$$

Eur. Phys. J. C (2022) 82:287
https://doi.org/10.1140/epjc/s10052-022-10244-w

THE EUROPEAN PHYSICAL JOURNAL C

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Letter

Neutrino spin operator and dispersion in moving matter

Alexander Grigoriev[1,a], Alexander Studenikin[2,3,4,b], Alexei Ternov[1,c]

[1] Department of Theoretical Physics, Moscow Institute for Physics and Technology, 141700 Dolgoprudny, Russia
[2] Department of Theoretical Physics, Moscow State University, 119992 Moscow, Russia
[3] Dzhelepov Laboratory of Nuclear Problems, Joint Institute for Nuclear Research, 141980 Dubna, Russia
[4] Interdisciplinary Scientific and Educational School of Moscow University "Fundamental and Applied Space Research", Moscow State University,

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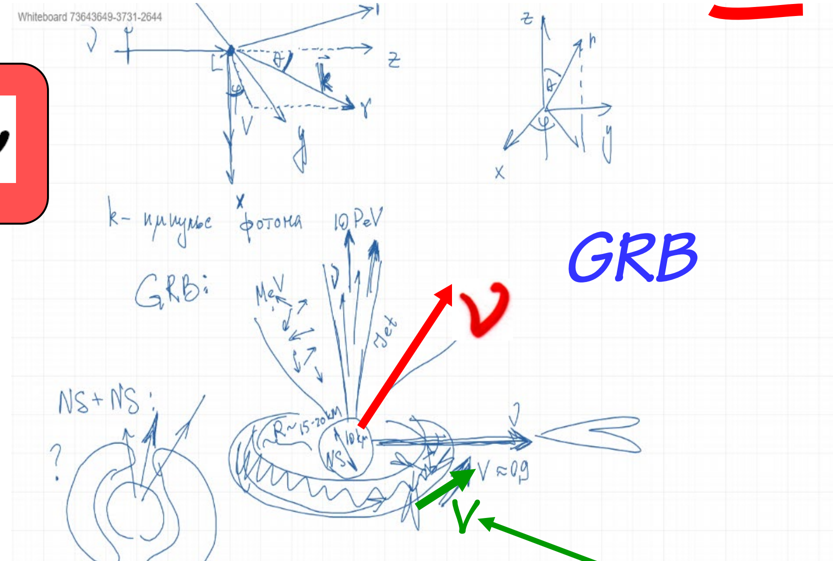
$SL\nu$

$\omega_{SL\nu}$ depends on $(\vec{V}_{matt} \vec{P}_{\nu})$

Abstract We found the spin integral of motion for neutrinos propagating in moving and polarized matter. Contrary to all previous studies this is the exact spin operator commuting with the Hamiltonian for a neutrino in matter which moves in an arbitrary direction relative to the direction of neutrino propagation. The operator obtained opens up the possibility of consistent classification of neutrino states in such a medium and, as a consequence, a systematic description of the related physical phenomena. Using the operator, we obtain a dispersion relation for neutrinos in arbitrary moving matter and consider its particular cases.



$$\left\{ i\gamma_\mu \partial^\mu - \frac{1}{2}\gamma_\mu(1+\gamma^5) f^\mu - m \right\} \Psi(x) = 0$$

$$\frac{1}{2} f^\mu = \tilde{n}_0 v^\mu, \quad \tilde{n}_0 = \frac{1}{2\sqrt{2}} G_F (1 + 4\sin^2\theta_W) n_0$$

$$n = \gamma n_0 \quad \gamma = 1/\sqrt{1 - \mathrm{v}^2}$$

$$\mathrm{S} = \gamma \left[\gamma^5 \gamma^0 m - \gamma^5 (\tilde{\mathrm{H}} - (\tilde{\mathbf{p}}\mathbf{v})) - m\gamma^0 (\mathbf{\Sigma}\mathbf{v}) \right]$$

$$E_{s=+1} = \sqrt{(\mathrm{p} - \tilde{n})^2 + m^2} + \tilde{n},$$

$$E_{s=-1} = \sqrt{\mathrm{p}^2 + 4\tilde{n}^2 \mathrm{v}^2 + m^2} + 2\tilde{n}$$

Flavour oscillations $\nu_e \leftrightarrow \nu_\mu$ **in moving matter**

Wofenstein term

$$A \to A' = \sqrt{2} G_F n + 8 \frac{G_F^2 n^2 \mathrm{v}^2}{\mathrm{p}} \sin^2\theta_W$$

Probability of oscillations

$$P = \frac{1}{2} \frac{\Delta^2 \sin^2 2\theta}{(\Delta \cos 2\theta - 2\mathrm{p}A')^2 + \Delta^2 \sin^2 2\theta}$$

Resonance condition

$$\frac{\Delta}{2\mathrm{p}} \cos 2\theta = A + 8 \frac{G_F^2 n^2 \mathrm{v}^2}{\mathrm{p}} \sin^2\theta_W$$

two terms are of same order for extra-dense matter $n \sim 10^{41}\ \mathrm{cm}^{-3}$ and $V \sim 0.9$, $P \sim 10$ keV

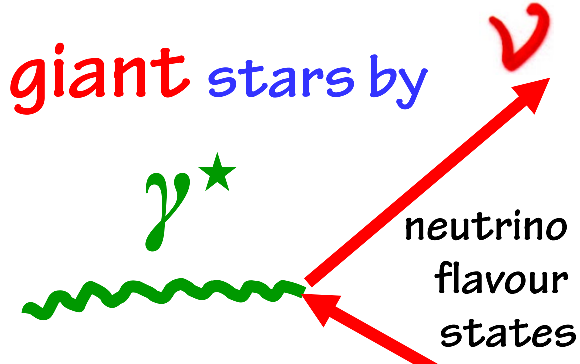
Electromagnetic ν in astrophycis and bounds on μ_ν and q_ν

Astrophysical bounds on μ_ν

② Astrophysical bound on μ_ν

G.Raffelt, PRL 1990

comes from cooling (observed luminosity) of **red giant** stars by plasmon decay $\gamma^\star \to \nu\nu$



$$L_{int} = \frac{1}{2}\sum_{a,b}\left(\mu_{a,b}\bar{\psi}_a\sigma_{\mu\nu}\psi_b + \epsilon_{a,b}\bar{\psi}_a\sigma_{\mu\nu}\gamma_5\psi_b\right)$$

J.Bernstein et al 1963 ν : $\mu_\nu \sim 10^{-10}\mu_B$ ν

solar

Matrix element

$$|M|^2 = M_{\alpha\beta}p^\alpha p^\beta, \quad M_{\alpha\beta} = 4\mu^2(2k_\alpha k_\beta - 2k^2\epsilon^*_\alpha\epsilon_\beta - k^2 g_{\alpha,\beta}), \quad \epsilon_\alpha k^\alpha = 0$$

Decay rate

$$\Gamma_{\gamma\to\nu\bar{\nu}} = \frac{\mu^2}{24\pi}\frac{(\omega^2-k^2)^2}{\omega}$$

= 0 in vacuum $\omega = k$

In the classical limit $\gamma^\star$ - like a massive particle with $\omega^2 - k^2 = \omega^2_{pl}$

Energy-loss rate per unit volume

$$Q_\mu = g\int\frac{d^3k}{(2\pi)^3}\omega f_{BE}\Gamma_{\gamma\to\nu\bar{\nu}}$$

$$\mu^2 \to \sum_{a,b}\left(|\mu_{a,b}|^2 + |\epsilon_{a,b}|^2\right)$$

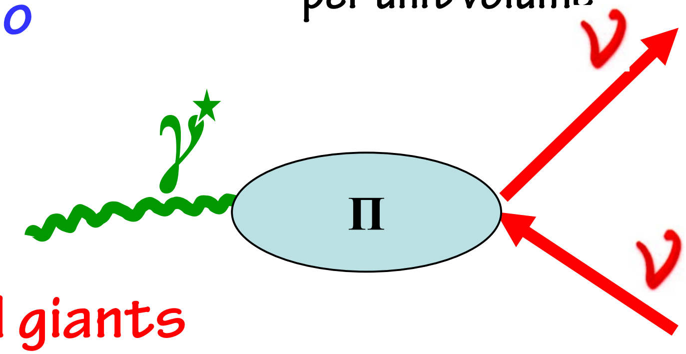
distribution function of plasmons

Astrophysical bound on μ_ν

$$Q_\mu = g \int \frac{d^3k}{(2\pi)^3} \omega f_{BE} \Gamma_{\gamma \to \nu\bar{\nu}}$$

Energy-loss rate per unit volume

Magnetic moment plasmon decay enhances the Standard Model photo-neutrino cooling by photon polarization tensor



more fast star cooling

slightly reducing the core temperature

delay of helium ignition in low-mass red giants (due to nonstandard ν losses)

astronomical observable

can be related to luminosity of stars before and after helium flash

… in order not to delay helium ignition in an unacceptable way (a significant brightness increase is constraint by observations …)

… best astrophysical limit on ν magnetic moment…

$$\mu \leq 3 \times 10^{-12} \mu_B$$

G.Raffelt, PRL 1990

D+M

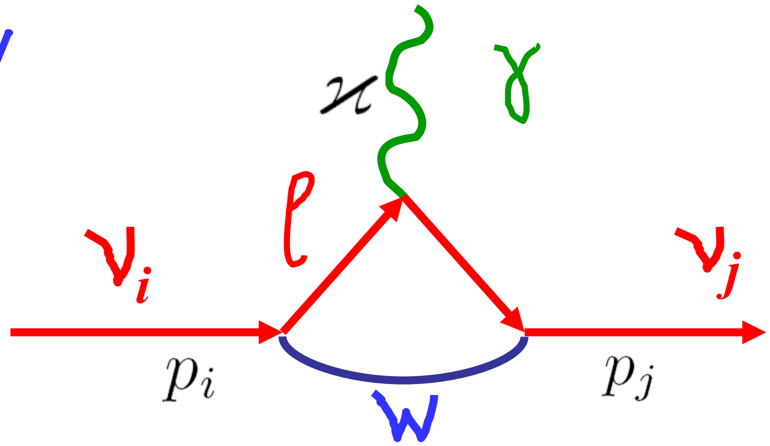
$$\mu^2 \to \sum_{a,b} \left(|\mu_{a,b}|^2 + |\epsilon_{a,b}|^2 \right)$$

Neutrino radiative decay

$\nu_i \longrightarrow \nu_j + \gamma$

$m_i > m_j$

$$L_{int} = \frac{1}{2}\bar{\psi}_i \sigma_{\alpha\beta}(\sigma_{ij} + \epsilon_{ij}\gamma_5)\psi_j F^{\alpha\beta} + h.c.$$



Radiative decay rate

Petkov 1977; Zatsepin, Smirnov 1978;
Bilenky, Petkov 1987; Pal, Wolfenstein 1982

$$\Gamma_{\nu_i \to \nu_j + \gamma} = \frac{\mu^2_{eff}}{8\pi}\left(\frac{m_i^2 - m_j^2}{m_i^2}\right)^3 \approx 5\left(\frac{\mu_{eff}}{\mu_B}\right)^2\left(\frac{m_i^2 - m_j^2}{m_i^2}\right)^3\left(\frac{m_i}{1\ eV}\right)^3 s^{-1}$$

$$\mu^2_{eff} = |\mu_{ij}|^2 + |\epsilon_{ij}|^2$$

- Radiative decay has been constrained from absence of decay photons:

1) reactor $\bar{\nu}_e$ and solar ν_e fluxes, — Raffelt 1999

2) SN 1987A ν burst (all flavours), — Kolb, Turner 1990;

3) spectral distortion of CMBR — Ressell, Turner 1990

... important for astrophysics consequence of μ_ν is appearance ν_R ... examples 1-3 ...

1

a) helicity change in ν magnetic moment scattering on e (p, n)

(active) $\nu_L \Rightarrow \nu_R$ (sterile)

$$\left(\frac{d\sigma}{dT}\right)_{\mu_\nu} = \frac{\pi\alpha_{em}^2}{m_e^2}\left[\frac{1-T/E_\nu}{T}\right]\mu_\nu^2$$

effective μ_ν

$$\mu_\nu^2(\nu_l, L, E_\nu) = \sum_j \left| \sum_i U_{li} e^{-iE_iL} \mu_{ji} \right|^2$$

b) spin (spin-flavor) precession in $B_\perp$

$$\mu_{ij} \rightarrow |\mu_{ij} - i\epsilon_{ij}|$$

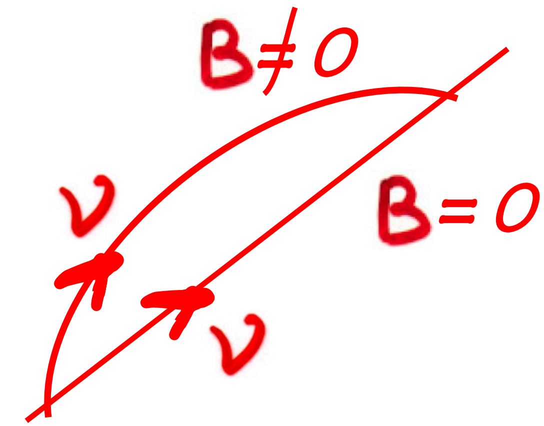
electric dipole moment

c) spin (spin-flavor) precession in transversal matter currents $j_\perp$ or polarization $\zeta_\perp$

$\nu_L \Leftrightarrow \nu_R$ $\quad\quad$ $\nu_L \Leftrightarrow \nu_R$

2

... important for astrophysics consequence of $q_\nu \neq 0$ is ν deviation from a rectilinear trajectory



... example 4 ...

Astrophysics bounds on μ_ν

1) SN 1987A provides energy-loss limit on μ_ν related to observed duration of ν signal

(also d_ν and transition moments)

...in magnetic moment scattering $\nu_e^L + e \to \nu_e^R + e$

due to change of helicity $\nu_L \Rightarrow \nu_R$

Dar, Nussinov & Rephaeli, Goldman et al, Notzol, Voloshin, Ayla et al, Balantekin et 1988

proto-neutron star formed in core-collapse SN can cool faster since ν_R are sterile and not trapped in a core like ν_L for a few sec

- escaping ν_R will cool the core very efficient and fast (~ 1 s)

the observed 5-10 s pulse duration in Kamioka II and IMB is in agreement with the standard model ν_L trapping ...

$\Downarrow$

$$\mu_\nu^D \sim 10^{-12}\mu_B$$

... inconsistent with SN1987A observed cooling time

Barbieri, Mahapatra Lattimer, Cooperstein, 1988
Raffelt, 1996

Astrophysics bounds on μ_ν

... example 5...

2) SN 1987A provides energy-loss limit on μ_ν
related to observed ν energies

... helicity change in ν magnetic moment scattering $\nu_e^L + e \to \nu_e^R + e$
on e (p, n)

ν_R from inner SN core have larger energy than ν_L emitted from neutrino sphere

then $\nu_R \overset{B}{\leftrightarrow} \nu_L$ in galactic B and higher-energy ν_L would arrive to detector as a signal of SN 1987A

⇒ from absence of anomalous high-energy ν — Nötzold 1988

$$\mu_\nu^D \sim 10^{-12} \mu_B$$

Astrophysics bounds on μ_ν

$$\mu_\nu(\text{astro}) < 10^{-10} – 10^{-12}\ \mu_B$$

Mostly derived from consequences of helicity-state change in astrophysical medium:

- available degrees of freedom in BBN
- stellar cooling via plasmon decay
- cooling of SN1987a

Red Giant lumin.
$\mu_\nu < 3 \cdot 10^{-12} \mu_B$
G. Raffelt, D. Dearborn, J. Silk, 1989.

Bounds depend on

- modeling of astrophysical system,
- on assumption on he neutrino properties .

Generic assumption:

- absence of other nonstandard interactions accept for μ_ν

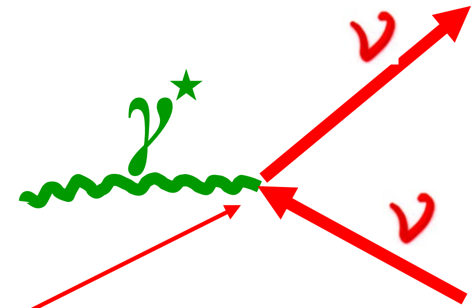
A global treatment would be desirable, incorporating oscillations and matter effects, as well as the complications due to interference and competitions among various channels

Astrophysical bounds on q_ν

Constraints on neutrino millicharge from red giants cooling

- **Plasma process** (photon decay)

$$\gamma^{\star} \rightarrow \nu\nu$$



Interaction Lagrangian

$$L_{int} = -iq_{\nu}\bar{\psi}_{\nu}\gamma^{\mu}\psi_{\nu}A^{\mu}$$

millicharge

Decay rate

$$\Gamma_{q_{\nu}} = \frac{{q_{\nu}}^2}{12\pi}\omega_{pl}\left(\frac{\omega_{pl}}{\omega}\right)$$

Dobroliubov, Ignatiev 1990;
Babu, Volkas 1992;
Mohapatra, Nussinov 1992 ...

Delay of helium ignition in low-mass red gians due to nonstandard ν losses

$$q_{\nu} \leq 2 \times 10^{-14} e$$

...to avoid delay of helium ignition in low-mass red giants

Halt, Raffelt, Weiss, PRL1994

$$q_{\nu} \leq 3 \times 10^{-17} e$$

... absence of anomalous energy-dependent dispersion of SN1987A ν signal, most model independent

- $$q_{\nu} \leq 3 \times 10^{-21} e$$

... from "charge neutrality" of neutron...

- … astrophysical bound on millicharge q_ν from

ν energy quantization in rotating magnetized star

Grigoriev, Savochkin, Studenikin, Russ. Phys. J. 50 (2007) 845
Studenikin, J. Phys. A: Math. Theor. 41 (2008) 164047
Balantsev, Popov, Studenikin, J. Phys. A: Math. Theor. 44 (2011) 255301
Balantsev, Studenikin, Tokarev, Phys. Part. Nucl. 43 (2012) 727
Phys. Atom. Nucl. 76 (2013) 489

- Studenikin, Tokarev, Nucl. Phys. B 884 (2014) 396

Millicharged ν in rotating magnetized star

Balatsev, Tokarev, Studenikin,
Phys.Part.Nucl., 2012,
Phys.Atom.Nucl., Nucl.Phys. B, 2013,

- Studenikin, Tokarev, Nucl.Phys.B (2014)

Modified Dirac equation for ν wave function

$$\left(\gamma_\mu(p^\mu + q_0 A^\mu) - \frac{1}{2}\gamma_\mu(c_l+\gamma_5)f^\mu - \frac{i}{2}\mu\sigma_{\mu\nu}F^{\mu\nu} - m\right)\Psi(x) = 0$$

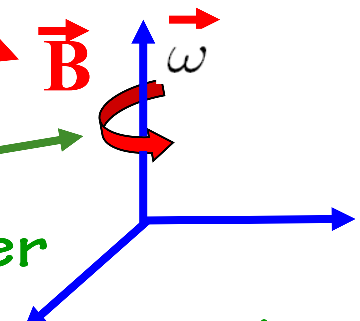
external magnetic field

$$V_m = \tfrac{1}{2}\gamma_\mu(c_l+\gamma_5)f^\mu$$

$$c_l = 1$$

matter potential

rotating matter



rotation angular frequency

$$f^\mu = -Gn_n(1, -\epsilon y\omega, \epsilon x\omega, 0)$$

matter density

- ν energy is quantized in rotating and magnetized star

A.Studenikin, I.Tokarev, Nucl.Phys.B (2014)

$$G = \frac{G_F}{\sqrt{2}}$$

$$p_0 = \sqrt{p_3^2 + 2N|2Gn_n\omega - \epsilon q_\nu B| + m^2} - Gn_n - q\phi$$

$N = 0, 1, 2, \ldots$ integer number

matter rotation frequency

millicharge

scalar potential of electric field

! ν energy is quantized in rotating matter like electron energy in magnetic field (Landau energy levels):

- $$p_0^{(e)} = \sqrt{m_e^2 + p_3^2 + 2\gamma N}, \quad \gamma = eB, \quad N = 0, 1, 2, \ldots$$

In quasi-classical approach

✓ quantum states in rotating matter

✓ motion in circular orbits

$$R = \int_0^\infty \Psi_L^\dagger \, \mathrm{r} \, \Psi_L \, d\mathbf{r} = \sqrt{\frac{2N}{|2Gn_n\omega - \epsilon q_0 B|}}$$

$N = 1,2,3 \ldots$

due to effective Lorentz force

$$\mathbf{F}_{eff} = q_{eff}\mathbf{E}_{eff} + q_{eff}\left[\boldsymbol{\beta} \times \mathbf{B}_{eff}\right]$$

A. Studenikin,
J.Phys.A: Math.Theor.
41(2008) 164047

$$q_{eff}\mathbf{E}_{eff} = q_m\mathbf{E}_m + q_0\mathbf{E}$$

$$q_{eff}\mathbf{B}_{eff} = |q_m B_m + q_0 B|\mathbf{e}_z$$

where

$$q_m = -G, \quad \mathbf{E}_m = -\boldsymbol{\nabla} n_n, \quad \mathbf{B}_m = 2n_n\boldsymbol{\omega}$$

matter density

matter rotation frequency

matter induced "charge", "electric" field, "magnetic" fields

=

... we predict :

A.Studenikin, I.Tokarev, Nucl.Phys.B (2014)

$E \sim 1\ eV$

1) low-energy ν are trapped in circular orbits inside rotating neutron stars

$$R = \sqrt{\frac{2N}{Gn\omega}} < R_{NS} = 10\ km$$

$$R_{NS} = 10\ km$$
$$n = 10^{37} cm^{-3}$$
$$\omega = 2\pi \times 10^{3}\ s^{-1}$$

2) rotating neutron stars as filters for low-energy relic ν ?

$$T_\nu \sim 10^{-4}\ \mathrm{eV}$$

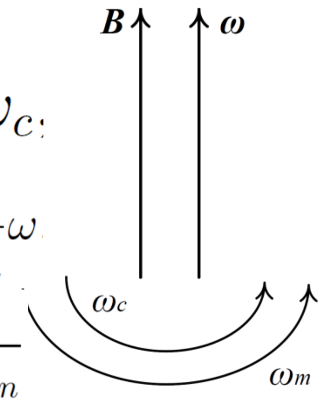
Millicharged ν as star rotation engine

- Single ν generates feedback force with projection on rotation plane
- $F = (q_0 B + 2Gn_n\omega)\sin\theta$

$$\Omega = \omega_m + \omega_c;$$

$$\omega_m = \frac{2Gn_n}{p_0 + Gn_n}\omega,$$

$$\omega_c = \frac{q_0 B}{p_0 + Gn_n}$$

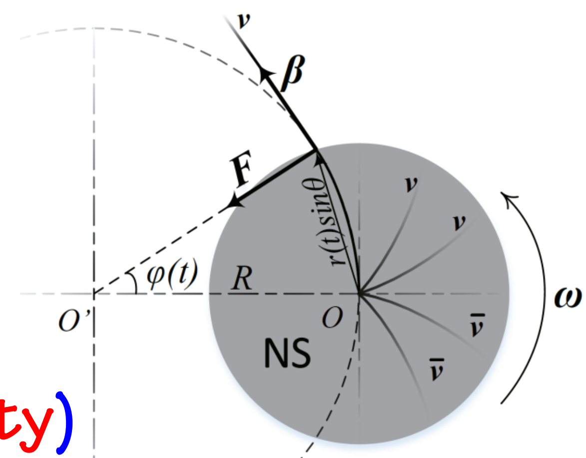


single ν torque

- $$M_0(t) = \sqrt{1 - \frac{r^2(t)\Omega^2\sin^2\theta}{4}}\,Fr(t)\sin\theta$$

total N_ν torque

$$M(t) = \frac{N_\nu}{4\pi}\int M_0(t)\sin\theta d\theta d\varphi$$



- Should effect initial star rotation (shift of star angular velocity)

$$|\Delta\omega| = \frac{5N_\nu}{6M_S}(q_0 B + 2Gn_n\omega_0)$$

$$\Delta\omega = \omega - \omega_0$$

A.Studenikin, I.Tokarev, Nucl.Phys.B (2014)

- **ν Star Turning mechanism (νST)**

Studenikin, Tokarev, Nucl. Phys. B 884 (2014) 396

Escaping millicharged νs move on curved orbits inside magnetized rotating star and feedback of effective Lorentz force should effect initial star rotation

- **New astrophysical constraint on ν millicharge**

$$\frac{|\Delta\omega|}{\omega_0} = 7.6\varepsilon \times 10^{18} \left(\frac{P_0}{10\ \mathrm{s}}\right)\left(\frac{N_\nu}{10^{58}}\right)\left(\frac{1.4M_\odot}{M_S}\right)\left(\frac{B}{10^{14}G}\right)$$

- $|\Delta\omega| < \omega_0$! ...to avoid contradiction of νST impact with observational data on pulsars ...

$$q_0 < 1.3 \times 10^{-19} e_0$$

.. best astrophysical bound ... !

= New developments in ν spin and flavour oscillation

... new astrophysical probes of ν

(1) generation of ν spin (flavour) oscillations by interaction with transversal matter current $j_{\perp}$

P.Pustoshny, Studenikin,
Neutrino spin and spin-flavour oscillations in transversal matter currents with standard and non-standard interactions

- Phys. Rev. D98 (2018) 113009

Studenikin, Neutrino in electromagnetic fields and moving matter

- Phys.Atom.Nucl. 67 (2004) 993-1002

(2) inherent interplay of ν spin and flavour oscillations in B

A. Popov, Studenikin,
Neutrino eigenstates and flavour, spin and spin-flavor oscillations in a constant magnetic field

- Eur. Phys. J. C79 (2019) 144

= Main steps in ν oscillations

67 years!

early history of ν oscillations

1. $\nu_e \overset{vac}{\longleftrightarrow} \bar{\nu}_e$, B. Pontecorvo, 1957

2. $\nu_e \overset{vac}{\longleftrightarrow} \nu_\mu$, Z. Maki, M. Nakagava, S. Sakata, 1962

3. $\nu_e \overset{matter, \rho = const}{\longleftrightarrow} \nu_\mu$, L. Wolfenstein, 1978

4. $\nu_e \overset{matter, \rho \neq const}{\longleftrightarrow} \nu_\mu$, S. Mikheev, A. Smirnov, 1985

• resonances in ν flavour oscillations ⇒ MSW-effect, solution for $\nu_\odot$ - problem

5. $\nu_{e_L} \overset{B_\perp}{\longleftrightarrow} \nu_{e_R}$, A. Cisneros, 1971; M. Voloshin, M. Vysotsky, L. Okun, 1986, $\nu_\odot$

6. $\nu_{e_L} \overset{B_\perp}{\longleftrightarrow} \nu_{e_R}, \nu_{\mu_R}$, E. Akhmedov, 1988; C.-S. Lim & W. Marciano, 1988

• resonances in ν spin (spin-flavour) oscillations in matter

Bruno Pontecorvo 1913–1993

only in $B_\perp$ and matter at rest

ν spin and spin-flavour oscillations in $B_\perp$

$$\nu_{e_L} \leftrightarrow \nu_{\mu_R} \qquad B = |\mathbf{B}_\perp| e^{i\phi(t)}$$

- ... twisting magnetic field ...

- $$P_{\nu_L \nu_R} = \sin^2\beta \, \sin^2\Omega z \qquad \sin^2\beta = \frac{(\mu_{e\mu} B)^2}{(\mu_{e\mu} B)^2 + \left(\frac{\Delta_{LR}}{4E}\right)^2}$$

$$\Omega^2 = (\mu_{e\mu} B)^2 + \left(\frac{\Delta_{LR}}{4E}\right)^2$$

$$\Delta_{LR} = \frac{\Delta m^2}{2}(\cos 2\theta + 1) - 2EV_{\nu_e} + 2E\dot{\phi}$$

- Resonance amplification of oscillations in matter:

$$\Delta_{LR} \to 0 \quad \Rightarrow \quad \sin^2\beta \to 1$$

Akhmedov, 1988
Lim, Marciano

... similar to MSW effect

Neutrino oscillations in the magnetic field of the sun, supernovae, and neutron stars

G. G. Likhachev and A. I. Studenikin

M. V. Lomonosov Moscow State University, 119899 Moscow, Russia
(Submitted 10 March 1995)
Zh. Éksp. Teor. Fiz. **108**, 769–782 (September 1995)

We examine the feasibility of oscillations of Dirac and Majorana neutrinos in a strong magnetic field (assuming a nonvanishing neutrino magnetic moment). We determine the critical magnetic field $\tilde{B}_{cr}(\Delta m_\nu^2, \theta, n_{eff}, E_\nu, \dot{\phi}(t))$ as a function of the neutrino mass difference, the vacuum mixing angle, the effective mass density, the neutrino energy, and the angle specifying the variation of the magnetic field in the plane transverse to the neutrino's motion. The conditions under which magnetic field-induced neutrino oscillations are significant are discussed. We study the possibility that such oscillations come about in supernova explosions, neutron stars, the sun, and the interstellar medium. We analyze the possible conversion of half the active neutrinos in a beam into sterile neutrinos when the beam emerges from the surface of a neutron star (cross-boundary effect), as well as when it crosses the interface between internal layers of a neutron star.

! "cross-boundary effect" →

$\tilde{B}_{cr}$ "critical magnetic field"

$$\tilde{B}_{cr} = \left| \frac{1}{2\tilde{\mu}} \left(\frac{\Delta m_\nu^2}{2E_\nu} A - \sqrt{2} G_F n_{eff} + \dot{\phi} \right) \right|$$

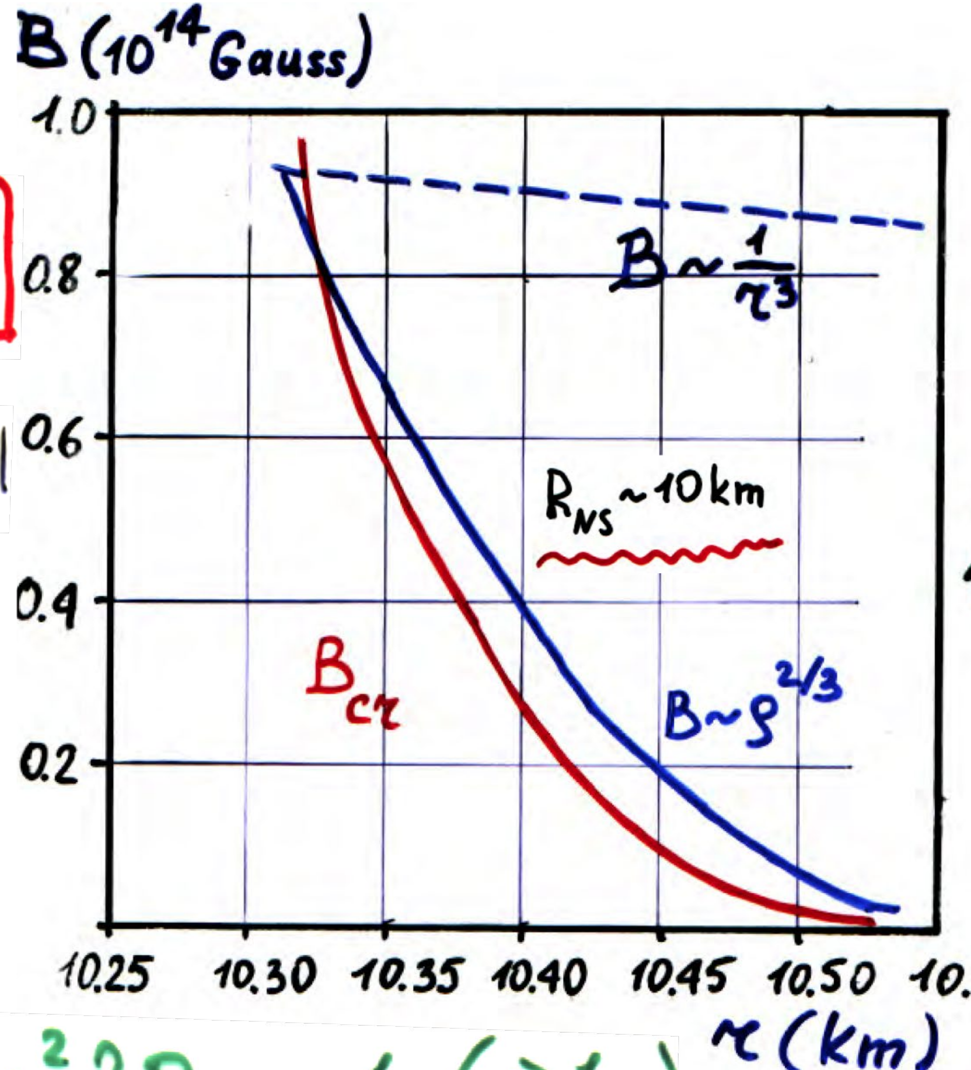
$$\tilde{B} = 10^2 \frac{\mu_B}{\tilde{\mu}} \left| - \frac{n_{eff}}{10^{31} cm^{-3}} + 0.4 A \left(\frac{\Delta m^2}{1 eV^2} \right) \left(\frac{1 MeV}{E} \right) + \frac{1 m}{L_{\dot{\phi}}} \right|$$

$\nu_L \to \nu_R$

active neutrino → sterile neutrino

$$\bar{P}_{\nu_R} = \frac{1}{2} \sin^2 2\theta_{eff}$$

If $B \gtrsim \tilde{B}_{cr}$ => $\sin^2 2\theta_{eff} \sim 1$ $(\gtrsim \frac{1}{2})$



"cross-boundary effect"

Likhachev, Studenikin,
Preprint ICTP, IC/94/70, 1994
Sov. Phys. JETP 108 (1995)

PHYSICAL REVIEW LETTERS **132,** 101005 (2024)

Resonances of Supernova Neutrinos in Twisting Magnetic Fields

Sudip Jana[1,*] and Yago Porto[2,†]
[1]*Max-Planck-Institut für Kernphysik, Saupfercheckweg 1, 69117 Heidelberg, Germany*
[2]*Instituto de Física Gleb Wataghin - UNICAMP, 13083-859, Campinas, São Paulo, Brazil*

(Received 28 March 2023; revised 20 June 2023; accepted 14 February 2024; published 5 March 2024)

We investigate the effect of resonant spin conversion of the neutrinos induced by the geometrical phase in a twisting magnetic field. We find that the geometrical phase originating from the rotation of the transverse magnetic field along the neutrino trajectory can trigger a resonant spin conversion of Dirac neutrinos inside the supernova, even if there were no such transitions in the fixed-direction field case. We have shown that, even though resonant spin conversion is too weak to affect solar neutrinos, it could have a remarkable consequence on supernova neutronization bursts where very intense magnetic fields are quite likely. We demonstrate how the flavor composition at Earth can be used as a probe to establish the presence of non-negligible magnetic moments, potentially down to $10^{-15}\mu_B$ in upcoming neutrino experiments like the Deep Underground Neutrino Experiment and the Hyper-Kamiokande. Possible implications are analyzed.

DOI: 10.1103/PhysRevLett.132.101005

(1)

Neutrino spin $\nu_e^L \Leftarrow (j_\perp) \Rightarrow \nu_e^R$ and

spin-flavour $\nu_e^L \Leftarrow (j_\perp) \Rightarrow \nu_\mu^R$

oscillations engendered

by transversal matter current $j_\perp$

! without ($\mu_\nu B$)

A. Studenikin,
Neutrino in electromagnetic fields and moving matter,
Phys. Atom. Nucl. 67 (2004) 993-1002

P. Pustoshny, A. Studenikin,
Neutrino spin and spin-flavour oscillations in transversal matter currents with standard and non-standard interactions, Phys. Rev. D98 (2018) 113009

Physics of Atomic Nuclei, Vol. 67, No. 5, 2004, pp. 993–1002. Translated from Yadernaya Fizika, Vol. 67, No. 5, 2004, pp. 1014–1024.

Phys.Atom.Nucl. 67 (2004) 993-1002

Neutrino in Electromagnetic Fields and Moving Media

A. I. Studenikin*

Moscow State University, Vorob'evy gory, Moscow, 119899 Russia

Received March 26, 2003; in final form, August 12, 2003

The possible emergence of neutrino-spin oscillations (for example, $\nu_{eL} \leftrightarrow \nu_{eR}$) owing to neutrino interaction with matter under the condition that there exists a nonzero transverse current component or matter polarization (that is, $\mathbf{M}_{0\perp} \neq 0$) is the most important new effect that follows from the investigation of neutrino-spin oscillations in Section 4. So far, it has been assumed that neutrino-spin oscillations may arise only in the case where there exists a nonzero transverse magnetic field in the neutrino rest frame.

STUDENIKIN PHYSICS OF ATOMIC NUCLEI Vol. 67 No. 5 2004

Consider spin, spin-flavour

$$\nu_{e_L} \to \nu_{e_R}, \quad \nu_{e_L} \to \nu_{\mu_R}$$

$$P(\nu_i \to \nu_j) = \sin^2(2\theta_{\text{eff}}) \sin^2 \frac{\pi x}{L_{\text{eff}}}, \quad i \neq j$$

$$L_{\text{eff}} = \frac{2\pi}{\sqrt{E_{\text{eff}}^2 + \Delta_{\text{eff}}^2}}$$

$$\sin^2 2\theta_{\text{eff}} = \frac{E_{\text{eff}}^2}{E_{\text{eff}}^2 + \Delta_{\text{eff}}^2}, \quad \Delta_{\text{eff}}^2 = \frac{\mu}{\gamma_\nu}\left|\mathbf{M}_{0\parallel} + \mathbf{B}_{0\parallel}\right|. \quad E_{\text{eff}} = \mu\left|\mathbf{B}_\perp + \frac{1}{\gamma_\nu}\mathbf{M}_{0\perp}\right|,$$

A.Studenikin, "Neutrinos in electromagnetic fields and moving media", Phys. Atom. Nucl. 67 (2004)

• transversal current $\mathbf{j}_\perp$

$$\vec{M}_0 = \gamma_\nu \rho n_e \left(\vec{\beta}_\nu (1 - \vec{\beta}_\nu \vec{v}_e) - \frac{1}{\gamma_\nu} \vec{v}_{e\perp}\right),$$

speed of ν; transversal speed of matter; matter density; $\parallel$; $\perp$

$$\gamma_\nu = \frac{E_\nu}{m_\nu},$$

where

$$\rho = \frac{G_F}{2\mu_\nu\sqrt{2}}(1 + 4\sin^2\theta_W)$$

… the effect of ν helicity conversions and oscillations induced by transversal matter currents has been confirmed in studies of ν propagation in astrophysical media:

$\nu_{e_L} \to \nu_{e_R}, \quad \nu_{e_L} \to \nu_{\mu_R}$

- J. Serreau and C. Volpe, Neutrino-antineutrino correlations in dense anisotropic media, Phys. Rev. D90 (2014) 125040
- V. Cirigliano, G. M. Fuller, and A. Vlasenko, A new spin on neutrino quantum kinetics Phys. Lett. B747 (2015) 27
- A. Kartavtsev, G. Raffelt, and H. Vogel, Neutrino propagation in media: flavor-, helicity-, and pair correlations, Phys. Rev. D91 (2015) 125020 …

,

Neutrino spin (spin-flavour) oscillations in transversal matter currents

… quantum treatment …

- ν spin evolution effective Hamiltonian in moving matter ? transversal and longitudinal currents $\vec{j}_{\perp} + \vec{j}_{\parallel}$
- two flavor ν with two helicities: $\nu_f = (\nu_e^+, \nu_e^-, \nu_\mu^+, \nu_\mu^-)^T$

ν interaction with matter composed of neutrons: $n = \frac{n_0}{\sqrt{1-v^2}}$ neutron number density in laboratory reference frame

$\mathbf{v} = (v_1, v_2, v_3)$ velocity of matter

- $$L_{int} = -f^\mu \sum_l \bar{\nu}_l(x)\gamma_\mu \frac{1+\gamma_5}{2}\nu_l(x) = -f^\mu \sum_i \bar{\nu}_i(x)\gamma_\mu \frac{1+\gamma_5}{2}\nu_i(x)$$ $l = e,\ or\ \mu$, $i = 1,\ 2$

$$f^\mu = -\frac{G_F}{2\sqrt{2}} j_n^\mu$$

- $$j_n^\mu = n(1, \mathbf{v})$$

$$\nu_e^\pm = \nu_1^\pm \cos\theta + \nu_2^\pm \sin\theta,$$
$$\nu_\mu^\pm = -\nu_1^\pm \sin\theta + \nu_2^\pm \cos\theta$$

ν flavour and mass states

P. Pustoshny, A. Studenikin,
Phys. Rev. D98 (2018) 113009

ν (2 flavours x 2 helicities) evolution equation

$$i\frac{d}{dt}\nu_f^s = \Big(H_0 + \Delta H_0^{SM} + \Delta H_{j_{||}+j_\perp}^{SM} + \Delta H_{B_{||}+B_\perp}^{SM} + \Delta H_0^{NSI} + \Delta H_{j_{||}+j_\perp}^{NSI}\Big)\nu_f^s$$

vacuum — matter at rest — moving matter — B — matter at rest — moving matter

Standard Model — Non-Standard Interactions

Resonant amplification of ν oscillations:

- $\nu_e^L \Leftarrow (j_\perp) \Rightarrow \nu_e^R$ by longitudinal matter current $j_{||}$
- $\nu_e^L \Leftarrow (j_\perp) \Rightarrow \nu_e^R$ by longitudinal $B_{||}$
- $\nu_e^L \Leftarrow (j_\perp) \Rightarrow \nu_\mu^R$ by matter-at-rest effect
- $\nu_e^L \Leftarrow (j_\perp^{NSI}) \Rightarrow \nu_\mu^R$ by matter-at-rest effect

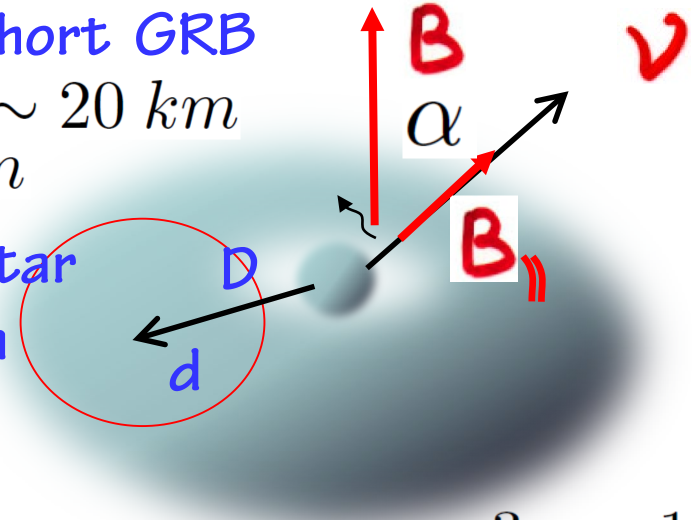
P. Pustoshny, A. Studenikin, Phys. Rev. D98 (2018) 113009

$$\boxed{\nu_e^L \Leftarrow (j_\perp) \Rightarrow \nu_e^R}$$

a model of short GRB

$D \sim 20\ km$

$d \sim 20\ km$



- Consider ν escaping central neutron star with inclination angle α from accretion disk: $B_{\parallel} = B \sin\alpha \sim \frac{1}{2} B$
- Toroidal bulk of rotating dense matter with $\omega = 10^3\ s^{-1}$
- transversal velocity of matter

$v_\perp = \omega D = 0.067$ and $\gamma_n = 1.002$

- Perego et al, Mon.Not.Roy.Astron.Soc. 443 (2014) 3134
- Grigoriev, Lokhov, Studenikin, Ternov, JCAP 1711 (2017) 024

$$E_{eff} = \left(\frac{\eta}{\gamma}\right)_{ee} \tilde{G} n v_\perp = \frac{\cos^2\theta}{\gamma_{11}} \tilde{G} n v_\perp \approx \tilde{G} n_0 \frac{\gamma_n}{\gamma_\nu} v_\perp$$

$$\Delta_{eff} = \left| \left(\frac{\mu}{\gamma}\right)_{ee} \boldsymbol{B}_{||} + \eta_{ee} \tilde{G} n \boldsymbol{\beta} \right| \approx \left| \frac{\mu_{11}}{\gamma_\nu} B_{||} - \tilde{G} n_0 \gamma_n \right|$$

$$\boldsymbol{B}_{||} \boldsymbol{\beta} = -1$$

$$\boxed{E_{eff} \geq \Delta_{eff}}$$

resonance condition

$$\boxed{\left| \frac{\mu_{11} B_{||}}{\tilde{G} n_0 \gamma_n} - \gamma_\nu \right| \leq 1}$$

Resonance amplification of spin-flavor oscillations (in the absence of $j_{\parallel}$)

$$\nu_e^L \Leftarrow (j_\perp, B_\perp) \Rightarrow \nu_\mu^R$$

$$\vec{B} = \vec{B}_\perp + \vec{B}_\parallel \rightarrow 0$$

Criterion – oscillations are important:

$$\sin^2 2\theta_{\text{eff}} = \frac{E_{\text{eff}}^2}{E_{\text{eff}}^2 + \Delta_{\text{eff}}^2} \geq \frac{1}{2}$$

$$E_{eff} = \left| \mu_{e\mu} B_\perp + \left(\frac{\eta}{\gamma}\right)_{e\mu} \tilde{G} n v_\perp \right| \geq \left| \Delta M - \frac{1}{2}\left(\frac{\mu_{11}}{\gamma_{11}} + \frac{\mu_{22}}{\gamma_{22}}\right) B_\parallel - \tilde{G} n (1 - \boldsymbol{v\beta}) \right|$$

neglecting $\vec{B} = \vec{B}_\perp + \vec{B}_\parallel \rightarrow 0$:

$$L_{eff} = \frac{\pi}{\left(\frac{\eta}{\gamma}\right)_{e\mu} \tilde{G} n v_\perp} \qquad \left(\frac{\eta}{\gamma}\right)_{e\mu} \approx \frac{\sin 2\theta}{\gamma_\nu}$$

$$\left| \left(\frac{\eta}{\gamma}\right)_{e\mu} \tilde{G} n v_\perp \right| \geq \left| \Delta M - \tilde{G} n (1 - \boldsymbol{v\beta}) \right| \Rightarrow \tilde{G} n \sim \Delta M$$

$$\Delta m^2 = 7.37 \times 10^{-5} \ eV^2$$
$$\sin^2 \theta = 0.297$$
$$p_0^\nu = 10^6 \ eV$$

$$\tilde{G} = \frac{G_F}{2\sqrt{2}} = 0.4 \times 10^{-23} \ eV^{-2}$$

$$\Rightarrow \Delta M = 0.75 \times 10^{-11} eV$$

$$n_0 \sim \frac{\Delta M}{\tilde{G}} = 10^{12} \ eV^3 \approx 10^{26} \ cm^{-3}$$

$$L_{eff} = \frac{\pi}{(\frac{\eta}{\gamma})_{e\mu} \tilde{G} n v_\perp} \approx 5 \times 10^{11} \ km$$

- $L_{eff} \approx 10 \ km$ (within short GRB) if $n_0 \approx 5 \times 10^{36} \ cm^{-3}$

2 ν

A.Popov, A.Studenikin, Eur. Phys .J. C79 (2019) 144

"Neutrino eigenstates and
flavour, spin and spin-flavour oscillations
in a constant magnetic field"

- $\boxed{\nu_e^L \leftrightarrow \nu_\mu^L}$ $\boxed{\nu_e^L \leftrightarrow \nu_e^R}$ $\boxed{\nu_e^L \leftrightarrow \nu_\mu^R}$

Consider two flavour ν with two helicities as superposition of helicity mass states $\nu_i^{L(R)}$

$$\nu_e^{L(R)} = \nu_1^{L(R)} \cos\theta + \nu_2^{L(R)} \sin\theta,$$
$$\nu_\mu^{L(R)} = -\nu_1^{L(R)} \sin\theta + \nu_2^{L(R)} \cos\theta$$

however, $\nu_i^{L(R)}$ are not stationary states in magnetic field $\mathbf{B} = (B_\perp, 0, B_\parallel)$

$$\nu_i^L(t) = c_i^+ \nu_i^+(t) + c_i^- \nu_i^-(t),$$
$$\nu_i^R(t) = d_i^+ \nu_i^+(t) + d_i^- \nu_i^-(t)$$

$\Leftarrow$ $\nu_i^{-(+)}$ stationary states in B

- Dirac equation $(\gamma_\mu p^\mu - m_i - \mu_i \boldsymbol{\Sigma} \boldsymbol{B}) \nu_i^s(p) = 0$ in a constant B

$$\hat{H}_i \nu_i^s = E \nu_i^s \qquad \hat{H}_i = \gamma_0 \boldsymbol{\gamma} \boldsymbol{p} + \mu_i \gamma_0 \boldsymbol{\Sigma} \boldsymbol{B} + m_i \gamma_0 \quad (s = \pm 1)$$

$\mu_{ij}(i \neq j) = 0$ •

ν spin operator that commutes with $\hat{H}_i$:

"bra-ket" products

- $$\hat{S}_i = \frac{1}{N} \left[\boldsymbol{\Sigma} \boldsymbol{B} - \frac{i}{m_i} \gamma_0 \gamma_5 [\boldsymbol{\Sigma} \times \boldsymbol{p}] \boldsymbol{B} \right]$$

$$\hat{S}_i \left| \nu_i^s \right\rangle = s \left| \nu_i^s \right\rangle, \; s = \pm 1$$

$$\langle \nu_i^s | \nu_k^{s'} \rangle = \delta_{ik} \delta_{ss'}$$ •

$$\frac{1}{N} = \frac{m_i}{\sqrt{m_i^2 \boldsymbol{B}^2 + \boldsymbol{p}^2 B_\perp^2}}$$

- ν energy spectrum

$$E_i^s = \sqrt{m_i^2 + p^2 + {\mu_i}^2 \mathbf{B}^2 + 2\mu_i s \sqrt{m_i^2 \mathbf{B}^2 + p^2 B_\perp^2}}$$

Probabilities of ν oscillations (flavour, spin and spin-flavour)

$\nu_e^L \leftrightarrow \nu_\mu^L$ $\quad P_{\nu_e^L \to \nu_\mu^L}(t) = \left|\langle \nu_\mu^L | \nu_e^L(t)\rangle\right|^2$

$\mu_\pm = \frac{1}{2}(\mu_1 \pm \mu_2)$ magnetic moments of ν mass states

flavour

$$P_{\nu_e^L \to \nu_\mu^L}(t) = \sin^2 2\theta \Big\{ \cos(\mu_1 B_\perp t)\cos(\mu_2 B_\perp t)\sin^2\frac{\Delta m^2}{4p}t + \sin^2(\mu_+ B_\perp t)\sin^2(\mu_- B_\perp t)\Big\}$$

spin

$$P_{\nu_e^L \to \nu_e^R} = \Big\{ \sin(\mu_+ B_\perp t)\cos(\mu_- B_\perp t) + \cos 2\theta \sin(\mu_- B_\perp t)\cos(\mu_+ B_\perp t)\Big\}^2 - \sin^2 2\theta \sin(\mu_1 B_\perp t)\sin(\mu_2 B_\perp t)\sin^2\frac{\Delta m^2}{4p}t.$$

spin-flavour

$$P_{\nu_e^L \to \nu_\mu^R}(t) = \sin^2 2\theta \Big\{ \sin^2 \mu_- B_\perp t \cos^2(\mu_+ B_\perp t) + \sin(\mu_1 B_\perp t)\sin(\mu_2 B_\perp t)\sin^2\frac{\Delta m^2}{4p}t\Big\}$$

... interplay of oscillations on vacuum $\omega_{vac} = \frac{\Delta m^2}{4p}$ and on magnetic $\omega_B = \mu B_\perp$ frequencies !

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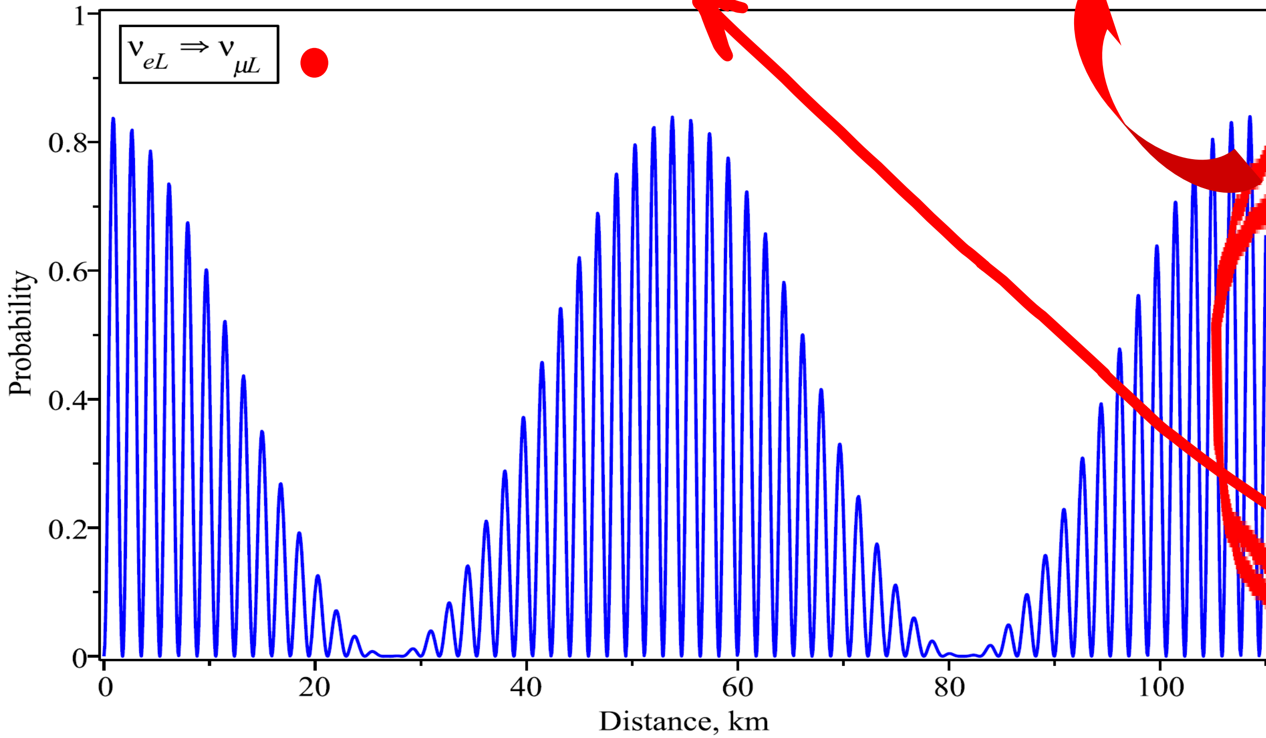
- For the case $\mu_1 = \mu_2$: probability of flavour oscillations

$$P_{\nu_e^L \to \nu_\mu^L} = \left(1 - \sin^2(\mu B_\perp t)\right) \sin^2 2\theta \sin^2 \frac{\Delta m^2}{4p} t = \left(1 - P^{cust}_{\nu_e^L \to \nu_e^R}\right) P^{cust}_{\nu_e^L \to \nu_\mu^L}$$

flavour

no spin oscillations



... amplitude of flavour oscillations on vacuum frequency $\omega_{vac} = \frac{\Delta m^2}{4p}$ is modulated by magnetic frequency $\omega_B = \mu B_\perp$

Fig. 1 The probability of the neutrino flavour oscillations $\nu_e^L \to \nu_\mu^L$ in the transversal magnetic field
- $B_\perp = 10^{16}$ G for the neutrino energy $p = 1$ MeV, $\Delta m^2 = 7 \times 10^{-5}$ eV^2 and magnetic moments $\mu_1 = \mu_2 = 10^{-20} \mu_B$.

Popov, Studenikin, Eur. Phys. J. C79 (2019) 144

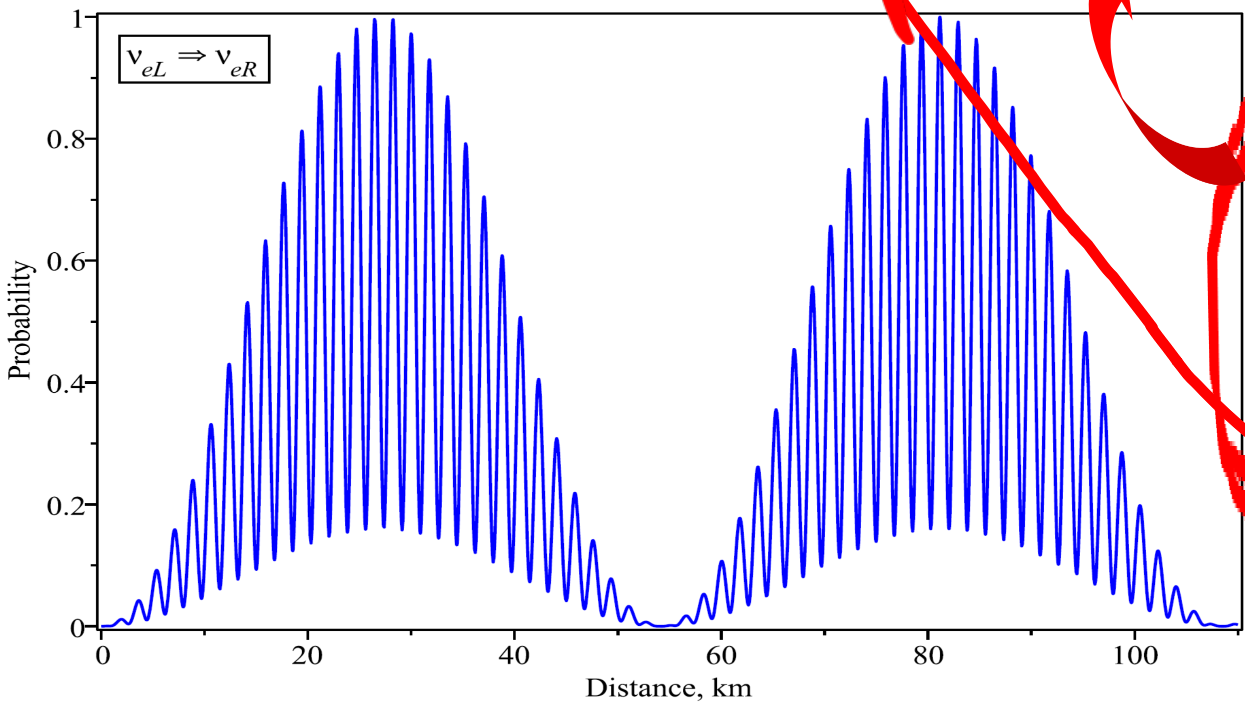
Chotorlishvili, Kouzakov, Kurashvili, Studenikin, Spin-flavor oscillations of ultrahigh-energy cosmic neutrinos n interstellar space: The role of neutrino magnetic moments, Phys. Rev. D96 (2017) 103017

For the case $\mu_1 = \mu_2$: probability of spin oscillations

no flavour oscillations

- $$P_{\nu_e^L \to \nu_e^R} = \left[1 - \sin^2 2\theta \sin^2\left(\frac{\Delta m^2}{4p}t\right)\right] \sin^2(\mu B_\perp t) = \left(1 - P^{cust}_{\nu_e^L \to \nu_\mu^L}\right) P^{cust}_{\nu_e^L \to \nu_e^R}$$

spin



… amplitude of spin oscillations on magnetic frequency $\omega_B = \mu B_\perp$ is modulated by vacuum frequency $\omega_{vac} = \frac{\Delta m^2}{4p}$

Fig. 2 The probability of the neutrino spin oscillations $\nu_e^L \to \nu_e^R$ in the transversal magnetic field $B_\perp = 10^{16}$ G for the neutrino energy $p = 1\ MeV$, $\Delta m^2 = 7 \times 10^{-5}\ eV^2$ and magnetic moments $\mu_1 = \mu_2 = 10^{-20}\mu_B$.

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- For the case $\mu_1 = \mu_2$: probability of spin-flavour oscillations

$$P_{\nu_e^L \to \nu_\mu^R} = \sin^2(\mu B_\perp t) \sin^2 2\theta \sin^2 \frac{\Delta m^2}{4p} t = P^{cust}_{\nu_e^L \to \nu_e^R} P^{cust}_{\nu_e^L \to \nu_\mu^L}$$

spin-flavour

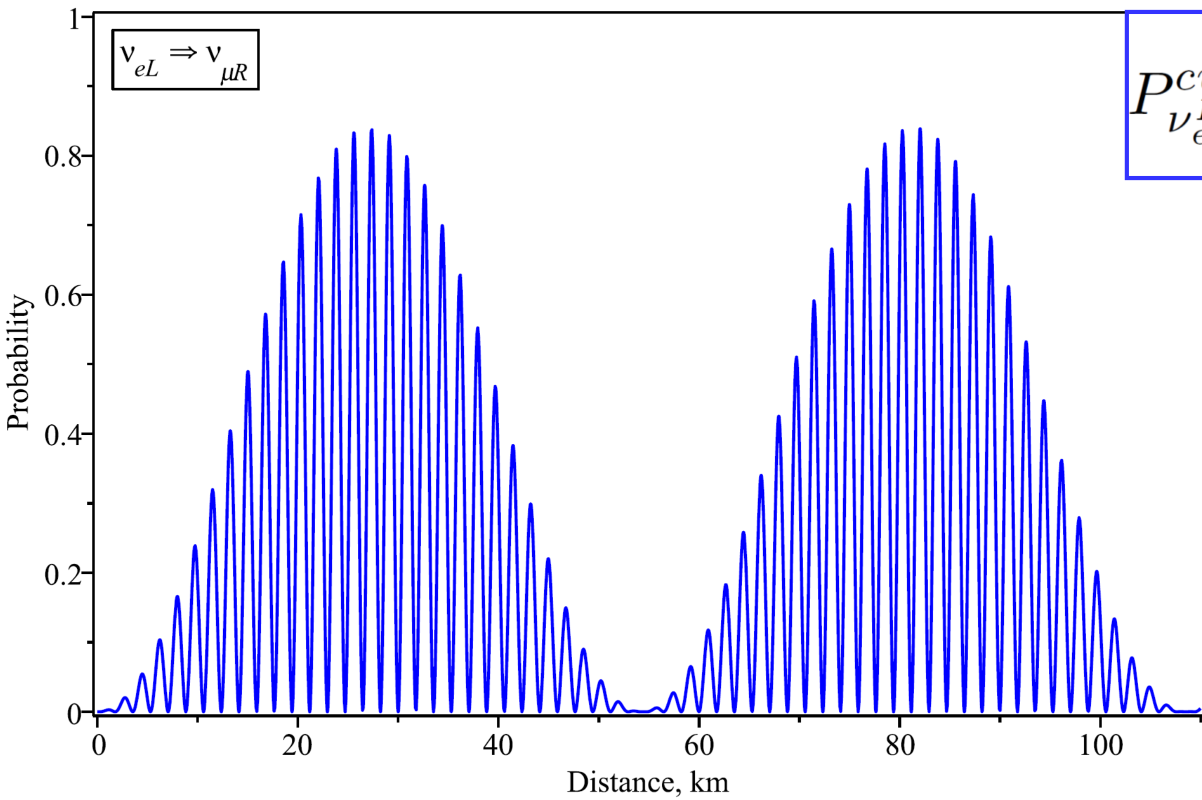

Fig. 3 The probability of the neutrino spin flavour oscillations $\nu_e^L \to \nu_\mu^R$ in the transversal magnetic field $B_\perp = 10^{16}\ G$ for the neutrino energy $p = 1\ MeV$, $\Delta m^2 = 7 \times 10^{-5}\ eV^2$ and magnetic moments $\mu_1 = \mu_2 = 10^{-20} \mu_B$.

$$P^{cust}_{\nu_e^L \to \nu_\mu^L} = \sin^2 2\theta \sin^2 \frac{\Delta m^2}{4p} t$$

$$P^{cust}_{\nu_e^L \to \nu_e^R} = \sin^2(\mu B_\perp t)$$

... interplay of oscillations on vacuum $\omega_{vac} = \frac{\Delta m^2}{4p}$ and on magnetic $\omega_B = \mu B_\perp$ frequencies

... in literature usually:

- $P_{\nu_e^L \nu_\mu^R} = \sin^2(\mu_{e\mu} B_\perp t) = 0$

$$\mu_{e\mu} = \frac{1}{2}(\mu_2 - \mu_1) \sin 2\theta$$

$$\mu_1 = \mu_2, \quad \mu_{ij} = 0, \ i \neq j$$

- ... M. Dvornikov, J. Maalampi, Phys. Lett. B 657 (2007) 217

- For completeness: ν survival $\nu_e^L \leftrightarrow \nu_e^L$ probability

... depends on μ_ν and B

$$P_{\nu_e^L\to\nu_e^L}(t) = \Big\{ \cos(\mu_+ B_\perp t)\cos(\mu_- B_\perp t) - \cos 2\theta \sin(\mu_+ B_\perp t)\sin(\mu_- B_\perp t) \Big\}^2 - \sin^2 2\theta \cos(\mu_1 B_\perp t)\cos(\mu_2 B_\perp t)\sin^2 \frac{\Delta m^2}{4p} t$$

$\sum$ of all probabilities (as it should be...) :

$$P_{\nu_e^L\to\nu_\mu^L} + P_{\nu_e^L\to\nu_e^R} + P_{\nu_e^L\to\nu_\mu^R} + P_{\nu_e^L\to\nu_e^L} = 1$$

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the discovered correspondence between flavour and spin oscillations in B can be important in studies of ν propagation in astrophysical environments

3 New effect in ν flavor oscillation in moving matter

$$\nu_e^L \Leftarrow (j_{||}, j_\perp) \Rightarrow \nu_\mu^L \qquad \boldsymbol{j}_\perp = n\boldsymbol{v}_\perp$$

longitudinal matter currents, transversal matter currents; invariant number density

Studenikin, to be published 2024 + arXiv: 1912.12491

- Equal role of $\boldsymbol{j}_\perp$ and $\boldsymbol{B}_\perp$ in generation of

$\nu_e^L \Leftarrow (j_\perp, B_\perp) \Rightarrow \nu_e^R$ spin oscillations

$\nu_e^L \Leftarrow (j_\perp, B_\perp) \Rightarrow \nu_\mu^R$ spin-flavour oscillations

- Probability of ν flavor oscillations $\nu_e^L \Leftarrow (j_{||}, j_\perp) \Rightarrow \nu_\mu^L$ in moving matter

$$P^{(j_{||}+j_\perp)}_{\nu_e^L\to\nu_\mu^L}(t) = \left(1 - P^{(j_\perp)}_{\nu_e^L\to\nu_e^R} - P^{(j_\perp)}_{\nu_e^L\to\nu_\mu^R}\right) P^{(j_{||})}_{\nu_e^L\to\nu_\mu^L}$$

$$P^{(j_{||})}_{\nu_e^L\to\nu_\mu^L}(t) = \sin^2 2\theta_{eff}\sin^2\omega_{eff}t, \quad \omega_{eff} = \frac{\Delta m^2_{eff}}{4p_0^\nu}$$

probability of flavor oscillations in $j_{||}$

probability of spin survival (not spin flip)

$$P^{j_\perp}_{\nu_e^L\to\nu_e^R}(t) = \frac{\left(\frac{\eta}{\gamma}\right)^2_{ee} v_\perp^2}{\left(\frac{\eta}{\gamma}\right)^2_{ee} v_\perp^2 + (1-\boldsymbol{v}\boldsymbol{\beta})^2}\sin^2\omega^{j_\perp}_{ee}t$$

spin oscillations in $\boldsymbol{j}_\perp$

$$P^{j_\perp}_{\nu_e^L\to\nu_\mu^R}(t) = \frac{\left(\frac{\eta}{\gamma}\right)^2_{e\mu} v_\perp^2}{\left(\frac{\eta}{\gamma}\right)^2_{e\mu} v_\perp^2 + \left(\frac{\Delta M}{\widetilde{G}n} - (1-\boldsymbol{v}\boldsymbol{\beta})\right)^2}\sin^2\omega^{j_\perp}_{e\mu}t$$

spin-flavor oscillations in $\boldsymbol{j}_\perp$

... is modulated by two "matter" frequencies ...

$$\omega^{j_\perp}_{ee} = \widetilde{G}n\sqrt{\left(\frac{\eta}{\gamma}\right)^2_{ee} v_\perp^2 + (1-\boldsymbol{v}\boldsymbol{\beta})^2}$$

$$\omega^{j_\perp}_{e\mu} = \widetilde{G}n\sqrt{\left(\frac{\eta}{\gamma}\right)^2_{e\mu} v_\perp^2 + \left(\frac{\Delta M}{\widetilde{G}n} - (1-\boldsymbol{v}\boldsymbol{\beta})\right)^2}$$

$$\left(\frac{\eta}{\gamma}\right)_{ee} = \frac{\cos^2\theta}{\gamma_{11}} + \frac{\sin^2\theta}{\gamma_{22}} \qquad {\gamma_{\alpha\alpha'}}^{-1} = \frac{1}{2}(\gamma_\alpha^{-1} + \gamma_{\alpha'}^{-1}) \qquad \gamma_\alpha^{-1} = \frac{m_\alpha}{E_\alpha}$$

$$\left(\frac{\eta}{\gamma}\right)_{e\mu} = \frac{\sin 2\theta}{\tilde{\gamma}_{21}} \qquad \tilde{\gamma}^{-1}_{\alpha\alpha'} = \frac{1}{2}(\gamma_\alpha^{-1} - \gamma_{\alpha'}^{-1})$$

PHYSICAL REVIEW D **103,** 115027 (2021)

Manifestations of nonzero Majorana *CP*-violating phases in oscillations of supernova neutrinos

Artem Popov[*]
Department of Theoretical Physics, Moscow State University, Moscow 119991, Russia

Alexander Studenikin[†]
Department of Theoretical Physics, Moscow State University, Moscow 119991, Russia and Joint Institute for Nuclear Research, Dubna 141980, Russia

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We investigate effects of nonzero Dirac and Majorana *CP*-violating phases on neutrino-antineutrino oscillations in a magnetic field of astrophysical environments. It is shown that in the presence of strong magnetic fields and dense matter, nonzero *CP* phases can induce new resonances in the oscillations channels $\nu_e \leftrightarrow \bar{\nu}_e$, $\nu_e \leftrightarrow \bar{\nu}_\mu$, and $\nu_e \leftrightarrow \bar{\nu}_\tau$. We also consider all other possible oscillation channels with ν_μ and ν_τ in the initial state. The resonances can potentially lead to significant phenomena in neutrino oscillations accessible for observation in experiments. In particular, we show that neutrino-antineutrino oscillations combined with Majorana-type *CP* violation can affect the $\bar{\nu}_e/\nu_e$ ratio for neutrinos coming from the supernovae explosion. This effect is more prominent for the normal neutrino mass ordering. The detection of supernovae neutrino fluxes in the future experiments, such as JUNO, DUNE, and Hyper-Kamiokande, can give an insight into the nature of *CP* violation and, consequently, provides a tool for distinguishing the Dirac or Majorana nature of neutrinos.

DOI: 10.1103/PhysRevD.103.115027

I. INTRODUCTION

CP symmetry implies that the equations of motion of a system remain invariant under the *CP* transformation, that is a combination of charge conjugation (*C*) and parity inversion (*P*). In 1964, with the discovery of the neutral kaon decay [1], it was confirmed that *CP* is not an underlying symmetry of the electroweak interactions theory, thus opening a vast field of research in *CP* violation. Currently, *CP* violation is a topic of intense studies in particle physics that also has important implications in cosmology. In 1967, Sakharov proved that the existence of *CP* violation is a necessary condition for generation of the baryon asymmetry through baryogenesis in the early Universe [2]. A review of possible baryogenesis scenarios can be found in [3].

Today we have solid understanding of *CP* violation in the quark sector, that appears due to the complex phase in the Cabibbo-Kobayashi-Maskawa matrix parametrization. Its magnitude is expressed by the Jarlskog invariant $\mathcal{J}_{\rm CKM} = (3.18 \pm 0.15) \times 10^{-5}$ [4], which seems to be excessively small to engender baryogenesis at the electroweak phase transition scale [3]. However, in addition to experimentally confirmed *CP* violation in the quark sector, *CP* violation in the lepton (neutrino) sector hypothetically exists (see [5] for a review). Leptonic *CP* violation is extremely difficult to observe due to weakness of neutrino interactions. In 2019, a first breakthrough happened when NOνA [6] and T2K [7] collaborations reported constraints on the Dirac *CP*-violating phase in neutrino oscillations. Hopefully, future gigantic neutrino experiments, such as DUNE [8] and Hyper-Kamiokande [9], also JUNO [10] with detection of the atmospheric neutrinos, will have a good chance significantly improve this results. Note that leptonic *CP* violation plays an important role in baryogenesis through leptogenesis scenarios [11].

The *CP*-violation pattern in the neutrino sector depends on whether neutrino is a Dirac or Majorana particle. The Pontecorvo-Maki-Nakagawa-Sakata (PMNS) neutrino mixing matrix in the most common parametrization has the following form

[*]ar.popov@physics.msu.ru
[†]studenik@srd.sinp.msu.ru

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Phys. Rev. D103 (2021) 115027

… the role of Majorana CP-violating phases in neutrino oscillations

$$\nu_e \leftrightarrow \bar{\nu}_{e,\mu,\tau}$$

in strong B and dense matter of supernovae for two mass hierarchies

… Majorana CP phases induce new resonances

… a tool for distinguishing Dirac-Majorana nature of ν

… see also presentation by Artem Popov, May 20, 2024

PHYSICAL REVIEW D **101,** 056004 (2020)

Neutrino quantum decoherence engendered by neutrino radiative decay

Konstantin Stankevich*
Department of Theoretical Physics, Faculty of Physics, Lomonosov Moscow State University, 119992 Moscow, Russia

Alexander Studenikin†
Department of Theoretical Physics, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia
and Dzhelepov Laboratory of Nuclear Problems, Joint Institute for Nuclear Research, Dubna 141980, Moscow Region, Russia

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A new theoretical framework, based on the quantum field theory of open systems applied to neutrinos, has been developed to describe the neutrino evolution in external environments accounting for the effect of the neutrino quantum decoherence. The developed new approach enables one to obtain the explicit expressions of the decoherence and relaxation parameters that account for a particular process, in which the neutrino participates, and also for the characteristics of an external environment and of the neutrino itself, including the neutrino energy. We have used this approach to consider a new mechanism of the neutrino quantum decoherence engendered by the neutrino radiative decay to photons and dark photons in an astrophysical environment. The importance of the performed studies is highlighted by the prospects of the forthcoming new large volume neutrino detectors that will provide new frontier in high-statistics measurements of neutrino fluxes from supernovae.

DOI: 10.1103/PhysRevD.101.056004

I. INTRODUCTION

Half a century ago Gribov and Pontecorvo derived [1] the first analytical expression for the neutrino oscillation probability that has opened a new era in the theoretical and experimental studies of the neutrino oscillation phenomenon. The neutrino oscillation patterns can be modified by neutrino interactions with external environments including electromagnetic fields that can influence neutrinos in the case neutrinos have nonzero electromagnetic properties [2]. The phenomenon of neutrino oscillations can proceed only in the case of the coherent superposition of neutrino mass states. An external environment can modify a neutrino evolution in a way that conditions for the coherent superposition of neutrino mass states are violated. Such a violation is called quantum decoherence of neutrino states and leads to the suppression of flavor neutrino oscillations. It should be noted that the quantum neutrino decoherence differs from the standard neutrino decoherence that appears due to separation of neutrino wave packets, the effect that is not considered below.

The quantum neutrino decoherence has attracted a growing interest during the last 15 years. Within reasonable amount of the performed studies, the method based on the Lindblad master equation [3,4] for describing neutrino evolution has been used. This approach is usually considered as the most general one that gives a possibility to study neutrino quantum decoherence as a consequence of standard and nonstandard interactions of a neutrino system with an external environment [5–15].

The Lindblad master equation can be written in the following form (see, for instance, [13]):

$$\frac{\partial \rho_\nu(t)}{\partial t} = -i[H_S, \rho_\nu(t)] + D[\rho_\nu], \qquad (1)$$

where ρ_ν is the density matrix that describes the neutrino evolution, H_S is the Hamiltonian, and the dissipation term (or dissipator) is given by

$$D[\rho_\nu(t)] = \frac{1}{2}\sum_{k=1}^{N^2-1}[V_k, \rho_\nu V_k^\dagger] + [V_k\rho_\nu, V_k^\dagger], \qquad (2)$$

where V_k are dissipative operators that arise from interaction between the neutrino system and the external

*kl.stankevich@physics.msu.ru
†studenik@srd.sinp.msu.ru

Stankevich, Studenikin,
Neutrino quantum decoherence engendered by neutrino radiative decay
Phys.Rev.D 101 (2020) 056004

… ν radiative decay as a source of quantum decoherence in extreme astrophycal environments

⇓

… observable consequences for SN ν
(JUNO, DUNE, Hyper-Kamiokande)

① Electromagnetic Properties of ν

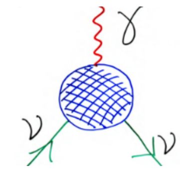
MSU Alexander Studenikin NCPM

C.Giunti, A.Studenikin, " ν electromagnetic interactions: A window to new physics", Rev.Mod.Phys, 2015

A.Studenikin, "Overview of ν electromagnetic properties 2022, arXiv:2301.06071

National Center FOR PHYSICS AND MATHEMATICS

① νEP theory - ν vertex function



matrices in ν mass eigenstates space

$$\Lambda^{if}_\mu(q) = f^{if}_Q(q^2)\gamma_\mu + f^{if}_M(q^2) i\sigma_{\mu\nu}q^\nu + f^{if}_E(q^2)\sigma_{\mu\nu}q^\nu\gamma_5 + f^{if}_A(q^2)(q^2\gamma_\mu - q_\mu \not{q})\gamma_5,$$

electric charge, magnetic moment, electric moment, anapole moment

form factors $f^{if}_X(q^2)$ at $q^2 = 0$ — static EP of ν ⇒

Dirac	ν	Majorana	
q_{if}, μ_{if}, ε_{if}, a_{if}	$\neq 0$	$q_{if} = 0$, μ_{if} $(i \neq f)$, ε_{if} $(i \neq f)$, a_{if}	CPT + charge concervation

Hermiticity and discrete symmetries of EM current $\langle \nu(p')|J^{EM}_\mu|\nu(p)\rangle = \bar{u}(p')\Lambda_\mu(q)u(p)$ put constraints on form factors

② $$\mu^D_{jj} = \frac{3e_0 G_F m_j}{8\sqrt{2}\pi^2} \approx 3.2 \times 10^{-19}\mu_B \left(\frac{m_j}{1\text{ eV}}\right)$$

Fujikawa & Shrock, 1980

- much greater values are Beyond Minimally Extended SM
- transition moments $\mu_{i\neq f}$, $\varepsilon_{i\neq f}$ are GIM suppressed

③ ν EMP experimental bounds

$\mu^{eff}_\nu < 2.9$ / 2.8 / ~ 0.1 $\times 10^{-11} \mu_B$

GEMMA 2012

Borexino 2017 ~ XENON1T 2020

astrophys., Raffelt ea 1988, 2020

Arcoa Dias ea 2015

$q_{\nu_e} <\sim 10^{-12}$, $\sim 10^{-19}$, $\sim 10^{-21}$ e_0

reactor ν scattering — AS '14, Chen ea '14

AS '14 (astrophysics)

neutrality of matter

• charge rad. $\langle r^2_\nu \rangle$ is most accessible for exp. observations •

2 ν electromagnetic properties: Future prospects

- new constraintson μ_ν (and q_ν) from GEMMA-2/ νGeN and Borexino (?)
- XENON1T an excess in electronic recoil events in < 7 keV (2-3 keV) over known backgrnds ⟹ $\mu_\nu \in (1.4,\ 2.9) \times 10^{-11}\,\mu_B$

E. Aprile et al, XENON1T coll., Phys Rev. D 102 (2020) 072004

! XENONnT

O.Miranda, D. Papoulias, M. Tórtola, J. W. F. Valle, XENON1T signal from transition neutrino magnetic moments, Phys.Lett. B 808 (2020) 135685

- new improved limit from stellar evolution data for global cluster ω-Centauri ⟹ $\mu_\nu < 2.2 \times 10^{-12}\mu_B$

S. Arceo-Diaz, K.-P.Schroder, K.Zuber, D.Jack, Astropart.Phys. 70 (2015) 1

- new improved limit $\mu_\nu < 1.2 \times 10^{-12}\mu_B$

F.Capozzi, G.Raffelt, Phys. Rev. D 102 (2020) 083007

comes from improved new calibrations of tip of red-giant branch which allows one to constrain novel energy losses

- new setup to observe coherent elastic neutrino-atom scattering using electron antineutrinos from tritium decay and a liquid helium target ⟹ upper limit

M.Cadeddu, F.Dordei, C.Giunti, K.Kouzakov, E.Picciau, A.Studenikin, Potentialities of a low-energy detector based on 4 He evaporation to observe atomic effects in coherent neutrino scattering and physics perspectives, Phys. Rev. D100 (2019) no.7, 073014

$\mu_\nu < 7 \times 10^{-13}\mu_B$

⟹ XENONnT Collaboration ... Most recent results

- Upgraded experiment has managed to rule out the so-called XENON1T excess by using a new larger liquid xenon (LXe)

Implications of first LUX-ZEPLIN and XENONnT results

A.Khan, Light new physics and neutrino electromagnetic interactions in XENONnT, Phys.Lett.B 837 (2023) 137650

M.Atzori Corona, W.Bonivento, M.Cadeddu, N.Cargioli, F.Dordei, New constraint on neutrino magnetic moment from LZ dark matter search results, Phys.Rev.D107 (2023) 053001

K.ShivaSankar, A.Majumdar, D.Papoulias, H.Prajapati, R. Srivastava, Phys.Lett. B 839 (2023) 137742

C.Giunti, C.Ternes, Testing neutrino electromagnetic properties at current and future dark matter experiments, Phys. Rev. D 108 (2023) 095044

... new stringent upper limits on effective and transition $\mu_\nu \sim few \times 10^{-12} \mu_B$

... new stringent upper limits on $q_\nu \sim 10^{-13}\ e_0$

2023 + soon ...

νGeN experiment ... low threshold ... $T \sim 200\ \ eV$

$\mu_\nu \sim (5-9) \times 10^{-12} \mu_B$

I.Alekseev et al., First results of the νGeN experiment on coherent elastic neutrino-nucleus scattering, Phys.Rev.D 106 (2022) 5, L051101

⟹ A.Studenikin, Europhys. Lett. 107 (2014) 21001 $|q_\nu| < 1.1 \times 10^{-13} e_0$

The search for coherent elastic neutrino-atom scattering and neutrino magnetic moment in Sarov

Sarov Tritium Neutrino Experiment = SATURNE Collaboration)

Matteo Cadeddu (INFN, Cagliari), Francesca Derdei (INFN, Cagliari), Carlo Giunti (INFN, Turin),
Konstantin Kouzakov (MSU), Bayarto Lubsandorzhiev (INR RAS), Oleg Moskalev, (VNIIEF, Sarov),
Ivan Stepantsov (MSU), Alexander Studenikin (MSU), Vladimir Trofimov (JINR),
Maxim Vyalkov (MSU), Arkady Yukhimchuk (VNIIEF, Sarov)

Potentialities of a low-energy detector based on ^{4}He evaporation to observe atomic effects in coherent neutrino scattering and physics perspectives

M. Cadeddu,[1,*] F. Dordei,[2,†] C. Giunti,[3,‡] K. A. Kouzakov,[4,§] E. Picciau,[1,∥] and A. I. Studenikin[5,6,¶]

[1]Università degli studi di Cagliari and Istituto Nazionale di Fisica Nucleare (INFN), Sezione di Cagliari, Complesso Universitario di Monserrato—S.P. per Sestu Km 0.700, 09042 Monserrato (Cagliari), Italy

[2]Istituto Nazionale di Fisica Nucleare (INFN), Sezione di Cagliari, Complesso Universitario di Monserrato—S.P. per Sestu Km 0.700, 09042 Monserrato (Cagliari), Italy

[3]Istituto Nazionale di Fisica Nucleare (INFN), Sezione di Torino, Via P. Giuria 1, I-10125 Torino, Italy

[4]Department of Nuclear Physics and Quantum Theory of Collisions, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia

[5]Department of Theoretical Physics, Faculty of Physics, Lomonosov Moscow State University, Moscow 119991, Russia

[6]Joint Institute for Nuclear Research, Dubna 141980, Moscow Region, Russia

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We propose an experimental setup to observe coherent elastic neutrino-atom scattering (CEνAS) using electron antineutrinos from tritium decay and a liquid helium target. In this scattering process with the whole atom, that has not been observed so far, the electrons tend to screen the weak charge of the nucleus as seen by the electron antineutrino probe. The interference between the nucleus and the electron cloud produces a sharp dip in the recoil spectrum at atomic recoil energies of about 9 meV, reducing sizably the number of expected events with respect to the coherent elastic neutrino-nucleus scattering case. We estimate that with a 60 g tritium source surrounded by 500 kg of liquid helium in a cylindrical tank, one could observe the existence of CEνAS processes at 3σ in 5 yr of data taking. Keeping the same amount of helium and the same data-taking period, we test the sensitivity to the Weinberg angle and a possible neutrino magnetic moment for three different scenarios: 60, 160, and 500 g of tritium. In the latter scenario, the Standard Model (SM) value of the Weinberg angle can be measured with a statistical uncertainty of $\sin^2\vartheta_W^{\rm SM}{}^{-0.015}_{-0.016}$. This would represent the lowest-energy measurement of $\sin^2\vartheta_W$, with the advantage of being not affected by the uncertainties on the neutron form factor of the nucleus as the current lowest-energy determination. Finally, we study the sensitivity of this apparatus to a possible electron neutrino magnetic moment and we find that using 60 g of tritium it is possible to set an upper limit of about $7\times10^{-13}\mu_B$ at 90% C.L., that is more than one order of magnitude smaller than the current experimental limit.

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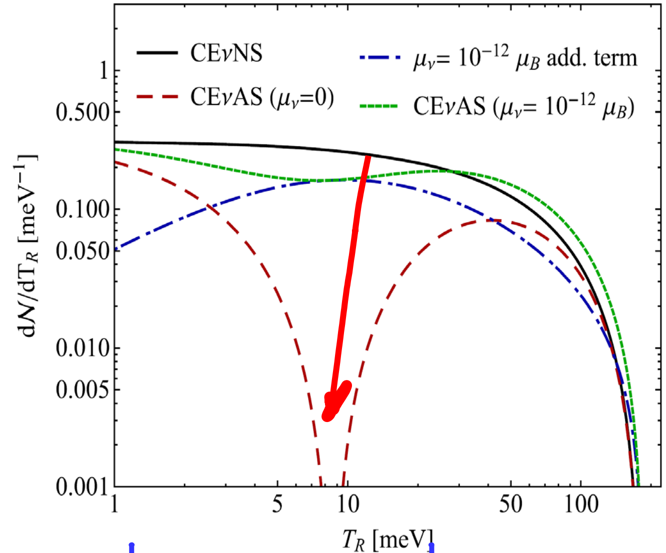
I. INTRODUCTION

Coherent elastic neutrino-nucleus scattering (CEνNS) has been recently observed by the COHERENT experiment [1,2], after many decades from its prediction [3–5].

This observation triggered a lot of attention from the scientific community and unlocked a new and powerful tool to study many and diverse physical phenomena: nuclear physics [6,7], neutrino properties [8–10], physics beyond the Standard Model (SM) [11–17], and electroweak interactions [18,19]. The experimental challenge related to the CEνNS observation is due to the fact that in order to meet the coherence requirement $qR \ll 1$ [20], where $q = |\vec{q}|$ is the three-momentum transfer and R is the nuclear radius, one has to detect very small nuclear recoil energies E_R, lower than a few keV.

At even lower momentum transfers, such that $qR_{\rm atom} \ll 1$, where $R_{\rm atom}$ is the radius of the target atom including the electron shells, the reaction can be viewed as taking place on the atom as a whole [21]. This effect should be visible for $qR_{\rm atom} \sim 1$, i.e., for momentum

*matteo.cadeddu@ca.infn.it
†francesca.dordei@cern.ch
‡carlo.giunti@to.infn.it
§kouzakov@gmail.com
∥emmanuele.picciau@ca.infn.it
¶studenik@srd.sinp.msu.ru



In our paper we have proposed an experimental setup to observe coherent elastic neutrino-atom scattering (CEνAS) using electron antineutrinos from tritium decay and a supefluid ^{4}He target.

In this scattering process with the whole atom, that has not been observed so far, the electrons tend to screen the weak charge of the nucleus as seen by the electron antineutrino probe.

$\mu_\nu \sim 10^{-13}\mu_B$

supefluid ^{4}He target technology (HeRALD) for direct detection of sub-GeV DM has been recently proposed in: S.Hartel et al., Phys.Rev.D 100 (2019) 9, 092007

LOMONOSOV MOSCOW STATE UNIVERSITY

National Center FOR PHYSICS AND MATHEMATICS

The Sarov experiment for probing coherent elastic neutrino-atom scattering and neutrino electromagnetic interactions

Konstantin Kouzakov

on behalf of

The SATURNE Collaboration

… see presentation of May 21, 2024

Thank you